

PROBING FOR EXOPLANETS HIDING IN DUSTY DEBRIS DISKS: DISK IMAGING, CHARACTERIZATION, AND EXPLORATION WITH *HST*/STIS MULTI-ROLL CORONAGRAPHY

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ABSTRACT

Spatially resolved scattered-light images of circumstellar debris in exoplanetary systems constrain the physical properties and orbits of the dust particles in these systems. They also inform on co-orbiting (but unseen) planets, the systemic architectures, and forces perturbing the starlight-scattering circumstellar material. Using *Hubble Space Telescope* (*HST*)/Space Telescope Imaging Spectrograph (STIS) broadband optical coronagraphy, we have completed the observational phase of a program to study the spatial distribution of dust in a sample of 10 circumstellar debris systems and 1 “mature” protoplanetary disk, all with *HST* pedigree, using point-spread-function-subtracted multi-roll coronagraphy. These observations probe stellocentric distances ≥ 5 AU for the nearest systems, and simultaneously resolve disk substructures well beyond corresponding to the giant planet and Kuiper Belt regions within our own solar system. They also disclose diffuse very low-surface-brightness dust at larger stellocentric distances. Herein we present new results inclusive of fainter disks such as HD 92945 ($F_{\text{disk}}/F_{\text{star}} = 5 \times 10^{-5}$), confirming, and better revealing, the existence of a narrow inner debris ring within a larger diffuse dust disk. Other disks with ring-like substructures and significant asymmetries and complex morphologies include HD 181327, for which we posit a spray of ejecta from a recent massive collision in an exo-Kuiper Belt; HD 61005, suggested to be interacting with the local interstellar medium; and HD 15115 and HD 32297, also discussed in the context of putative environmental interactions. These disks and HD 15745 suggest that debris system evolution cannot be treated in isolation. For AU Mic’s edge-on disk, we find out-of-plane surface brightness asymmetries at ≥ 5 AU that may implicate the existence of one or more planetary perturbers. Time-resolved images of the MP Mus protoplanetary disk provide spatially resolved temporal variability in the disk illumination. These and other new images from our *HST*/STIS GO/12228 program enable direct inter-comparison of the architectures of these exoplanetary debris systems in the context of our own solar system.

Key words: methods: observational – planet–disk interactions – stars: individual (HD 15115, HD 15745, HD 32297, HD 53143, HD 61005, HD 92945, HD 107146, HD 139664, HD 181327, Au Mic, MP Mus)

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1. INTRODUCTION

Understanding how planetary systems, including our own, form and evolve frame one of the most fundamentally compelling areas of inquiry in contemporary astrophysics (Beuther et al. 2014). Over the last few decades, optical, infrared (IR), and millimeter observations have shown that most stars are surrounded at birth by circumstellar (CS) accretion disks of primordial material. Separately, more than 1000 confirmed, and thousands more candidate, extrasolar planets detected to date, mostly by photometric transit and radial velocity (RV) methods, are believed to originate in such disks with Jovian-mass planets forming before CS accretion disks become gas-poor over time.

The evolutionary transformation of gas-rich disks to systems with small amounts of dust and little residual gas occurs over the first ~ 10 Myr of the host star’s lifetime, with warm dust being rare in systems after ~ 3 Myr (Haisch et al. 2001), gas being rare after 5–10 Myr, and stellar mass accretion declining to $\sim 10^{-11} M_{\odot} \text{ yr}^{-1}$ over the same epoch (Fedele et al. 2010). Much of the dust may then be tied up in smaller rocky or icy bodies that can interact with co-orbiting giant planets or other small bodies and collide, thus replenishing the CS environments (transformed into “debris disks”) with second-generation material. Such collisional replenishment of dust-rich CS particles can continue throughout the lifetime of the system (Carpenter et al. 2005; Roccatagliata et al. 2009; Panić et al. 2013). In our own solar system, even after 4.6 Gyr, this is evidenced by the presence of our own zodiacal dust cloud, asteroid and Kuiper

¹³ Deceased.

Belts, and population of comets. Indeed, this dust-production paradigm was reinforced with *Hubble Space Telescope* (*HST*) and *Rosetta* imaging of the aftermath of the apparent dynamical collision or rotational disruption of the minor planet P/2010 A2, releasing a wake of trailing debris (Jewitt et al. 2010), even as competing forces (e.g., Poynting–Robertson drag and radiation pressure “blow out”) attempt to clear the circum-solar environment of such particles. However, because these dynamical processes occur on much shorter timescales than a host star’s lifetime, a quasi-equilibrium condition may exist where dusty debris is continuously generated through collisions and the net solid mass declines with age (Wyatt et al. 2007a, 2007b; Wyatt 2008). The outcomes of such interactions will analogously affect the compositional content and spatial distribution of dusty debris in exoplanetary systems on spatial scales large compared to planets.

Dynamical interactions of planets with residual planetesimals likely play a vital role in shaping the architectures of planetary systems. Recently formed terrestrial planets can generate copious quantities of dust in the early phases of planetary system formation (Kenyon & Bromley 2004). For example, in the history of our solar system, Jupiter is thought to have wrought considerable destruction in the early evolution of the asteroid belt (Bottke et al. 2005; Nesvorný et al. 2013), scattering volatile-rich planetesimals toward the inner solar system that may have contributed significant water to, and affected the habitability of, the early Earth (Raymond et al. 2004, 2014; Morbidelli et al. 2012; Albertsson et al. 2014). The Nice model (Tsiganis et al. 2005; Morbidelli et al. 2005) for the dynamical re-arrangement of the solar system explains (1) the migration of the outer planets; (2) the depletion of dust-producing planetesimals by factors of 10 in both the Asteroid and Kuiper Belts; and (3) evidence for a late-heavy bombardment in the lunar cratering record (Strom et al. 2005). High spatial resolution (\lesssim AU scale) and high fidelity maps of dusty CS debris in exoplanetary systems, in which the presence (or absence) of co-orbiting planetary-mass bodies may be imprinted or encoded on the imaged dust distribution, are needed to test the general applicability of such models around other stars.

CS dust disks (exo-Kuiper Belts) were first identified by thermal emission in excess of that expected from the stellar photosphere. This is how the majority of debris disks at $d < 100$ pc are known (Backman & Paresce 1993) and continue to be discovered (Eiroa et al. 2013). Most modeling of these disks has focused on the wavelength dependence of the IR excess emission (i.e., their spectral energy distributions; SEDs), but is not unique. Different spatial distributions of dust and dust particle properties can produce similar SEDs. Breaking these degeneracies requires spatially resolved imagery. With sufficient spatial resolution, we can potentially observe the effects of sculpting of the debris by dynamical interactions with planets and other phenomena. At long wavelengths, such imaging studies are currently limited to far-IR thermal emission from nearby ($d < 50$ pc) disks obtained with the 3.5 m *Herschel Space Observatory*, and the few debris disks studied to date with the Atacama Large Millimeter Array (ALMA) sensitive to larger particles. Those disks that are bright enough to be imaged in scattered near-IR or optical light are observable with ground-based 8–10 m class adaptive-optics (AO)-augmented telescopes with coronagraphs, or with *HST* coronagraphy. In particular, these CS debris systems are detectable and resolvable using the Space Telescope Imaging Spectrograph (STIS), with point-spread-function (PSF)-template-subtracted coronagraphy

(PSFTSC), and can uniquely provide detailed information about the systems and their architectures even when the individual planets remain elusive.

Initially, only $\sim 15\%$ of the thermally emissive debris disks observed by *HST* were detected in scattered light, although additional disks continue to be discovered from archival data as post-processing techniques continue to improve, and as resolved detections from other wavelengths continue to accumulate. The discovery images of these light-scattering disks exhibit a diversity of radially “confined” and non-axisymmetric features in the stellocentric regions thus far explored, including photo-centric offsets (Schneider et al. 2009; Kalas et al. 2005, 2008), warps (Heap et al. 2000), clumps (Holland et al. 1998; Greaves et al. 1998; Greaves 2012), non-coplanar features (Golimowski et al. 2006), spirality (Clampin et al. 2003), and brightness asymmetries (Telesco et al. 2000; Kalas et al. 2007a, 2007b). These are exactly the morphological features predicted to arise in starlight-scattering CS debris perturbed by co-orbiting planets (Wyatt et al. 1999; Augereau et al. 2001; Kuchner & Holman 2003; Wyatt 2003, 2005; Wolf et al. 2007; Ertel et al. 2012). Some of these features, however, may also result from different evolutionary scenarios positing the production of belts of colliding planetesimals. The challenge then is to distinguish between disks arising from different mechanisms, e.g., between disks with planetesimals captured into mean motion resonances with planets, and disks where planetesimals are self-stirred (Wyatt 2008).

Distinguishing between these scenarios can best be achieved with images that reveal and spatially resolve most of the debris system substructures. Most of the “discovery” images, however, have been plagued with optical artifacts due to systematic effects from incomplete starlight suppression (i.e., residuals from limited-efficacy coronagraphy and PSF subtraction), and are of insufficient (or questionable) fidelity. To overcome this problem, we have obtained new, much higher fidelity and sensitivity, visible-light observations of an *HST*-selected sample of CS disks using a multi-roll PSFTSC imaging technique with STIS. Our new STIS observations achieve inner working angles (IWAs) of $r \geq 0.3$, equal to that previously provided by *HST*’s Near Infrared Camera and Multi-Object Spectrometer (NICMOS) coronagraph, but with a field of view (FOV) and visible-light spatial resolution (≈ 60 mas) comparable to that of *HST*’s Advanced Camera for Surveys (ACS) with its much larger effective IWAs. (Today, neither the NICMOS nor ACS coronagraphs are operational.) These new STIS images probe the interior regions of these debris systems, with inner working distances below ≈ 8 AU for about half the stars in this sample, corresponding to the giant planet and Kuiper Belt regions within our own solar system. This enables direct inter-comparison of the architectures of these exoplanetary debris systems in the context of our own solar system, and also studies of disk dynamics, grain properties, and systemic evolution. These in turn can be used to deduce (or exclude) the existence of planets, and, more importantly, can set constraints on planetary masses, orbital distances, eccentricities, and evolutionary history (Chiang et al. 2009; Rodigas et al. 2014b).

Our target sample is discussed in Section 2 of this paper. In Section 3, we discuss the observational motivation and goals of the program predicated on the history of “lessons learned” from nearly two decades of *HST* coronagraphy and PSF subtraction. In Section 4, we detail our observational approach to greatly improve upon a priori anticipated sensitivity limiting instrumental and extrinsic systematic effects that arise from

Table 1
Survey Targets*

Target	V mag	$B - V$	Spec	Dist. (pc)	Age (Myr)	Disk $L_{\text{IR}}/L_{\text{star}}$ (%)	Initial <i>HST</i> Imaging	
							Instrument	References
HD 15115	7.150	+0.35	F2	45.2	12? ^a	0.05	ACS	Kalas et al. (2007b)
HD 15745	7.824	+0.32	F2V	63.5	~ 100 ? ^b	0.12	ACS	Kalas et al. (2007a)
HD 32297	8.304	+0.20	A5V ⁺	112.4	~ 10 ^c	0.27	NIC	Schneider et al. (2005)
HD 53143	6.820	+0.80	G9V	18.3	1000 ± 300 ^d	0.025	ACS	Kalas et al. (2006)
HD 61005	8.932	+0.71	G8V	34.4	90 ± 40 ^e	0.25	NIC	Hines et al. (2007)
HD 92945	8.592	+0.89	K1V	21.4	$80\text{--}150/300$ ^f	0.076	ACS	Golimowski et al. (2011)
HD 107146	7.622	+0.62	G2V	27.5	$80\text{--}200$ ^g	0.12	ACS	Ardila et al. (2004, 2005)
HD 139664	4.64	+0.40	F5V	17.4	$300(-200, +700)$ ^h	0.009	ACS	Kalas et al. (2006)
HD 181327	7.475	+0.48	F6V	51.8	$\sim 12\text{--}20$ ⁱ	0.25	NIC/ACS	Schneider et al. (2006)
AU Mic	10.277	+1.44	M1V	9.9	$6\text{--}20$ ^j	0.44	ACS [†]	Krist et al. (2005)
MP Mus	10.44	+0.94	K1Ve	~ 86	13 ± 5		NIC	Cortes et al. (2009)

Notes.

* For additional characteristics of the CS debris systems resulting from this study, see Table 5.

Age estimation: ^a Moór et al. 2006; ^b Kalas et al. 2007a; ^c Schneider et al. 2005; ^d Decin et al. 2000; Song et al. 2000; Nordstrom et al. 2004; ^e Hines et al. 2007; ^f López-Santiago et al. 2006; ^g Williams et al. 2004; ^h Nordström et al. 2004; ⁱ Schneider et al. 2006; ^j Zuckerman et al. 2001.[†] AU Mic ground-based discovery imaging: Kalas et al. (2004).

+ Updated spectral type since Schneider et al. (2005) by Fitzgerald et al. (2007).

incomplete or imperfect PSF subtractions (if observationally unconstrained or uncontrolled). Therein we elaborate upon the observations that define and enable our multi-roll PSFTSC by utilizing time-constrained multiple field (spacecraft roll) orientation (“ N ”-roll, with $N = 6$ by optimal design). In Section 5, we present our image data reduction, and PSF-subtraction methods applied across the target set. In Section 6, we place the STIS multi-roll PSFTFC observations in the context of CS disk observations with other instruments, facilities, and at other wavelengths obtained in space and on the ground. In Section 7, we present, at a high level across the target set, the observational results of the imaging program, with commentary in Section 8. Finally, in the Appendix, we present and discuss the imaging results and derived properties and characteristics of each individual target system in detail. Detailed analysis and modeling of the image data for individual targets will be presented in follow-up papers, (e.g., Stark et al. 2014 for HD 181327, D. C. Hines et al. 2014 (in preparation) for HD 61005, J. H. Debes et al. 2014 (in preparation) for HD 32297, and others).

2. THE *HST*/GO 12228 TARGET SAMPLE

We observed an *HST*-selected ensemble of 10 CS debris disks and 1 “mature” protoplanet-ary/transition disk of an age comparable to the youngest debris disks in our sample (see Table 1). All targets were previously observed with *HST* coronagraphy in heterogeneously diverse discovery or survey programs. We re-visited disk targets whose surface brightnesses (SBs) and geometries were sufficiently well determined (in the regions observed) from those previous ACS optical or NICMOS short wavelength near-IR images to ensure robust imaging with our multi-roll PSFTSC strategy.

Scattered-light images prior to this study, obtained with NICMOS, and/or ACS, were limited in the stellocentric regions probed as follows.

1. Only the outer regions (typically $r \gtrsim 2''$) were probed due to the limiting effective IWA of the ACS coronagraph; or
2. Spatial resolution was limited (two to four times worse than optical coronagraphy) and PSF-subtraction artifacts remain for those disks probed as closely as $r = 0.3$ with NICMOS,

due to the longer (than optical) wavelength employed in the near-IR and two-roll only observations secured; or

3. Sensitivity to low-SB dust (though detected either at low signal-to-noise ratio (S/N) or only where the dust was relatively brightest) was limited due to the relatively shallow exposure depths used in the imaging survey programs for many of the disk discovery images; or
4. Compositionally diagnostic “color” information (wavelength-dependent spectral reflectivity of the grains), both in bulk and spatially resolved within the disks, is currently lacking because observations were conducted in only the optical or only in the near-IR (but not both) in the very limited regions (if any) of spatial overlap, and not at all in the inner regions of the disks.

This sample includes 10 (42%) of the of 24 currently known light-scattering debris disks, and includes disks with host stars spanning spectral types from A0 to M1, distances from ~ 10 to 100 pc, and ages from ~ 10 to 1000 Myr. These disks span a range of inclinations to the line-of-sight from edge-on ($i = 90^\circ$; the most favorable geometry for discovery detection), to nearly face-on ($i = 0^\circ$). These are the benchmark disks for thermal IR imaging and SED modeling (Lebreton et al. 2012; Donaldson et al. 2013), some of the earliest to be observed by ALMA (MacGregor et al. 2013; Walsh & Testi 2012), and anticipated for the Gemini Planet Imager (GPI) and the Spectro-Polarimetric High-contrast Exoplanet Research (SPHERE) coronagraph.

3. OBSERVATIONAL MOTIVATION/GOALS

Over the past 16 yr, *HST* coronagraphy (with PSF subtraction) has provided some remarkable, even paradigm-changing, images of CS disks through the investment of many hundreds of *HST* orbits (see the representative discovery images in Figure 1). They are, however, derived from a diversity of heterogeneously planned and executed survey programs that typically have been limited in depth, inner (stellocentric) working angle, and are characteristically “polluted” with artifacts from non-optimal PSF subtractions; this impedes unambiguous interpretation and often provides only loose constraints for well-informed modeling, making comparison within the sample and with models challenging. However, studies of individual objects have

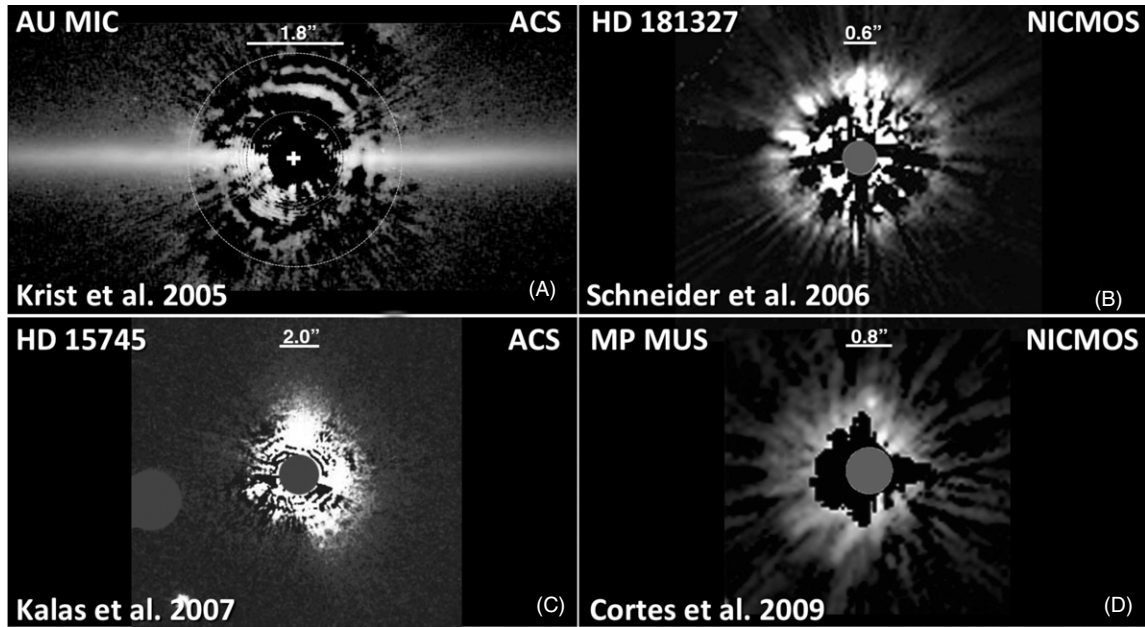


Figure 1. Prior *HST* PSF-subtracted optical (ACS) and near-IR (NICMOS) coronagraphic images of a representative sample of CS disks from prior programs are limited in image quality due to “pollution” by optical artifacts from PSF-subtraction residuals well beyond the physical limits of their image plane central coronagraphic obscurations.

demonstrated ways to improve upon the discovery imaging, both by using appropriate observational strategies (see Section 4) and advanced data reduction techniques (see Sections 5.4–5.8, and Section 6.1).

“Raw” *HST* coronagraphy alone, with any of the telescope’s coronagraphically augmented instruments, has insufficient starlight suppression to study even highly efficiently starlight-scattering CS dust without further “disk-to-star” contrast enhancement via PSF subtraction. PSF subtraction significantly enhances the imaging dynamic range (disk-to-star contrast) by two to three orders of magnitude, further removing residual starlight from the PSF halo that is incompletely suppressed by the *HST* coronagraphs alone. To enable quantitative metrical comparisons with other instrumental systems, in this paper, we specifically define “contrast” as the flux density contained within a resolution element at any stellocentric angular distance arising from the stellar PSF halo ratioed to the flux density within the central resolution element of stellar PSF core unsuppressed by coronagraphy and/or PSF subtraction. *HST* PSF-subtracted coronagraphic images are often degraded or compromised by optical artifacts resulting from the subtraction of imperfectly “matched” stellar PSF templates primarily due to chromatism (differences in photopheric SEDs), particularly under the very broad STIS unfiltered bandpass (Grady et al. 2005), and by variable focus wavefront error caused by spacecraft thermal instabilities on orbital and longer timescales (also known as “breathing” of the optical telescope assembly). Additional contributors to such artifacts that impact different spatial scales and correlation scale lengths in PSF-subtracted images are discussed by Schneider et al. (2011). To a priori minimize these in defining our GO 12228 observation plan, we adopted an approach with the following.

1. Very close color-matching (in $\Delta[B - V]$, and $\Delta[V - R]$ when possible) of PSF template stars to the science targets, as indicated in Grady et al. (2005);
2. Selection of PSF template stars to minimize target-to-template spacecraft slew distance (very similar Sun angles with attitude changes);

3. Interleaving of PSF observations with the science target in sequential, non-interruptible, orbits (as learned from experience with ACS; see Clampin et al. 2003);
4. Maintaining similar target:PSF orbit orientation (thus similar β angles at nominal roll);
5. Using six target and two PSF orbits per star to obtain data sampling of both inter/intra-orbit breathing phase;
6. Obtaining observations at different spacecraft roll angles so that the final “analysis-quality,” multi-roll-combined, image covers as much of the 360° field about the star as possible; and that several images cover the same portions of the FOV, enabling a robust separation of PSF and detector artifacts from the disk imagery;
7. Obtain/combine data from different locations (and integrated exposure depths) along the STIS coronagraphic tapered wedge to sample the disk from an IWA from $0''.3 \leq r \leq$ at least $12''.5$.

Details are given in Table 2.

4. OBSERVATIONS—PLANNING AND EXECUTION

4.1. Strategy

We obtained multiple orientation STIS coronagraphic imaging of the disks listed in Table 1, and single orientation observations otherwise obtained in the same manner, of a color-matched set of PSF template stars (Table 2), one for each science target. Our goals were to obtain high spatial resolution images of the light-scattering CS material in these disk systems over a wide stellocentric radial and azimuthal range. We imaged as close to the star as possible in stellocentric angle with the STIS occulting Wedge A (limited by the $0''.3$ half-width at the smallest, standard position of WedgeA-0.6), and also exposed deeply commensurate with a requisite “multiple-roll” strategy to minimize PSF-subtraction residuals to image tenuous low-SB dust in the outermost reaches of these CS debris systems (at WedgeA-1.0 position) previously unimaged with ACS. The multiple-orientation observing approach, importantly and simultaneously, provides both rejection (by anti-coincidence through

Table 2
Planned, Contemporaneously Interleaved PSF Template Stars

Disk Target	PSF Star	PSF Spec	PSF V	PSF $B - V$	$\Delta(BV)^a$	PSF Slew ^b	PSF ΔONR^c	
HD 15115	HD 16647	F3V	6.26	+0.37	−0.02	3°5	1°7	−8°6
HD 15745	HD 16199 ^d	F2	7.47	+0.34	−0.02	28°3	−38°3	
	HD 20499	F2	7.27	+0.35	−0.04	9°2		33°3
HD 32297	HD 33403	B9V	8.31	+0.01	−0.19	4°9	−9°8	5°6
HD 53154	HD 59780 ^e	G8III	6.52	+0.92	−0.12	4°8	...	3°6
HD 61005	HD 56161	G5IV	6.94	+0.75	−0.04	4°8	−4°5	7°1
HD 92945	HD 89585	G5III	7.41	+0.88	+0.01	9°3	0°8	0°8
HD 107146	HD 120066	G0.5V	6.30	+0.66	−0.04	23°8	17°5	11°1
HD 139664	HD 99353	F7V	5.17	+0.50	−0.10	39°7	48°4	68°55
HD 181327	HD 180134	F7V	6.37	+0.45	+0.03	1°4	0°5	0°7
AU Mic	HD 191849	M0V	7.97	+1.46	−0.02	15°1	2°9	40°9
MP Mus	CD-701017	...	9.58	+1.08	−0.07	38°4	5°6	0°5
Brt. Calib. ^f	Alpha Pic	A8V	3.30	+0.18	N/A	N/A	N/A	N/A

Notes.

^a Difference in $B - V$ color index between target and PSF template stars.

^b *HST* attitude change in slewing between target and PSF template star.

^c ONR: “off nominal roll”; ΔONR : roll angle change between the second target and PSF star visits in contemporaneous set.

^d Previously undetected close binary companion—unsuitable for PSF-subtraction template.

^e PSF template obtained at in the second of two planned sets of observations—applied to both.

^f Bright PSF template test/calibration target (not directly used for science target PSF subtractions).

median filtering) of optical artifacts that could otherwise be mis-construed as intrinsic substructures in CS disks, and completely (or nearly so) azimuthally samples “around” the otherwise obscuring STIS coronagraphic wedge and unapodized *HST* diffraction spikes. The effectiveness of this approach for debris disk imaging into sub-arcsecond IWA was demonstrated previously to lower efficacy with only two spacecraft roll angles (e.g., Schneider et al. 2009). In this study we planned six spacecraft roll angles, to provide complete azimuthal coverage at $r > 0''.3$.

4.2. STIS Coronagraphic Exposures

For optimal S/N at the edge of the $r = 0''.3$ occulting wedgeA-0.6, our individual short exposures were designed on a target by target basis to reach (but not exceed) 90% of the full-well depth of the STIS CCD from the stellar light PSF halo that significantly dominates over starlight scattered by CS dust before PSF subtraction. For this we scaled (based on stellar V and B band magnitudes) from previous observations of STIS coronagraphic PSF template stars as well as prior observations of protoplanetary and debris disk hosts. Multiple “short” exposure images (target specific integration times of a few to less than tens of seconds depending upon stellar brightness) were obtained to improve the S/N as previously demonstrated for HR 4796A per Schneider et al. (2009). We supplemented these “shortest” exposures with much deeper exposures (typically ~ 10 times each, multiply CR-split for cosmic ray rejection and compensation) to map the middle and outer regions of the disks where the CS disk SB is (in general) much lower (due to the r^{-2} dilution of the stellar radiation field), and imaging is otherwise photon, rather than contrast, limited. We additionally employ the $0''.5$ half-width position of WedgeA-1.0, advantageously using the larger unvignetted coronagraphic FOV, to further investigate external perturbations to the CS dust in the form of “tidal tails” and other structures from

posited companions and interactions with the local interstellar medium (ISM), and blow-out of small scattering particles at large stellocentric distance.

The WedgeA-0.6 images were designed to (1) image as close to the central star as possible without saturating at the edges of the wedge in the pre-PSF-subtracted images; and (2) provide full (or nearly so) azimuthal coverage unaffected by the wedge and *HST* diffraction spikes in multiple-orientation combination, to $r \leq 3''.3$. Deeper images (of necessity saturating the regions close to the star) extending to much larger angular distances (using a commensurately larger sub-array read out) were separately obtained using WedgeA-1.0. These WedgeA-1.0 images were designed to simultaneously provide (1) deep images of the outer portions of the target disks to stellocentric distances of $\geq 6''$; (2) unsaturated imaging overlapping into the smaller stellocentric regions (to $r \gtrsim 2''$) separately covered in the non-overlapping domain by the WedgeA-0.6 images to the physical limit of the occulting wedge edge. See Table 2 for target-specific details with additional information available in the as-executed *HST* Phase 2 observation plan.¹⁴

All data (following standard mode 2 coronagraphic target acquisitions) were obtained with STIS instrumental CCDGAIN = 4, and the longer (deeper) exposures with CR-SPLIT = 3–5, optimized for the length of target (declination) specific visibility windows. We employ sub-array readouts tuned to the maximum anticipated angular extent of the outer regions of the light-scattering disks detectable to the anticipated levels of the background noise to minimize dead time and maximize observational efficiency. For details, see Table 3.

4.3. STIS Coronagraphic Orbits

For each science target, we obtained a sequence of short and long coronagraphic exposures (see above), in a total of

¹⁴ <http://www.stsci.edu/hst/phase2-public/12228.pro>

Table 3
HST GO 12228 Observation and Data Log

Target (Disk/PSF)	UT Date Obs. Start	Orient ($^{\circ}$)	W0.6 # Exp	W0.6 T_{EXPT} All Visits (s)	W1.0 # Exp	W1.0 T_{EXPT} All Visits (s)	Visit Data ID ^a
HD 15115	2011 Oct 8	195.1, 220.1, 245.1	60	1026	12	4566.8	41, 42, 44
HD 16647	2011 Oct 8	221.7	24	261.6	7	1439.9	43
HD 15115	2011 Dec 19	0.1, 27.6, 30.1	60	1026	12	4866	45, 46, 48
HD 16647	2011 Dec 19	19.0	24	261.6	7	1439.9	47
HD 15745	2012 Jan 2	20.1, 40.1, 60.1	42	1436.4	9	5022.7	71, 72, 74
HD 16199 ^b	2012 Jan 2	1.7	14	521.4	3	1639.8	73
HD 15745	2012 Nov 3	105.1, 127.1, 154.1	42	1436.4	9	4199.4	B5, B6, B8
HD 20499	2012 Nov 3	160.4	11	304.7	4	1644.8	B7
HD 32297	2011 Dec 22	325.1, 355.1, 25.1	27	1652.4	9	4961.7	05, 06, 08
HD 33403	2011 Dec 22	345.3	8	576	3	1647.9	07
HD 32297	2011 Nov 6	222.1, 242.1, 262.1	27	1652.4	9	4961.7	01, 02, 04
HD 33403	2011 Nov 6	247.6	8	576	3	1647.9	03
HD 53154	2011 Jan 10	318.6, 336.1	40 ^c	688 ^c	12 ^c	3600 ^c	82, 84
HD 53154	2011 Apr 4	30.8, 48.2, 65.7	60	1032	18	5400	85, 86, 88
HD 59780 ^d	2011 Apr 4	45.7	24	333.6	8	1668	87
HD 61005	2011 Feb 17	340.1, 10.1, 20.1	30	1893	9	4748.4	55, 56, 58
HD 56161	2011 Feb 18	5.5	16	313.6	3	1773	57
HD 61005	2011 Dec 6	250.1, 270.1, 290.1	30	1893.6	9	4748.4	51, 52, 54
HD 56161	2011 Dec 6	277	16	313.6	3	1773	55
HD 92945	2012 Feb 29	25.7, 47.7, 67.7	41 ^e	1514.4 ^e	9	4098.3 ^e	61, 62, 64
HD 89585	2012 Feb 29	48.5	16	475.2	3	1614	63
HD 92945	2011 May 6	25.7, 46.7, 67.7	42	1575	9	4761	65, 66, 68
HD 89585	2011 May 6	48	16	475.2	3	1614	67
HD 107146	2011 Feb 22	192.1, 209.6, 227.1	48	1065.6	12	4764	25, 26, 28
HD 120066	2011 Feb 22	227.1	24	264	7	1421	27
HD 107146	2011 May 3	66.7, 81.7, 96.7	48	1060.8	12	4988.4	21, 22, 24
HD 120066	2011 May 3	86.7	24	264 s	7	1445	23
HD 139664	2011 Jul 31	24.1, 48.1, 66.1	48/24	115.2/1152	9	4469.4	91, 92, 94
HD 99353	2011 Jul 31	96.4	16/8	62.4/624	3	1410.0	93
HD 139664	2011 May 23	303.1, 327.1, 351.1	48/24	115.2/1152	9	4469.4	95, 96, 98
HD 99353	2011 May 23	35.6	16/8	62.4/624	3	1410.1	97
HD 181327	2011 May 20	222.7, 242.7, 262.7	60	1314	24	5124	11, 12, 14
HD 180134	2011 May 20	243.2	24	262.8	8	1656	13
HD 181327	2011 Jul 10	293.1, 313.6, 334.1	60	1314	24	5060	15, 16, 18
HD 180134	2011 Jul 10	314.2	24	285.6	8	1656	17
AU Mic	2010 Aug 9	311.1, 341.1, 10.1	36	1638	9	4887	35, 36, 38
HD 191849	2010 Aug 9	343.9	16	489.6	3	1713	37
AU Mic	2011 Jul 16	228.1, 257.1, 287.1	36	1638	9	4887	31, 32, 34
HD 191849	2011 Jul 16	298.0	16	489.6	3	1657.8	33
MP Mus	2011 Mar 3	278.1, 300.1, 321.1	21	8158.5	0 ^f	0 ^f	A1, A2, A4
CD-701017	2011 Mar 3	305.7	11	2681.1	0 ^f	0 ^f	A3
MP Mus	2011 Jun 6	359.1, 15.6, 36.1	21	8158.5	0 ^f	0 ^f	A5, A6, A8
CD-701017	2011 Jun 6	18.0	11	2681.1	0 ^f	0 ^f	A7
Alpha Pic ^g	2011 Jan 10	324.0	16/8	11.2/112	16/6	224/1530	83

Notes.

^a Visit level data set ID as assigned by MAST. GO 12228 data archived as obiw + dataset_id + *.

^b V73: close angular proximity background star and/or companion revealed—ill-suited as PSF-subtraction template.

^c V63: target acquisition failure—no data from this visit. #EXP, T_{expt} tabulated as executed, not as planned.

^d Contemporaneous (interleaved) PSF template observation only for HD 53143 only for V85–V88.

^e V62: interrupted by recoverable guide star loss-of-lock. #EXP, T_{expt} tabulated as executed, not as planned.

^f VA1–VA8: observations planned for only WedgeA-0.6 for this PMS star disk target and its PSF template.

^g V83: bright, blue, program calibration target. Obtained contemporaneously with V81–V84.

six orbits to observe at six different spacecraft roll (celestial field orientation) angles for spatial coverage, and to minimize PSF-subtraction residuals that are only partially correlated with different rolls (Schneider et. al. 2009). We “broke” the six orbits into two sets of three sequentially executed orbits with,

ideally (but seldom realized), relative orientations -30° , 0° , and $+30^{\circ}$ (with loose tolerance) from off-nominal roll. The two sets of three orbits were each scheduled (at different epochs) at nominal roll orientations differing by 90° – 150° to fully sample, (or nearly so) azimuthally beyond the $r = 0.3$

edge of the coronagraphic wedge, and probe beyond the $r = 3''/4$ edge of the STIS FOV on the narrow side of the wedge's taper in the WedgeA-0.6 position. This also had the advantage, by observing at two epochs and thus different Sun and β angles (stochastically, not by requirement), of guarding against the possibility of an anomalous “breathing” phase in a single set of observations that can otherwise significantly degrade PSF subtraction. We imposed no absolute orientation constraints on the three-orbit sets, even in cases where the celestial orientation of disk major axes are known. This eased the already constrained schedulability of the individual visits with no penalty or degradation for meeting our science goals. Along with each three-orientation visit set, we also identically observed, but only at nominal roll for one orbit, a bright target-color matched PSF template star to be used for contemporaneous PSF subtraction. Thus, a total of eight orbits were required for each disk–host/PSF–star pair. Hence, we consumed a total of 88 orbits to observe our 11 CS disk targets and matched PSF-subtraction template stars.

4.4. PSF Star Selection and Coronagraphic Performance

Extensive experience with *HST* broadband coronagraphy (with NICMOS $R = \Delta\lambda/\lambda_{\text{cent}} = 0.53$ with F110W and even more extreme with STIS unfiltered $R = 0.75$), clearly demonstrated the need for PSF stars very well matched in color indices under the instrumental bandpass to mitigate differential chromatic effects in PSF structure that otherwise hinder efficacious PSF subtraction (Grady et al. 2005; Schneider et al. 2010). Hence, whenever possible, given other constraints, we selected PSF template stars of very similar photospheric SEDs (and of sufficient brightness) well matched to each of our CS disk hosting target stars (with $[B - V] = +0.02$ to $+1.44$; see Table 2) differing by $|\Delta[B - V]| < 0.1$ for each target/PSF–template star pairs.

A second major contributor to optical artifacts in PSF-subtracted images arises from secular changes in the fine structure of the coronagraphic PSF over time (spacecraft orbit and orbital phase) and with spacecraft repointing (attitude and roll). Both differential chromatic and PSF instability effects have hindered ACS coronagraphic observations, very often having prevented even the much larger limiting IWAs of the ACS coronagraph's $r = 0''.9$ and $1''.8$ masks to have been reached as those observations had been conducted. To mitigate these issues to the greatest extent possible, along with each of our CS disk-hosting target stars, we interleaved (as the third of four sequential orbits), at nominal roll, a target-specific color-matched PSF-subtraction template star in each of the 90° – 150° rotationally offset visit sets. Whenever possible, so as not to excite “breathing” modes with thermal instabilities in the *HST* Optical Telescope Assembly (OTA), we constrained the cross-sky target/template slew distance to $< 10^\circ$ (to minimize spacecraft attitude changes), and kept roll angle differentials from target/template nominal roll as close to zero as possible; see Table 2 for details.

4.5. Details of the As-executed Observations

Table 3 provides an as-executed visit level observation summary for each disk target and its contemporaneously observed PSF star, each observed at two epochs in four contiguous *HST* orbits. In this program, each visit is comprised of a target acquisition exposure (not shown in Table 3) followed by multiple WedgeA-0.6 and WedgeA-1.0 exposures in a single orbit. At

each epoch, in chronological order, the first, second, and fourth visits are of the disk target with celestial field orientation angles scheduled as close to -30° , 0° , $+30^\circ$ from off-nominal roll as possible, given *HST* guide star selection and other scheduling constraints. The Table 3 “Orient” angles give the field angle measured from the 50CCD aperture's image +Y axis to celestial north counter-clockwise (through east). In the third visit at each epoch, the PSF star is imaged at its nominal roll angle, thus interleaving PSF observations between the second and third disk target orbits (orbits two and four in each set of target+template images). In most cases the absolute orientation angle of the spacecraft for the PSF orbit, at its nominal roll, is similar to the second disk target orbit to avoid driving breathing modes with large changes in spacecraft β angle. For each visit, Columns 4 and 6 in Table 3 indicate the total number of WedgeA-0.6 and WedgeA-1.0 exposures taken in the three same-epoch disk visits (or single PSF visit) followed in Columns 5 and 7 by the corresponding total integration time. In detail, from visit-to-visit and between the two epochs, the exposure times in particular for the WedgeA-1.0 images are not always identical for the same targets due to variances in target visibility intervals at different times of the year and at different celestial orientations. A detailed “phase 2” exposure-level observation plan, and orbit layout, with a higher level of detail than summarized in Table 3 is available from STScI.¹⁵

Each visit has a unique identifier (Table 3, Column 8), monotonic for epochal target/template pair, with numbers 1–4 or 5–8 corresponding to the first through fourth visits at that epoch. Throughout this paper, we refer to visit numbers as V## in accordance with the Column 8 visit data ID. The raw and pipeline processed exposure level data archived in the Mikulski Archive for Space Telescopes (MAST) are entered as program OBIW## followed by these visit-level data identifiers.

5. IMAGE DATA CALIBRATION, REDUCTION, AND PSF SUBTRACTION

5.1. Exposure-level Basic Instrumental Calibration

The GO 12228 observations were executed between 2011 May and 2012 November. All raw image data were retrieved soon after the execution of each target/epoch visit set from MAST. Basic image data calibration of each individual exposure in the raw data frames, including bias correction, dark subtraction, flat-fielding, linearity correction, and identification of instrumentally deficient and saturated pixels, was performed using the STSDAS *calstisa* software¹⁶ as routinely used in the *HST*/OPUS pipeline. Initial basic instrumental pipeline processing is usually done by STScI immediately following the epoch of observation with bias and dark reference files that often are “stale” and do not best reflect the observationally contemporaneous state of the instrument. Thus, prior to final reduction and PSF subtraction, all raw images were re-calibrated using later-obtained (by STScI as part of the Cycle 19/20 calibration program) bias and dark reference data closer to the dates of actual observations subsequently provided to us through STScI's Calibration Database System.¹⁷

5.2. Visit-level Reduction

All instrumentally calibrated frames were visually inspected to check for (compensate if possible, or vet if necessary) any

¹⁵ <http://www.stsci.edu/hst/phase2-public/12228.pdf>

¹⁶ <http://www.stsci.edu/hst/stis/software/>

¹⁷ <http://www.stsci.edu/hst/observatory/cdb>

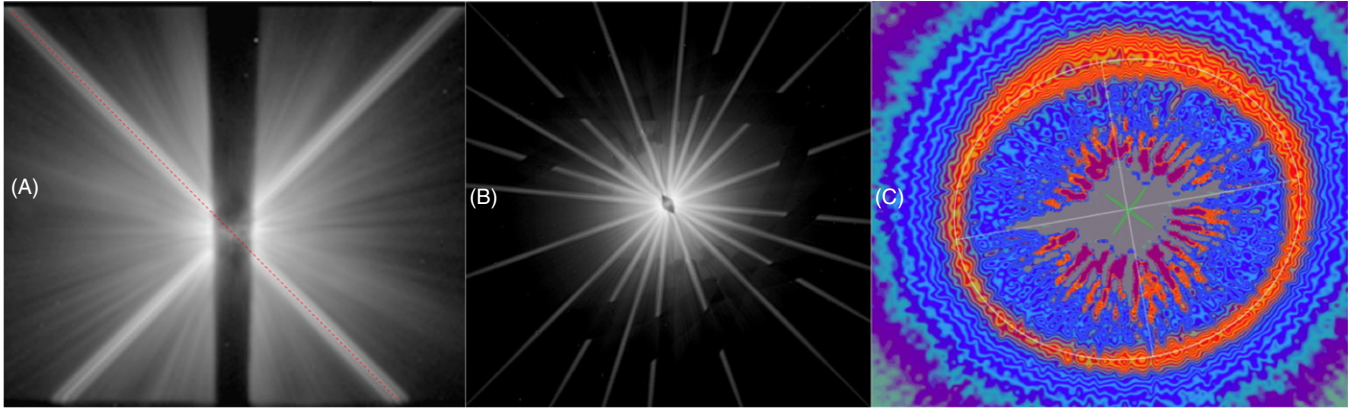


Figure 2. Panel (A): representative STIS coronagraphic stellar PSF (WedgeA-0.6 location) showing the best fit to the ridge of pixels along one of the two sets of orthogonal diffraction spikes that “point back” to the star (the other spikes are unannotated to better show their morphology). The intersection of the two linear fits locates the star to a precision typically of $\sim 4\text{--}5$ mas for the depth of integration systematically used in individual images (differently for WedgeA-0.6 and -1.0) throughout the GO 12228 program. Panel (B): in two-wedge combined multi-roll imaging (shown here before PSF subtraction for illustrative purposes), the multiple diffraction spikes in co-aligned images allow maintaining a systematic precision in co-registration at the same $4\text{--}5$ mas level. Panel (C): precise determination of the star location (green marker) enables high-precision differential astrometric measures and evaluating the significance of the measured positions of disk substructures and features with respect to the disk-hosting star, such as photocentrically determined offsets in CS debris rings (in this case for HD 181327 by elliptical isophote fitting with the best fit to the elliptical annulus of peak radial SB and corresponding minor/major axes overlaid; see Section A.1).

(A color version of this figure is available in the online journal.)

significant imaging performance anomalies such as exceptionally high cosmic-ray event hit rates (e.g., in the V45 third WedgeA-1.0 exposure) and loss of twoFGS fine-lock pointing control (e.g., in the V62 loss of two WedgeA-0.6 exposures and curtailment of one WedgeA-1.0 exposure). Sky backgrounds were estimated from manually selected regions in the periphery of the field far from the target star/disk and not coincident with any stars/galaxies in the field, and subtracted. Within each visit, the multiple-exposure data obtained at each of the two Wedge A occulter positions were separately median combined for cosmic-ray rejection and diminishment of random instrumental and photon noise compared to stellar/disk signal, and then converted to count rate images based on shutter open times. The visit-level median combined count rate images were then screened for background biases statistically unequal to zero within the pixel-to-pixel noise in the field periphery (and corrected in rare cases where outliers were detected). These and subsequent steps in image processing and later analysis were carried out using the IDL-based IDP3 software (Stobie et al. 2006).

5.3. Absolute Photometric Calibration

Throughout this paper, we adopt STScI’s absolute photometric calibration of the STIS unfiltered 50CCD instrumental response, applicable for all coronagraphic imaging, as codified through the STSDAS synphot/calcpot software package and calibration database reference files. We summarize in Table 4.

In terms of SB, then, for an HRC astrometric (“plate”) scale of $0.05077'' \text{ pixel}^{-1}$, then $1 \text{ count s}^{-1} \text{ pixel}^{-1} = 0.1765 \text{ mJy arcsec}^{-2}$.

5.4. Locating the Target Star

Because of systematic limitations in the *HST* on board autonomous target acquisition process, the final targeting slew to place the target star at the desired location in the STIS 50CCD field can deviate by $\sim \pm 1$ pixel (50 mas) from a fixed fiducial position with respect to the grid of detector pixels, and typically by $\sim \pm 0.25$ pixels from the deployable focal plane coronagraphic mask nominal target location. Thus, prior to

Table 4
STIS Coronagraphic Mode Absolute Photometric Calibration

Calibration Parameter	Parameter Value
Full spectral passband	2000–10500 Å
Pivot wavelength (λ_p)	5752 Å; $F_\lambda = g(F_v * c)/\lambda_p^2$
FWHM of unfiltered passband	4330 Å
Instrumental sensitivity ^a	1 count (ADU gain = 4.096) $\text{s}^{-1} = 4.55\text{E-}7 \text{ Jy}$ or = AB mag 26.386 (with gain = $4.096 \text{ e}^- \text{ ADU}^{-1}$)
(AB) Zero-point magnitude	3671 Jy

Note. ^a For a spectrally flat source.

later two Wedge position and multi-roll image combination, and subsequent star/disk relative astrometric measures, the location of the coronagraphically occulted target star in a visit-level median-combined science image aperture frame (SIAF) must be determined to astrometrically high precision. No direct (unocculted) image of the star itself (other than a target acquisition image at a different field position and with an attenuating filter in place, both insufficient for this purpose) was possible at its post-slew position used for science imaging due to saturation limits from the bright target stars. Thus, we used the previously demonstrated “X marks the spot” method (see Schneider et al. 2009, Section 4.2) to determine the location of the coronagraphically occulted star. We use the bright (but unsaturated) orthogonal *HST* diffraction spikes, unsuppressed by the STIS coronagraph, that are rotated $\sim \pm 45^\circ$ from the occulting wedge axis (see Figure 2(A)) and with respect to the STIS 50CCD SIAF x/y pixel grid.

The target star position in each visit-level frame is typically located with an experimentally determined precision with respect to the SIAF frame of $\sim 4\text{--}5$ mas (~ 0.1 pixels) rms. These diffraction-spike-determined stellar centroids are subsequently used to initially co-register intra-visit target images acquired at their differing WedgeA-0.6 and -1.0 positions. They are also used (following PSF-subtraction) for inter-visit (multiple-roll) image alignment at a common-origin, astrometrically determined stellocentric reference position (see Figure 2(B)) against

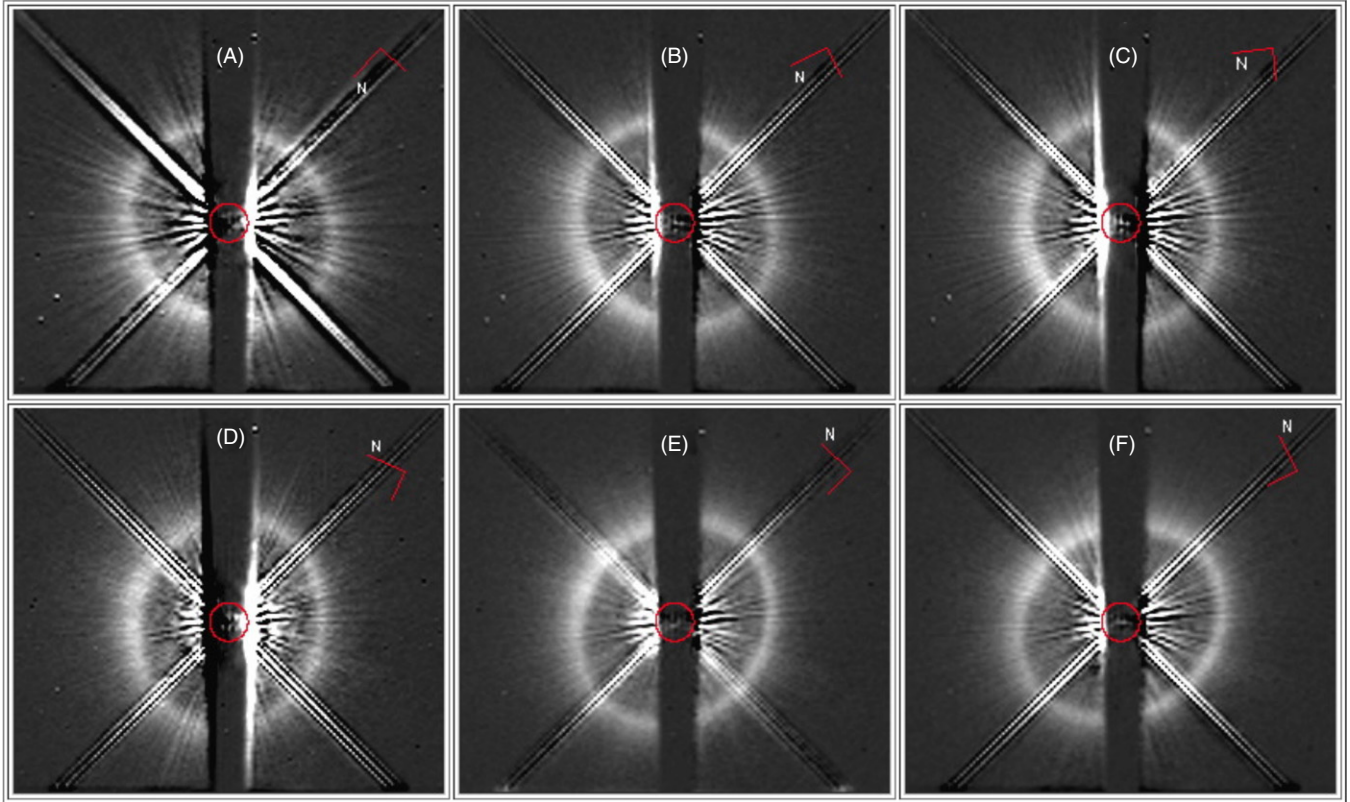


Figure 3. WedgeA-0.6 visit-level PSF subtractions—HD 181327. Top ((A)–(C)): visits 11, 12, and 14. Bottom ((D)–(F)): visits 15, 16, and 18. Top and bottom images use contemporaneously interleaved Visit 13 and 17 HD 180134 PSF templates, respectively (see Table 3 for details). $6''.2 \times 6''.0$ sub-array regions shown with a linear display from -1.0 (black) to $+7.0$ (white) counts $\text{s}^{-1} \text{ pixel}^{-1}$. The central red circle overlaid is $r = 0''.3$ centered on the occulted position of the star.

(A color version of this figure is available in the online journal.)

which disk substructures can be located with respect to the position of the host star (e.g., Figure 2(C)) with high precision.

5.5. Visit-level PSF Subtractions

The subtraction of PSF template images from those of the disk-hosting stars follows the process described in detail by Schneider et al. (2009). The iterative global PSF-subtraction process, treating template brightness and X/Y positions as free parameters, is initially done with target and template images in the detector frame orientation (so disk features rotate in the PSF-subtracted frames from visit to visit at different spacecraft roll angles). A representative full set of visit-level PSF-subtracted images (optimized to globally minimize residuals from PSF-subtraction) is shown for HD 181327 WedgeA-0.6 and -1.0 in Figures 3 and 4, respectively.

Not all visits for a given target result in PSF-subtracted images with the same image quality. For example, in Figure 3 the radial “tendrill”-like artifacts that typify differences in *HST* “breathing” (driving differential wavefront focus errors) are the most absent in the visit 16 (panel (E)) PSF-subtracted image of the HD 181327 debris ring, while the visit 11 (panel (A)) image is most affected by a breathing phase differential. It is not unusual for the first visit in an imaging set to have larger breathing residuals, in particular if the thermal state of the telescope is significantly changing following a slew from a different location in the sky from a preceding observation. The relatively low level of dispersion seen among the two sets of six HD 181327 images (WedgeA-0.6 and -1.0) is typical. However, significantly degraded results, despite all efforts for

“PSF matching” can occur with episodes of large “breathing” excursions.

Besides “breathing” residuals, the on board autonomous coronagraphic target acquisition process attempts to place targets on the mid-line of the occulting wedge. Because of the repeatability limitations in the STIS slit wheel that deploys a retractable focal plane mask used for coronagraphic observations, target placement “behind” the wedge can vary by typically $\sim 1/4$ pixels on initial placement in each visit, as we see in the GO 12228 data and as previously reported by Kriss et al. (2013). As a result, there is often a brightness differential from the PSF halo at the opposite side wedge edges that can change both in brightness and parity from acquisition to acquisition. With PSF subtraction, this can result in positive and negative edge gradients that can extend over several pixels, thus increasing the innermost working angle otherwise possible beyond the physical ($0''.3$ from the mid-line) limit of the WedgeA-0.6 mask, and the reach of the image saturation region with WedgeA-1.0. This can be seen, for example, in the visit 14 imaging of the HD 181327 debris ring (Figure 3, panel (C)) where wedge-edge brightness gradients saturate the regions immediately exterior to the wedge edges a few pixels beyond the $r = 0''.3$ physical limit of the wedge. In the visit 15 imaging (Figure 3, panel (D)) exactly the opposite parity (but with similar reach) is seen—indicative of a relative displacement of the target and template stars with respect to the mid-line of the wedge. Finally, in this regard, the visit 16 imaging (Figure 3, panel (E)) had no perceptible difference in target/template star relative decentering. Serendipitously, as previously noted, the visit 16 image was also least subject to breathing effects, so in combination this represents

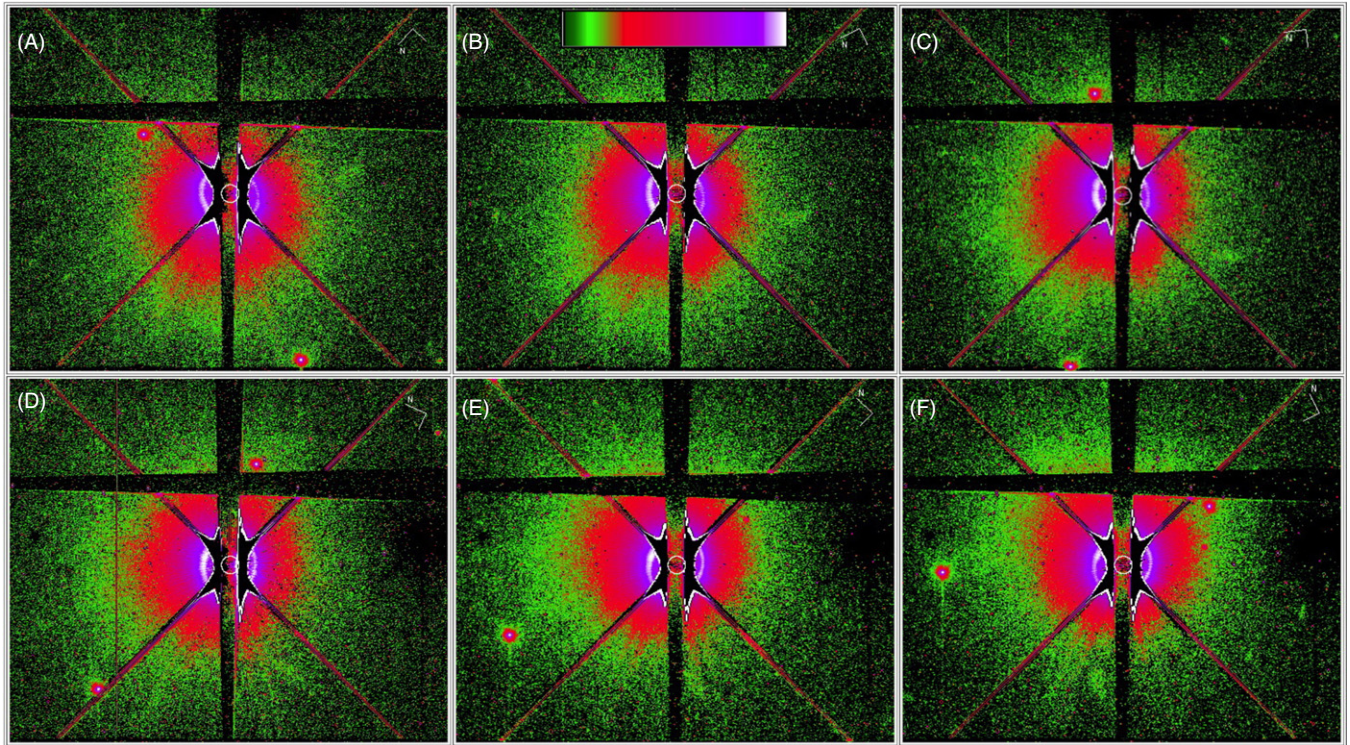


Figure 4. WedgeA-1.0 visit-level PSF subtractions—HD 181327. Top ((A)–(C)): visits 11, 12, and 14. Bottom ((D)–(F)): visits 15, 16, and 17. Top and bottom images use contemporaneously interleaved visit 13 and 17 HD 180134 PSF template, respectively (see Table 3 for details). $25''.3 \times 21''.0$ sub-array regions shown with a \log_{10} display from $[-2.0]$ to $[+0.75]$ counts $\text{s}^{-1} \text{pixel}^{-1} \text{dex}$ (inset color bar). The central white circle overlaid is $r = 0''.5$ at the occulted position of the star. The flanking saturated regions (black) extend beyond the edge of the wedge.

(A color version of this figure is available in the online journal.)

a near best-case for PSF-subtraction performance in the face of instrumentally induced residuals due to both breathing and decentering.

5.6. Chromatic Correction and Self-referential PSF Templates

Despite our best efforts in the selection of PSF template stars with “color matching” of optical color indices from catalog photometry, chromatic residuals in PSF subtraction can sometimes still appear. In such cases, the dominant stellocentric residual patterns manifest themselves as intensity modulated (bright and dark) circular “rings” that are readily apparent in debris systems viewed edge-on or nearly so where there is little or no disk flux in the chromatically affected off-disk background. These annular-like chromatic residuals are invariant in the frame of the detector, as the (spacecraft) orientation angle on the sky is changed (i.e., where the residuals are revealed in azimuthal angle relative to the position angle (P.A.) of the disk on the sky). By using disk-masked image combinations at multiple rolls to map and characterize these residuals in median combination, an empirical chromatic correction to apply after PSF subtraction can be obtained. We have done this in the cases of HD 32297 (Section A.3) and AU Mic (Section A.10) where we discuss this in detail. In the case of the edge-on AU Mic disk, we investigated and compared this approach to a “similar” one alternatively using the off-disk portions of the disk-masked AU Mic halo (including one with the disk fully “hidden” behind the occulting wedge) as its own, composited PSF template—thus fully eliminating any differential chromatic effects. We also discuss in Section A.10 the efficacy of this approach in light of potential concerns of “self-subtraction” of the disk if incompletely or imperfectly masked in the creation of a composite PSF template also.

5.7. WedgeA-0.6 and -1.0 and Multiple-roll Image Combination

For each target and separately its matched PSF template, whenever possible, we fully combine all WedgeA-0.6 and -1.0 visit-level count-rate images from all successfully executed visits. On an individual target basis, however, we reject images (or visits) that suffer performance degradation due to anomalous breathing excursions or other operational issues. Such cases are noted in the Table 3 footnotes, and discussed for individual targets in the Appendix.

With the optimization of the PSF-subtraction free parameters (ΔX , ΔY , and template brightness scaling) established at the visit-level in the SIAF image orientation, all WedgeA-0.6 and -1.0 images are translated to a common stellocentric origin based upon target image “X marks the spot” centroids, and rotated to a common celestial (“north up”) orientation (based upon down-linked engineering telemetry derived from the *HST* pointing control system). Image translation and rotation is accomplished by remapping the original images onto an output pixel grid via bi-cubic pixel interpolation apodized by a sinc-function of kernel width commensurate with the ~ 60 mas FWHM of the STIS/50CCD PSF (to prevent “ringing” with less than critical sampling in some pixels). Once all images are co-aligned and in a common celestial frame and PSF-subtracted, image-specific digital masks were created (“manually” crafted for each PSF-subtracted image using the IDP3 software “mask builder”) to individually exclude “bad” pixels from each image that were (1) obscured by the STIS coronagraphic wedges A or B, (2) corrupted by OTA diffraction spikes, (3) saturated close to the star (typically as planned in WedgeA-1.0 imaging, or due to non-ideal target centering), or (4) suspect in calibratability at or very

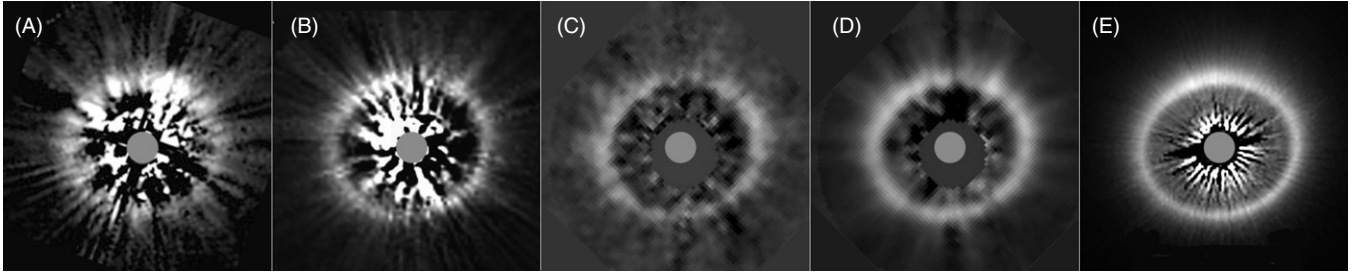


Figure 5. Comparison of PSF-subtraction methods revealing the HD 181327 debris ring in *HST* coronagraphic images with ((A)–(D)) derived from the same raw data (NICMOS 1.1 μm imagery) and (E) STIS six-roll (6R) contemporaneous observationally matched-PSF template-subtracted coronagraphy (PSFTSC). (A) NICMOS discovery image using 2 (of 10) non-contemporaneously observed PSF template stars (Schneider et al. 2006; *HST* GO program 10177). (B) “LAPLACE” (*HST* AR program 11279) re-processing and globally optimized re-reduction with PSF-matching from a down-selected 53 template ensemble (Schneider et al. 2011). (C) LOCI re-processing (without regularization) with a 232 LAPLACE recalibrated PSF template library. (D) KLIP re-processing (35 coefficients) using same PSF template library as panel (C) with regularization. (C) and (D) from Soummer et al. (2012; *HST* AR program 12652). (E) STIS 6R-PSFTSC (the result from this program is discussed in this paper). The gray circle indicates the location and size of the NICMOS $r = 0''.3$ coronagraphic circular obscuration.

near the physical edge of occulting wedge. (For an idealized illustration of (1) and (2) only using only a single wedge and optimal differential rolls, see, e.g., Figure 2 of Debes et al. 2013). These artifact-masked images are then median combined into an analysis quality (AQ) data image comprised of either just the WedgeA-0.6 or WedgeA-1.0 input images, or both, utilizing all unrejected data.

In final AQ image combination using both Wedge A-0.6 and -1.0 data, we typically truncate the radial extent of the WedgeA-0.6 images at a (target-dependent) radius of ~ 45 – 60 pixels ($r \approx 2''.5$ – $3''$) beyond which the short exposure time WedgeA-0.6 images typically become photon-starved and would contribute only noise in comparison to 10–20 times more deeply exposed WedgeA-1.0 images. With less complete spatial coverage, we also separately create “epoch 1” and “epoch 2” AQ images to enable comparative analysis for temporal variations between the two epochs within the disk (e.g., see MP Mus, Section A.11) or for identification of background objects with non-common proper motions.

In addition to these AQ images themselves, for analysis purposes, we also propagate and keep in ancillary images for every pixel (1) the number of images contributing to every output pixel, (2) the standard deviation about the masked-median median, and (with the matched data masks), and (3) the total exposure time in each pixel contributing to the final AQ images.

5.8. Mitigation of Differential Residuals from Quasi-static Wavefront Errors

In this paper, all images presented and results discussed are derived from the AQ images processed as described in Sections 5.1–5.7. In some cases, some additional improvement can be realized by the identification and further rejection of still-remaining PSF-subtraction residuals arising from identifiable quasi-static wavefront errors (i.e., from “breathing” of the *HST* optical telescope assembly). A multi-roll residual removal routine (MRRR) for doing so is presented, discussed, and demonstrated in detail by Stark et al. (2014). MRRR is enabled by the multi-roll correlation of the imprinted quasi-static PSF-residual patterns in the instrument (detector) frame oriented images. The applicability and efficacy of the MRRR approach is dependent upon the target scene and the roll-angle overlapping of unmasked field coverage at each roll. MRRR is best for disks where such imprinted patterns are independently detectable with significance in the relative absence of disk flux in the

instrument-frame images (before rotating to a common celestial orientation). In such images while the disk flux is re-oriented on the detector in the images acquired at different spacecraft rolls, the quasi-static PSF structure remains rotationally invariant. In the context of the well-suited observations of the HD 181327 debris ring, following the application of MRRR, Stark et al., *ibid*, found a reduction in the photometric measurement uncertainties along its debris ring of $\sim 5\%$.

6. STIS MULTI-ROLL PSFTFC OBSERVATIONS—COMPARATIVE CONTEXT

Herein we comparatively summarize some of the key (and unique) attributes of the STIS multi-roll PSFTSC observations in the context of prior *HST*, ground-based, and thermal IR observations of the CS disks observed in the GO 12228 sample.

6.1. Comparison with *HST*/NICMOS Reprocessed Archival Data

Recent advances in post-processing techniques such as the application of locally optimized combinations of images (LOCI), adaptive filtering via principal value decomposition, and recombination, e.g., with the Karhunen–Loeve eigenimage projection (KLIP) method (e.g., see Soummer et al. 2012) have demonstrated the ability to achieve incremental improvements in image quality through the partial identification and rejection of PSF-subtraction residuals even over more aggressive approaches in “classical” (global) PSF-subtraction such as have been enabled by the NICMOS Legacy Archive PSF Library (Schneider et al. 2011). We illustrate in Figure 5 with a comparison of the discovery image of the HD 181327 debris ring (panel (A)) subsequent image processing improvements for the NICMOS data with increasingly aggressive post-processing methods (panels (B)–(D)), and our STIS six-roll PSFTSC image (panel (E)). In the STIS image, for the first time, we clearly resolve (1) a crisp inner edge to the debris ring; (2) demonstrate its difference from the diffuse outer edge; (3) detect scattered light interior to the ring; and (4) measure brightness asymmetries around the ring that are not preserved in the LOCI or KLIP re-processed data; see Section A.9. Local optimization methods enhance the visibility of the ring, however, they (1) do not conserve low-SB flux density in spatially extended regions, and (2) require a large, observationally diverse “library” of template coronagraphic PSFs that simply do not yet exist for *HST* optical images.

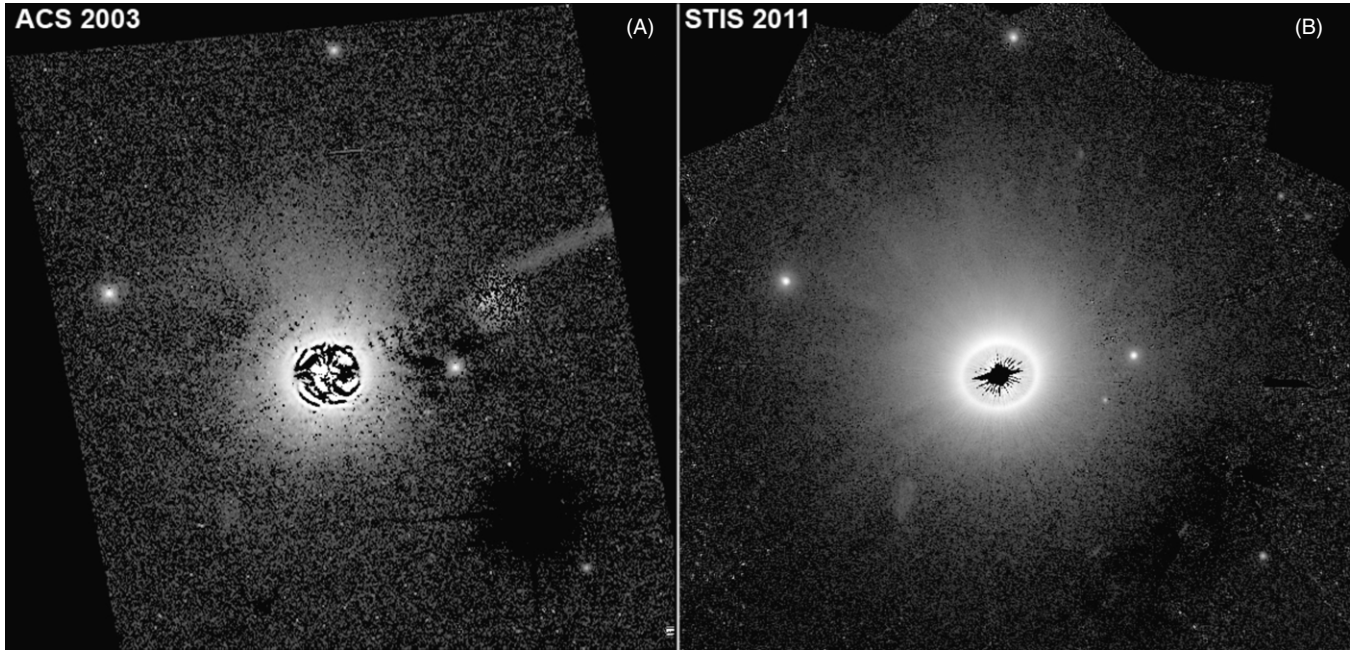


Figure 6. Smallest of the ACS coronagraphic masks (used in the HD 181327 scattered-light disk co-discovery imaging, panel (A)) provides $r \leq 0''.9$ IWA obscuration. However, ACS PSF-subtraction residuals completely dominate the disk light interior to the radius of the peak SB of the ring ($r = 1''.7$) revealed by NICMOS PSFTSC (see Figure 1(B)) using various methods of different aggressiveness. The residuals seen in the ACS image are largely suppressed to much smaller stellocentric angles with STIS 6R/PSFTSC beyond the effective six-roll combined WedgeA-0.6 limiting, $IWA_{\text{effective}}$ of $r = 0''.3$. Here, deeper (and better sampled) imaging of the outer portions of the disk with STIS 6R/PSFTSC (panel (B)) simultaneously provides high-fidelity, high-S/N imaging of the outer portions' low SB of the disk, the bright debris ring, and the largely cleared region in the ring interior. Both images: full FOV = $25'' \times 25''$, north up, east left.

6.2. Comparison with *HST*/ACS Observations

While the large angular extent of the ACS coronagraphic masks ($r = 0''.9$ and $r = 1''.8$) preclude CS observations at small IWAs, the instrument (unlike NICMOS) does provide a coronagraphic FOV comparable to STIS. This has been used advantageously for (angularly) large CS debris systems, but is significantly less efficient than STIS. For example, the full extent of the HD 181327 debris system (unseen with NICMOS) is revealed with STIS six-roll PSFTCS and compared in Figure 6, over the full dynamic range of imaging sensitivity, to a discovery epoch PSF-subtracted ACS image at a very similar central wavelength (from Schneider et al. 2006). Nebulosity in the STIS image is traced to stellocentric distances of $9''.5$ with more complete sampling about the star, better image fidelity, and higher sensitivity to low-SB light-scattering material in the outermost, photon-limited portions of the HD 181327 debris system. The STIS instrument's near full-throughput pupil (compared to $\sim 50\%$ for ACS in its coronagraphic mode), and unfiltered spectral sensitivity ($\Delta\lambda/\lambda = 75\%$, compared to 25% for ACS/F606W) together provide a gain in exposure depth of approximately six times per unit integration time. For these STIS observations, with the additional investment in exposure time of a factor ≈ 4.5 times over the ACS images, an improvement in exposure depth by a factor of ≈ 27 times, and photon-limited S/N of ≈ 5.2 times was realized in the outer-disk photon-limited regime (as closely predicted by the STSDAS/SYNPHOT synthetic photometry instrument models). STIS, however, having only an unfiltered coronagraphic mode, provides no diagnostic “color” information as with ACS multi-band imaging.

The improvement in image quality and gains in S/N with STIS multi-roll PSFTFC enable higher-precision measurements of the spatial and SB distributions of the CS materials in, and possibly escaping from, these debris system with more

comprehensive, and better observationally constrained, analysis and modeling than possible with the original discovery images alone. For example, in the case of HD 181327, the differences now observable in scattering phase functions in the ring and in the outer halo further suggest the presence of an unseen, planetary-mass perturber that may lie exterior to the ring. Additionally, these improved data have resulted in the putative discovery of a debris “spiral,” or spray of ejecta, from a posited recent massive collision in the exo-Kuiper Belt of the HD 181327 debris system. Both of these findings are discussed in detail in Stark et al. (2014), with the observational data presented for this and the other disks in this sample herein. With these data, we are now at the point of directly imaging phenomena postulated as important in the evolution of our solar system, bringing the analysis of exoplanetary systems with *HST* into a new regime.

6.3. Comparison with Ground-based Observations

Ground-based coronagraphs on large, AO-augmented telescopes are currently capable of revealing the brightest features of CS debris systems that possess strong local contrast gradients (rather than diffuse, extended flux). This is well exemplified with the HD 61005 debris system. The HD 61005 (“the Moth”) debris system was first imaged with *HST*/NICMOS (Hines et al. 2007), and its brightest features with strong SB gradients have now been imaged from the ground. The NICMOS discovery image showed a bright, but diffuse, skirt of material apparently “blown” off an inner disk viewed nearly edge-on. The skirt of the escaping material could be traced $6''$ from the star, with a brightness (inferred surface density) enhancement posited from a bow-shock on the “front” edge of the debris system in the direction of its relative motion running into an ISM wind. ACS optical coronagraphic polarimetry better revealed the outer portions of

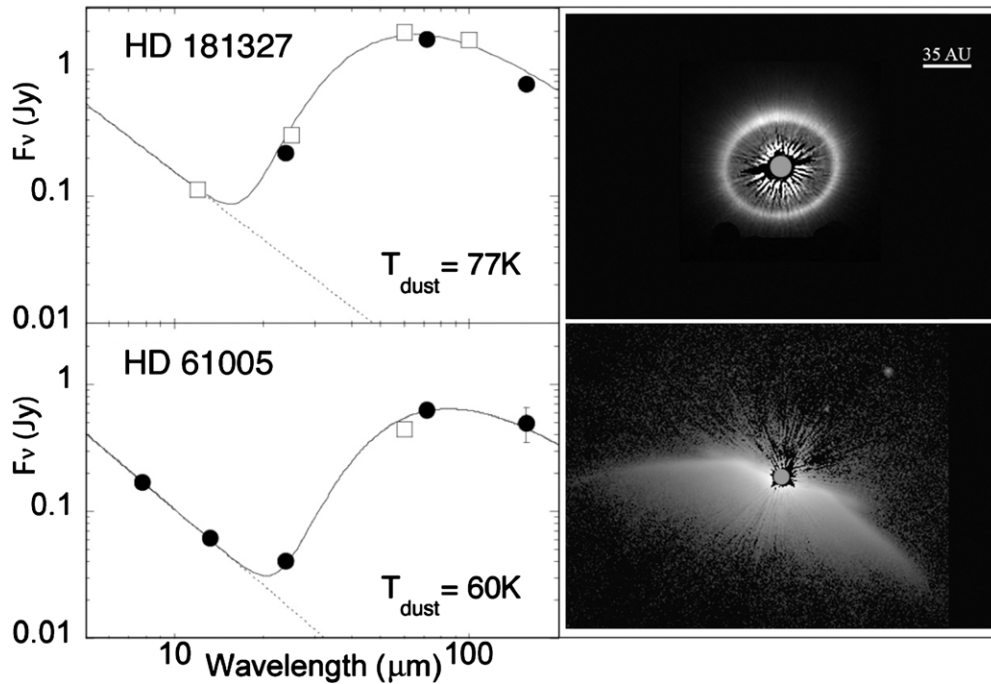


Figure 7. Photometric measures of thermal IR excesses (fit to single temperature blackbodies) from the HD 181327 (top) and HD 61005 (bottom) debris systems exhibit close similarity in their SEDs (left), though very different morphologies are revealed with STIS 6R/PSFTSC scattered-light imaging (right).

the debris disk, the fan of material, and the “limb-brightened” edge of the blow-out fan. Maness et al. (2009) found the optical-to-NIR disk color to be blue, and the polarization properties of the outer parts of the disk and the skirt were different; in particular, with the skirt dominated by small, compact, light-scattering particles. Subsequently, Buenzli et al. (2010) imaged the brightest (lowest contrast) features of the HD 61005 debris system using VLT/NaCo with angular differential imaging (ADI) and LOCI image processing, resolving the disk as a highly inclined debris ring, and measured its apparent eccentricity, and thus inclination, as 84.3 ± 1.0 and suggested a 2.75 ± 0.85 AU photocentric offset. A crisp inner edge to the ring was inferred, but the skirt of the blow-out material was not observed. Our STIS six-roll PSFTSC imagery (see Section A.5) fully discloses the skirt and the forward (higher optical-depth) limb-brightened bow-shock, resolves both the “forward” and “back” sides of the debris ring (similar on its bright side to that seen in Buenzli et al. 2010), the ring inner clearing, recovers the photocentric offset, and reveals, for the first time, that the disk is a three-component system. In addition to the ring-like disk and fan, there is an additional disk component of material beyond the launch point for the fan, which is interpreted as an analog to the scattered Kuiper Belt (D. C. Hines et al. 2014, in preparation). We also image the interior to the ring, finding a largely cleared zone down to 10 AU from the star. The debris ring, as seen in the STIS imagery, is coincident with the planetesimal belt resolved at sub-millimeter wavelengths (Ricarte et al. 2013).

6.4. “Comparison” with Thermal IR Excess as a Predictor of Optical Brightness

IR excess emission above stellar photospheric levels is the signature of re-radiating CS dust attributable to orbiting exoplanetary debris for the majority of systems. Thus, over the past 16 yr, *HST* coronagraphic imaging surveys designed to discover optical/near-IR light-scattering counterparts to (mostly) ther-

mally emissive debris disks have targeted nearby stars with the brightest 12–100 μm thermal IR excesses. These candidate systems were identified from space-based surveys (e.g., *Infrared Space Observatory*, *IRAS*, and *Spitzer* and *Herschel* photometry and/or SEDs). In nearly all cases, only IR excess sources with $L_{\text{IR}}/L_{\text{star}} \geq 10^{-4}$, as a presumed order of magnitude proxy to a priori unknown optical/near-IR scattering fractions ($F_{\text{disk}}/F_{\text{star}}$), were targeted with anticipation of high survey yields. This was a “conservative” selection criterion given the anticipated (but initially not well quantified) contrast-limited coronagraphic imaging sensitivity to disk-scattered starlight with *HST*’s coronagraphically augmented instruments. Collectively, however, (only) $\sim 15\%$ of candidate debris systems surveyed were found to possess CS dust with high enough SB to be detected with *HST* coronagraphy as conducted in a diversity disk-finding survey programs. Only two dozen spatially resolved, starlight-scattering debris disks emerged from 1998 to 2013, most recently with a few previously observed, initially undetected, but newly discovered with advanced re-processing techniques. This relatively low success rate underscores that thermal IR excess *alone* is not a good predictor of the fraction of starlight scattered by a CS debris disk into the observer’s line-of-sight (e.g., see Table 1 and Figure 7). IR excess data also provide no insight into how that total flux is actually spatially distributed (Booth et al. 2013) and thus visible in the face of instrumental sensitivities and systematics; i.e., the scattered-light component of most disks with $L_{\text{IR}}/L_{\text{star}} \geq 10^{-4}$ remained elusive.

For detectable disks, a loose correlation between IR excess and optical/near-IR scattering fraction exists (e.g., see Figure 8 for the debris systems observed in this sample). This is the case particularly for systems with ages ≤ 100 Myr, and excluding specific systems, like AU Mic, where radiation pressure blowout of small dust particles is negligible. On an individual target basis, with additional dependences, $L_{\text{IR}}/L_{\text{star}}$ is only a weak predictor

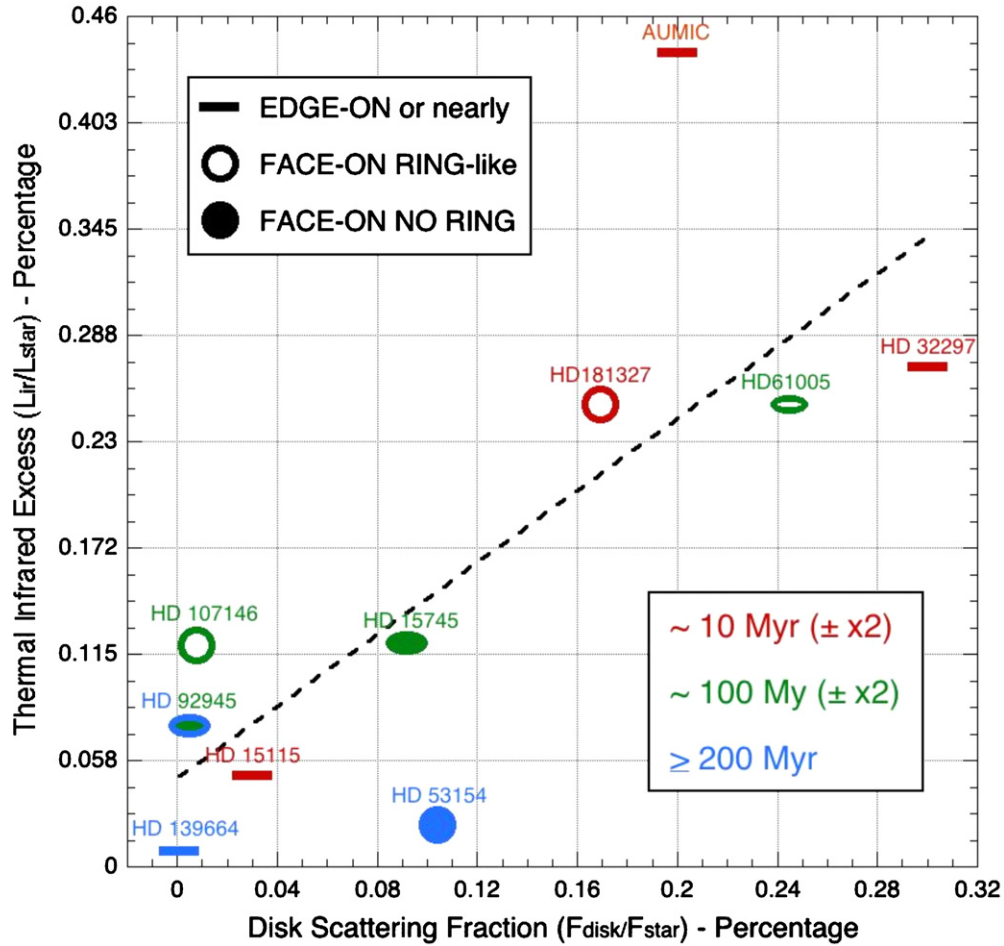


Figure 8. Overall, debris disk-scattering fractions and thermal excesses are cross-correlated with age, but with multi-parametric dependencies, $L_{\text{ir}}/L_{\text{star}}$ may be only a weak predictor of optical brightness for individual systems (e.g., HD 15115 in the GO 12228 sample).

(A color version of this figure is available in the online journal.)

of SB and detectability in spatially resolvable, light-scattering disks. Grain properties (optical constants), dust density distribution, viewing geometry of any disk, and systemic age can lower disk visibility (contrast) globally or for substructures within in a complex manner unpredictable without foreknowledge of those systemic parameters. Additionally, CS debris systems with very similar SEDs can have vastly different systemic morphologies, geometries, and dust density distributions uniquely informed by spatially resolved imaging.

7. OBSERVATIONAL RESULTS

The GO 12228 targets constitute a follow-up sample of previously discovered light-scattering CS debris systems which thus excludes “non-detections.” A gallery of disk images (discussed in detail in the [Appendix](#)) that highlight their morphological diversities is presented in Figure 9. The ability to discern with clarity substructures within these disks depends not only upon global or large spatial scale disk-to-star contrasts, but also on localized disk signal gradients in the presence of any remaining PSF-subtraction artifacts. For example, lower-contrast substructures (e.g., rings, gap edges, and perceptively edge-on mid-plane regions) are typically more readily detectable (with sufficient integration time) than diffuse, extended features. The latter remain elusive in ground-based images reliant on aggressively

advanced, locally optimized PSF-subtraction/calibration methods such as LOCI, KLIP, and (observationally) ADI that do not preserve disk flux on spatial scales larger than the dominant spatial frequencies of the detectable, lower-contrast, and spatially localized disk features. These methods, though of utility for disk imaging, simply do not preserve low-SB extended flux and typically render such features undetectable. Conversely, the gallery of GO 12228 debris disk images presented in Figure 9, derived from STIS multi-roll PSFTSC, are immune to these post-processing effects. These images may be used directly not only to characterize and quantify disk and disk grain properties, but as “truth” (within their own levels of efficacy and uncertainties) in assessing independent results from ground-based coronagraphs with advanced processing techniques that intrinsically spatially filter disk images. The disk images in Figure 9 are presented at differing angular scales due to the differences in target distances with our goal of well representing the debris system morphologies both individually and comparatively. Each disk, and its observationally derived characterizing parameters, is discussed in more detail in the [Appendix](#).

The primary morphological, geometric, and photometric characteristics of the disks, derived from these images, are summarized in Table 5. Disk inclinations (i), except for those very close to edge on or face on, and major-axis celestial P.A.s are photometrically estimated from stellocentric elliptical

Table 5
Disk Morphology, Geometry, Extent, and Brightness

Target	Morphology	Est. i ($^{\circ}$)	P.A. ($^{\circ}$)	IWD (AU)	Inner Clearing?	Radial Extent (AU)	Brightness Asymmetries	F_{disk} (mJ)	$F_{\text{disk}}/F_{\text{star}}$ (%)
HD 15115	Nearly edge-on ^a One side bifurcated (“The Needle”)	80 ^a	99.1 \pm 0.6	18	Obscured by geometry ??	E:320 W:570	E/W Extent E. side bifurcation W. side above-plane	2.04	0.030 \pm 0.004
HD 15745	Featureless (?), bi-lateral asymmetry	52	22.5 \pm 1.5	19	None to IWD limit = 19 AU	i: 255 σ : 430	Along minor axis	3.26	\sim 0.092
HD 32297	Edge-on bow concavity	Near edge-on	47.5 \pm 1.0	34	Obscured by geometry	\sim 1560 asymmetric	\square to disk plane. Disk SB asymmetry w.r.t. minor axis.	5.98	0.30 \pm 0.03
HD 53154	Symmetric inner, outer asymmetric faint “arcs”	<5	...	5.5	None to IWD limit = 5.5 AU	i: 37 σ : 70	No at <3” > 4” NW gap/arc	6.92	\sim 0.104
HD 61005	Low SB skirt (“The Moth”)	85 \pm \sim 1	70.6 \pm 1.2	14	Yes. Inner edge at \sim 45 AU	\sim 330	\square to disk plane. Ring SB asymmetry w.r.t. minor axis.	4.46	0.245 \pm 0.03
HD 92945	Broad halo + Narrow inner ring	28	100 \pm 2	6.4	Yes. Inside inner of inner ring	\sim 118 ring: 59	Inner ring better seen on east side Outer W brighter	0.143	\sim 0.0051
HD 107146	Nearly face-on Broad ring	18.5 \pm \sim 2	38 \pm 3.1	11	Yes. Shallow edge @ \sim 60 AU	\sim 220 symmetric	H-G scattering phase angle only	0.404	0.0077 \pm 0.0004
HD 139664	Near edge-on	\sim 90	75.5 \pm 0.5	20	Not observable	\sim 100	E/W Peak \sim x 1.6 (S/N limited?)	0.260	\sim 0.0005
HD 181327	Inclined Narrow ring + Diffuse outer halo	30.1 \pm 1.2	102 \pm 4	18	Yes. Sharp inner edge at \sim 80 AU	\sim 460 asymmetric	Non-H-G azimuthal. Inner/outer skew. 25% ansal Δ SB	7.81	0.17 \pm 0.015
AU Mic	Edge-on with one-side “bump” and warp	90	37.8 \pm 0.2	5	Obscured by geometry	\sim 130 symmetric	Out-of-plane. Warp. Substructures	2.51	0.20 \pm 0.02
MP Mus	Asymmetric, “Featureless”	27.3 \pm 3.3	10 \pm 2	30	None to IWD limit = 30 AU	Disk: 155 Halo: > 340	H-G. Front/back temporal variability	1.74	\sim 0.68

Note. ^a Inner “half-ring” $\rightarrow \sim 80^{\circ}$, extended un-bifurcated disk \rightarrow edge-on; see Figure 12.

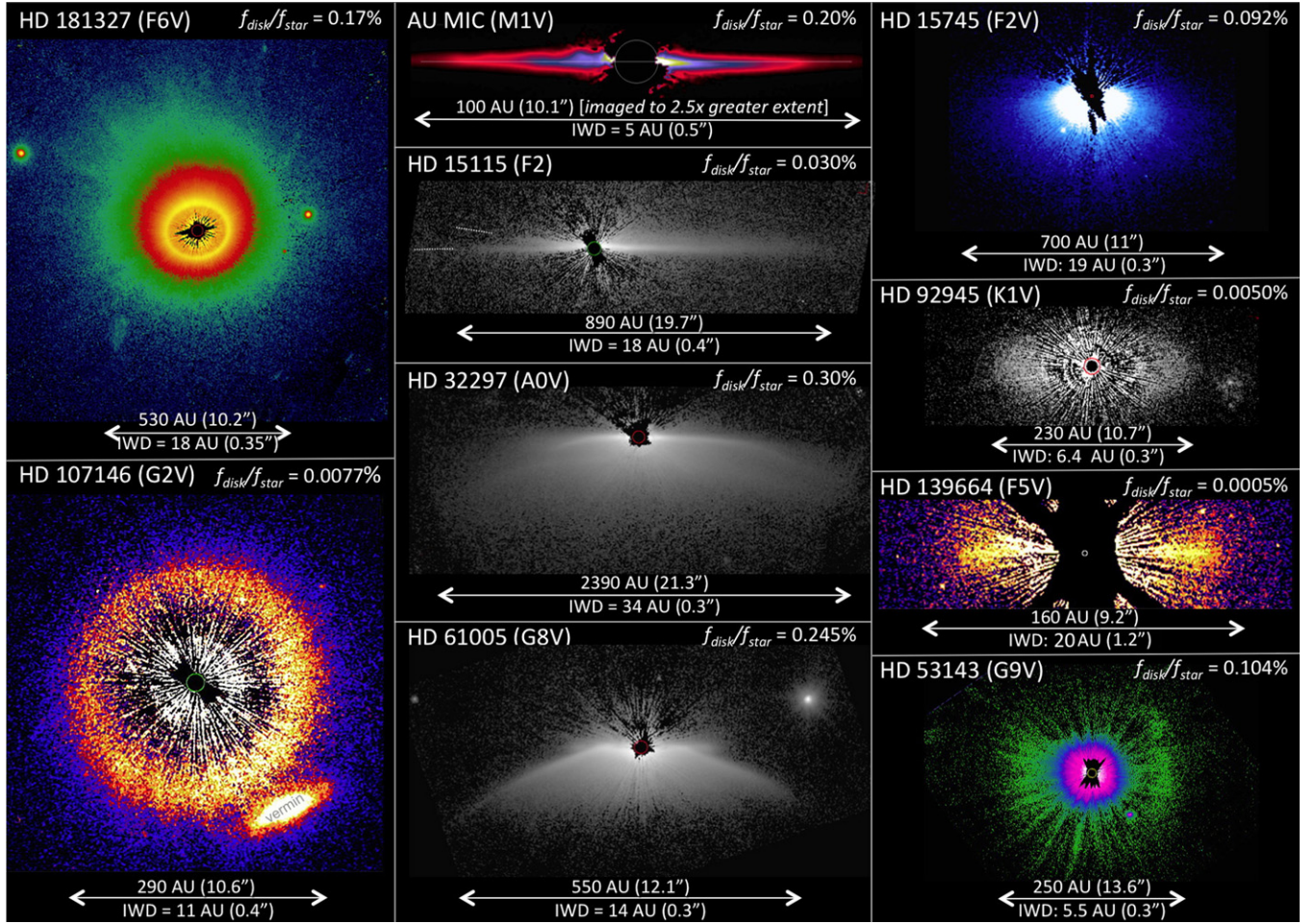


Figure 9. Analysis quality scattered-light images of the GO 12228 debris disks discussed in the [Appendix](#). The arrows indicate the full physical and angular extent of the disks (except AU Mic) in astronomical units and arcseconds (scaled differently for each disk), and below the inner working distances realized (though for all disks not at all azimuth angles) with PSFTSC imaging.

isophote fitting. Inner working distances (IWDs) are the closest stellocentric distances to which the CS brightness could be measured. The radial extent of the debris systems with large halos are expressed to where the disk flux declines to $+3\sigma$ resel^{-1} of the background far from the disk; inner (i) and outer (o) radii are separately estimated for disks with concentric components. Total (disk integrated) flux density measurements (F_{disk}) and scattering fractions ($F_{\text{disk}}/F_{\text{star}}$) at the STIS 50CCD pivot wavelength of $0.57 \mu\text{m}$ fully enclose the disks, but exclude the small areas unsampled at and interior to the IWD (e.g., see Figure 41). For F_{star} , we assume catalog V-band magnitudes, as given in Table 1, transformed into the STIS 50CCD spectral band (Table 4). Note that $F_{\text{disk}}/F_{\text{star}}$ is simply an observationally derived parameter along the line-of-sight to the observer. For moderate to highly inclined systems, the IWA could conceal a significant fraction of the potentially (otherwise) observable scattered flux, in particular if the dust is highly forward scattering.

In Figure 10, we plot the disk-to-star signal contrast as a function of stellocentric angular distance along the disk major axis where the disks are brightest (or nearly so) in CS azimuthal angle on both sides of the star. Figure 10 is presented in dimensionless contrast units (to facilitate comparison and scaling in other instrumental systems), as per the definition of image contrast in Section 4. To the extent that the CS light-scattering material may be spectrally neutral across the STIS

passband (which cannot be ascertained from these data alone, but in general deviations will only be small), these contrast curves may be closely transformed to photometric SBs at the 50CCD pivot wavelength of $0.57 \mu\text{m}$ by multiplying by $0.178 \text{ mJy arcsec}^{-2}$.

For most disks, their SBs decline radially from IWAs along the major axis of $\geq 0''.3$ to the outermost disk regions $\sim 10''$ (or where the noise level declines to $\sim 3\sigma$ resel^{-1} over the sky background) by typically ~ 4 dex. At common stellocentric distances, the diversity in debris disk SBs amongst the sample varies at the extrema from the highest to lowest contrast disks by ~ 3 dex. The highest contrast disks, e.g., HD 139664 and HD 92945, prove to be a contrast challenge to STIS PSFTSC at the smallest stellocentric distances and, in some cases, in particular with < 6 rolls, a mask-limited $\text{IWA}_{\text{minimum}} = 0''.3$ is not reached.

Figure 10 is of particular utility to those considering future coronagraph designs and disk imaging programs where instrumental performance is a prime consideration, hence it is cast in stellocentric angular distance, although each of the debris systems are at different physical distances from the Earth. Such curves may be applied directly to instrument performance models to assess disk visibilities and to derive integration time requirements for fainter disks with scalable S/N. Table 1 gives the GO 12228 target distances in parsecs to readily allow for the transformation of the stellocentric angle into astronomical units for each of the disks in Figure 10.

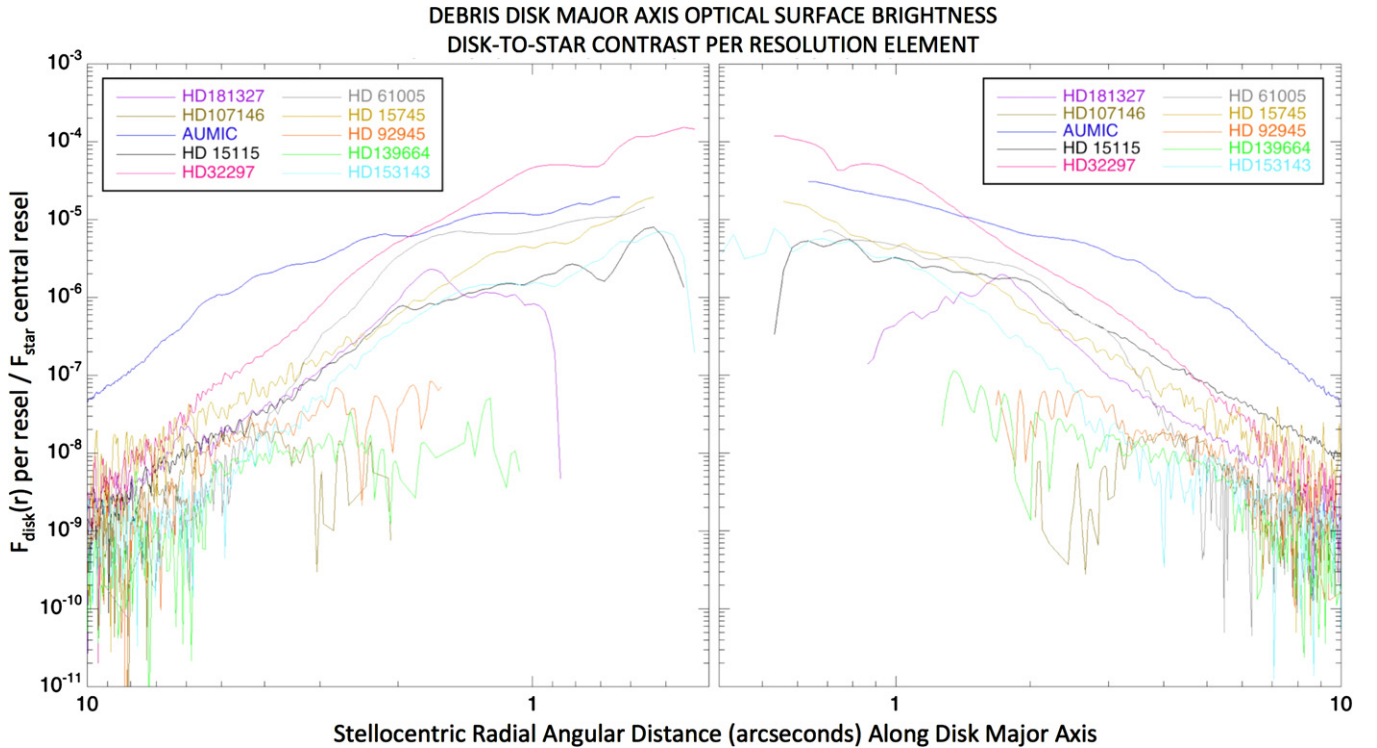


Figure 10. Radial SBs for the GO 12228 debris disks measured along the disk minor axes as a function of stellocentric distance expressed as resel^{-1} image contrasts, derived from AQ images shown in Figure 9. In detail, the minor axis may not be coincident with the smallest asymmetric IWA realized, so the smallest stellocentric distance from the star is not observed in all cases.

8. DISCUSSION

8.1. Placing *HST* Optical Coronagraphic Imagery in Context

High-fidelity ground-based imaging of CS material, in general, remains a difficult challenge even as the newest generation of “extreme” AO augmented coronagraphic imagers are coming on line. Historically, most ground-based CS disk imaging has relied on coronagraphy with polarimetric differential techniques (Hinkley et al. 2009), where light from the disk is simultaneously imaged in the orthogonal (ordinary and extraordinary) linear polarization components, with the unpolarized light of the star used as the PSF to achieve contrast augmentation; e.g., see Perrin et al. (2014b) for a comprehensive review. Images obtained in this mode can reveal the “polarized intensity” of the CS material imaged, i.e., the “ $p \cdot i$ ” product of the polarization fraction and “total intensity,” a useful diagnostic, but from which these fundamental quantities are inseparable.

Because of the nearly ideal properties (brightness, contrast, size, morphology) of the HR 4796A debris ring for high-contrast imaging, it has become somewhat of a de facto “truth test” target for ground-based AO systems (both “extreme” and less aggressive) that can be compared against either the *HST*/NICMOS NIR discovery imaging (Schneider et al. 1999) or the subsequent *HST*/STIS optical follow-up imaging (Schneider et al. 2009). For example, this debris ring has been spatially resolved from the ground with the Subaru High Contrast Instrument for Adaptive Optics (Thalmann et al. 2011), the Very Large Telescope (VLT) Naysmith Adaptive Optics System Near-Infrared Imager and Spectrograph coronagraph (NaCo; Lagrange et al. 2012), and Gemini South Near-Infrared Coronagraphic Imager (Wahhaj et al. 2014), none of which are extreme AO systems, and also very recently with the Clay Magellan AO

(MagAO; T. J. Rodigas et al. 2014, in preparation) instrument, with varying levels of fidelity.

A recent demonstration observation by GPI produced both total light and polarized intensity images of the exceptionally bright ($F_{\text{disk}}/F_{\text{star}} = 1.6 \times 10^{-3}$) HR 4796A CS debris ring (Perrin et al. 2014a) fully resolving the apparently ($p \cdot i$) brighter side of the ring beyond the instrument’s IWA at high (but uncalibrated) contrast and spatial resolution.¹⁸ These may be qualitatively compared with another first-light total intensity HR 4796A debris ring image from SPHERE.¹⁹

Today, ground-based “total intensity” imaging of CS debris disks, however, generally requires observing and/or post-processing methods, such as ADI or LOCI/KLIP, for contrast enhancement that does not preserve the flux (polarized, or in technically feasible cases, total intensity) in spatially extended regions and with shallow contrast gradients. Thus, the highest fidelity images with recoverable total-light photometry currently remains in the domain of space-based coronagraphy and, in particular, the highest efficacy yet demonstrated with STIS multi-roll PSFTSC.

With *HST*, PSF-subtraction residuals from chromatism and wavefront error instability (breathing) effects can be minimized with judicious PSF template star selection, and observation scheduling and design allowing the “effective” IWA to approach the physical limit of the occulting mask, i.e., $r = 0.3$ for STIS WedgeA-0.6. Thus, angularly resolving small inner debris rings, centrally cleared regions in CS disks, and constraining their inclinations may also be done with photometric robustness in total light and without the need for advanced processing techniques that do not spatially conserve the target flux. In comparison to techniques such as principal component analysis,

¹⁸ <http://www.gemini.edu/node/12113>

¹⁹ <http://www.eso.org/public/news/eso1417/>

our variant on classical (global) PSF subtraction does not require an existing a priori geometrical model for the debris system. This is likely to be particularly valuable for systems that have only been resolved at longer wavelengths or those whose debris disks are inferred only from IR excesses; e.g., the many newly identified debris systems discovered from WISE observations (Patel et al. 2014), or those that are spatially resolved by *Herschel*, but for which optical imagery does not exist.

Imaging the outer, faint, and diffuse regions of CS disks, where small particles may be escaping the system or interacting with the ISM is also difficult (if not problematic in most cases) from the ground where ADI facilitates observations. ADI data are obtained as suites of exposures where the rotation of the system on the sky enables the high-contrast detection of sources. ADI relies on azimuthal median filtering where the angular size of the source is less than the system rotation in individual images. Exposure times are limited by the need to avoid rotational smearing, which increases in tangential scale at larger stellocentric distances. ADI is very effective for point sources, i.e., exoplanets in close angular proximity to their host stars ($\sim 10\text{--}100$ s AU in physical scale in these disks). However, ADI-assisted detection of CS debris suffers for systems lacking small spatial scale substructures and without locally strong brightness gradients (e.g., face-on disks and those with large, diffuse outer disk halos), and in particular for stellocentric regions beyond about $2''$. These issues do not arise with *HST*. With *HST*, “maximum” exposure times are set by Earth-occultation interruptions, or splitting long integrations for cosmic-ray detection and correction, and are on the order of a kilosecond or longer. The total integration times for six-roll observations are typically >10 ks and thus with sufficient S/N in the outer disks to map faint outer halos, tracing either loosely bound particles, or those which are being actively ejected from the system.

8.2. All Debris Systems are Not the Same

The imagery presented in this paper shows systems with debris rings of different widths, morphological and SB asymmetries, substructures, broader debris disks, and debris systems with blow-out structures that do not imprint conspicuous signatures in their SEDs, in sub-millimeter observations, or in FIR thermal emission.

Detailed dynamical and compositional modeling of these and augmenting multi-wavelength data, beyond the intended scope of this paper, is required to fully exploit the new images obtained in this study. The detection of radial differences in the scattering phase functions in inner debris rings and outer disk halos, with (or without) pericenter offsets of debris belts from their host stars, offer new constraints to such models. Modeling efforts incorporating these data are still in their early phases, but have already yielded indirect signatures, suggesting the presence of wide exoplanets in HD 181327 (Stark et al. 2014) and HD 61005 (D. C. Hines et al. 2014, in preparation). The detection of spiral-structured debris from a posited major collision originating in a well-revealed exo-Kuiper Belt 86 AU from HD 181327 (Stark et al. *ibid*) challenges models for when larger bodies are present in such structures, and highlights debris disks as highly dynamical in nature. These data further suggest that environmental interactions of the outer portions of debris disks with the ISM also appear to be common (J. H. Debes et al. 2014, in preparation).

Posited implications for other debris system features and substructures, revealed by the images we have obtained, are

discussed on a target-by-target basis in the [Appendix](#). Therein we demonstrate the analytic and interpretive potential of such high fidelity optical coronagraphic imaging of CS material uniquely enabled with STIS multi-roll PSFTSC. These images highlight the fact that debris disks are highly dynamic and complex environments. The picture that emerges, even from this small sample of previously imaged debris disks, is one of system-to-system diversity, perhaps matching the diversity of exoplanets and exoplanetary systems known from RV and transit studies. Figure 9 clearly illustrates the diversity in debris system global morphologies and architectures (separate from those expected from simple line-of-sight inclination effects) as well as asymmetries and substructures. We comment on several below.

1. *Ring-like disks.* Ring-like features in CS debris systems are readily, and unambiguously, observable in CS disks with favorable (intermediate to face-on) viewing geometries when they exist, such as in HD 92945 (Section A.6), HD 107146 (Section A.7), and HD 181327 (Section A.9). The improved IWA of our data compared to previous *HST* observations has resulted in imaging rings also in the nearly edge-on systems HD 61005 (Section A.5) and (partially detected in) HD 15115 (Section A.1). Even in our small sample of systems, the ring architectures are diverse and range from broad and low SB in HD 107146, a narrow ring with an outer diffuse and very low-SB component in HD 92945, an extremely bright ring with an extended and highly asymmetric outer component in HD 61005, a narrow and very bright ring seen in HD 181327, and an (apparently) extreme “front/back” brightness asymmetry in HD 15115. These five disks are also extremely diverse in scattering fractions from $F_{\text{disk}}/F_{\text{star}} = 0.25\%$ for HD 61005 to 0.005% for HD 92945, though these two disks in particular have rings of approximately the same physical size ($r \sim 60$ AU) and are of similar estimated ages ($\sim 100\text{--}200$ Myr). Conversely, the broad ring of the sun-like star HD 107146, though with a similar spectral type host star as HD 61005, is ~ 2 times as large. Despite their specific differences, the prevalence of such structures of sufficient brightness to be imaged (seen in other debris disk as well, e.g., HR 4796A and Fomalhaut), suggest they are common features in debris systems.
2. *Highly asymmetric disks.* The HD 61005 debris ring was not resolved in the lower spatial resolution scattered-light discovery imaging in the near-IR with NICMOS (Hines et al. 2007), due to its nearly edge-on inclination (nor with ACS due to its large IWA). Those observations, however, revealed a large “skirt” of likely small-grain material (earning it the name of “the Moth” by its morphology), attributed to material escaping the system posited by ISM ram pressure ejection (Hines et al. 2007; Maness et al. 2009). This suggested that prior modeling and interpretation of debris disk dynamics and evolution, as isolated systems, was deficient. Environmental interactions could be important and must be considered, e.g., chemical evolution of the disk material by ISM “pollution,” compositional evolution and differentiation by stripping of small particles from the disk, and thus mass loss (disk erosion) over time by systemically extrinsic (as well as prior considered intrinsic) forces. The GO 12228 AQ images better revealed the HD 61005 debris system structures, but also three other examples of disks, two at an nearly edge on viewing geometry, HD 15115 and HD 32297 (Section A.3), and the prior described “fan-like”

disk of HD 15745 (Section A.2), where ISM/environmental interactions may be responsible for their morphologies. Low S/N detection of diffuse material also be beyond the outer edges of the HR 4796 A (Thalmann et al. 2011) and Fomalhaut (Kalas et al. 2013) debris rings, suggesting that such structures may be common and potentially important in understanding debris disk evolution and dynamics.

3. *Out-of-plane features/substructures in edge-on disks.* Disks viewed edge-on (or nearly so) while (at least partially) self-obscuring the in-plane inner disk regions, fortuitously provide a most favorable geometry for readily studying “out of plane” CS material without ambiguities. The prototypical β Pictoris edge-on debris system sports a strong “warp” (Heap et al. 2000), or inclined (out-of-plane) secondary disk (Golimowski et al. 2006; D. Apai et al. 2014, in preparation). Its origin is posited as possibly arising from a forcing interaction with its interior, imaged, co-orbiting giant planet, β Pic b and/or yet undetected additional planets in the system. In AU Mic (Section A.10), we detect an out-of-plane “bump” (SB and presumed dust density enhancement) on only one side of the disk at ~ 13 AU, but with an asymmetric warp about the disk minor axis that may be difficult to explain in the absence of one or more co-orbiting planets. HD 15115 is viewed nearly edge on (but not quite so by a few degrees). While no such features are seen in the inner disk, only on one side does the outer disk appear intriguingly bifurcated at a large stellocentric distance whose origin remains (at this time) only speculative.

8.3. Novel Structures and Behaviors in Older Protoplanetary Disks

While small-particle material originating in debris disks can be seen ejected to large stellocentric distances (with some likely escaping from some systems), distant material associated with protoplanetary disks informs on the evolutionary history, and particularly the star-forming environment, of the host-star system. Large, remnant protostellar envelopes have previously been observed at stellar ages of 4 Myr (AB Aur; Grady et al. 1999) and at 10 Myr (HD 100546; Grady et al. 2001), in association with intermediate-mass stars. The scale of these structures suggests that the system formed in isolation and without a history of dynamical perturbations, and also highlights structures that are inconspicuous at longer wavelengths and that do not leave footprints in the disk IR excess. The data for MP Mus (Section A.11) obtained in this survey not only provide a superb image of the disk of this system ready for detailed modeling that can be compared to co-eval systems in the TW Hya association, but also provides evidence for a remnant envelope, indicating that at least some T Tauri stars form in isolation as profound as some of the nearest Herbig Ae stars. Our observation of MP Mus also demonstrates the utility of the observational approach for studying spatially resolved temporal variability in disk-surface illumination.

8.4. Symbiosis with Other Wavelength/Instrumental System Data

Our STIS six-roll PSFTSC images directly reveal disk features and substructures unseen in precursor *HST* discovery data, and components that had been inferred from SED modeling, but not previously observed directly. These visible-light observations provide imagery with spatial resolution comparable to that which ALMA should soon provide, and an order of magnitude better than delivered by the *Herschel Space Observatory*.

Optical (*HST*) and sub-millimeter (ALMA) imaging at similar spatial resolutions are highly symbiotic, with *HST* tracing the small grain structures in the debris systems, and ALMA tracing the population of larger particles (see Boley et al. 2012 and Moór et al. 2013), and E-VLA sensitive to centimeter-sized material.

Extreme A/O systems are now coming online for ground-based 6.5–8 m class telescopes (e.g., SCExAO, MagAO, GPI, SPHERE), and offer higher ground-based Strehl ratio imagery than the previous generation of AO systems, though with more restricted FOVs than offered by *HST*/STIS. Nonetheless, they offer symbiotic capabilities for imaging moderate contrast features in the bright regions of angularly small disks, enhanced when limited to polarized flux, in particular in the NIR where Strehl ratios and imaging performance is improved (and with spatial resolution, e.g., in *K* band, comparable to STIS in the optical and thus providing, in combination, significant color diagnostics).

In addition to the remainder of the approximately dozen *HST*-discovery images of debris disks not followed up in this study, including a few recently found in NICMOS LAPLACE (Schneider et al. 2011) reprocessed data with KLIP (Choquet et al. 2014), there are now known debris disks at <50 pc that are resolved at long wavelengths by the *Herschel Space Observatory*; for example, HD 139664 (reported in this paper with STIS), HD 127821 and HD 13337 all $L_{\text{IR}}/L_{\text{star}} \sim 10^{-4}$ (Su et al. 2013), HD 207129 (Löhne et al. 2012), and Eta Crv (Duchêne et al. 2014). These *Herschel*-resolved debris disks are typically too large in angular extent to image with extreme AO systems (and in some cases too large for complete imaging by the *James Web Space Telescope* (JWST)), but can be fully imaged with *HST*/STIS. Some of these newly discovered, ring-like disks show IR pericenter glow in *Herschel* PACS imagery, but the low (order arcsecond) angular resolution of the *Herschel* data preclude a quantitative measurement of pericenter offsets that are classic signatures of exo-planets on eccentric orbits. Optical coronagraphy with *HST*/STIS, with spatial outer working angle coverage of the full exo-Kuiper Belt regions of such disks, provides an order of magnitude improvement in angular resolution compared to *Herschel* data. This also enables measurements of pericenter offsets <3 AU at $d = 27.5$ pc, and of outer belts inclinations to $\pm \sim 1^\circ$ that are an order of magnitude improved over those achievable from the thermal IR imagery (Greaves et al. 2014). The availability of such inclination constraints, in systems with known exoplanets with orbital data, would result in more robust estimates of the fraction of exoplanets with orbits that are close to coplanar with their host star Kuiper Belts, and the fraction that have been perturbed into higher inclination orbits (Greaves et al. 2014). Moreover, the frequency of pericenter offsets and their quantitative measures allow us to evaluate similar data for eccentric planets (see Beust et al. 2014 for application to Fomalhaut). High-quality inclination measures for systems further allow selection of disks to search for transiting exo-comets (Roberge et al. 2014).

8.5. Conclusion

High-fidelity scattered-light images of CS disks are needed to (1) constrain the modeling of individual systems, (2) provide reference data for the interpretation of observations obtained in thermal emission, or (3) directly reveal material that is inconspicuous at longer wavelengths (Lebreton et al. 2012). In some cases, the discovery images of the disks that were targeted in this study were sufficient to establish the presence of CS material, or to constrain models of the IR excess. However,

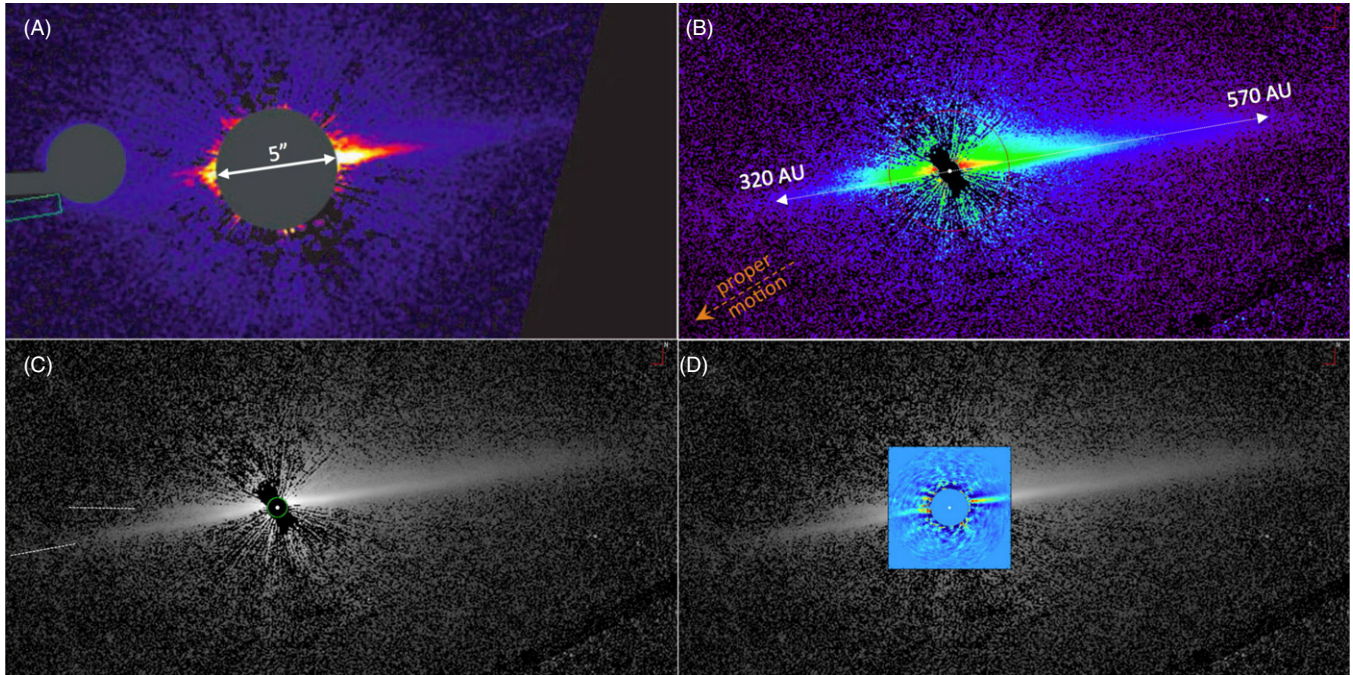


Figure 11. (A) HD 15115 scattered-light ACS discovery image reproduced from Kalas et al. (2007b). (B) STIS 6R/PSFTSC image \log_{10} displayed with 20,000 times imaging the dynamic range from $[-3.0]$ to $[+1.0]$ counts $\text{s}^{-1} \text{pixel}^{-1} \{\text{dex}\}$ ($0.177\text{--}3532 \mu\text{Jy arcsec}^{-2}$). The radial extents of the disk along the mid-plane to sensitivity-limited stellocentric distances at 3σ resel^{-1} over the background are indicated. Note the preponderance of starlight-scattering material preferentially “above” the disk mid-plane (only) to the NW of the star. The region interior to the $r = 2''.5$ red circle was unimaged with ACS. (C) Same data as panel (B) shown in grayscale with a \log_{10} stretch from $[-3.5]$ to $[+1.0]$ counts $\text{s}^{-1} \text{pixel}^{-1} \{\text{dex}\}$ ($0.056\text{--}3532 \mu\text{Jy arcsec}^{-2}$). The green circle indicates the $r = 0''.4$ IWA along the disk mid-plane. The mid-plane diverging, bifurcated, linear light-scattering structures on the east side of the disk are indicated. (D) Disk mid-plane from $\sim 1''.0$ to $2''.5$ (inset FOV $5'' \times 5''$ linear display stretch from -1 to $+2.5$ mJy arcsec^{-2}) resolved by Rodigas et al. (2012) at the K_s band with LBT/PICES overlaid on the STIS image. All images are north up, east left.

(A color version of this figure is available in the online journal.)

the discovery images were typically not suited for dynamical modeling of individual systems, or for identifying indirect signatures of planet presence: such studies require both the image fidelity and S/N provided by our multi-roll observing strategy. Such data enables us, for the first time, to obtain holistic views of exoplanetary systems that have yet to be achieved for our own solar system.

The combined image quality, photometric efficacy, and inner to large outer working angle depth (sensitivity) offered by *HST*/STIS multi-roll PSFTFC for CS disk imaging is unique and currently problematic with any other existing instrumental systems. We argue that as a foundational basis and legacy for both future multi-wavelength (for example, in combination with ALMA), and follow-up in development or conceived high-performance ground, or space-based disk-imaging coronagraphic facilities (e.g., *JWST*, NASA’s *Astrophysically Focused Telescope Assets*, the *Exoplanetary Circumstellar Environments and Disk Explorer* mission concept), new observations of as many disks as feasible should be carried out while *HST* and STIS are still available.

The *HST*/GO 12228 investigation team (co-authors of this paper) dedicate this paper to the memory of our late colleague and co-investigator, Bruce E. Woodgate (1939–2014), whose influence and insights in advancing astronomical instrumentation and astrophysics have touched so many. As Principal Investigator and architect of the Space Telescope Imaging Spectrograph, our investigation and so many others would not have been possible without his vision and foresight. He will be sorely missed by us all.

This study is based on observations made with the NASA/ESA *Hubble Space Telescope*, obtained at the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS 5–26555. These observations are associated with program No. 12228. Support for program No. 12228 was provided by NASA through a grant from the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS 5–26555. J.C. was supported by the Research Corporation for Science Advancement through a Cottrell College Science Award. C.C.S. acknowledges the support of a Carnegie Fellowship and an appointment to the NASA Postdoctoral Program at NASA Goddard Space Flight Center, administered by Oak Ridge Associated Universities through a contract with NASA.

APPENDIX

INDIVIDUAL OBJECTS

A.1. HD 15115

Introductory notes. A starlight-scattering debris disk with an “extreme” asymmetry about its F2 main sequence star, a posited member of the ~ 12 Myr old β Pic moving group, was discovered by Kalas et al. (2007b) from ACS *F606W* coronagraphic imaging data obtained by the GO program 10896 (reproduced in Figure 11(A)). Through contrast-limited by PSF-subtraction residuals at stellocentric angular distances $\leq 2''.5$, the ACS imagery revealed the outer disk’s edge-on geometry and highly asymmetric extent to ~ 315 AU on the east side of the star and ~ 550 AU on the west side. Follow-up *J*, *H*, and *K'* AO

imaging of the disk with the NIRC2 instrument on the Keck-II Telescope confirmed the existence of the disk asymmetry beyond $\sim 2''.5$, but suggested a more symmetric structure from $0''.7$ to $2''.5$. With combined *HST* and Keck photometry in commonly sampled stellocentric regions, Kalas et al. (2007b) suggested the optical-to-NIR color as being blue, with small dust particles postulated in a manner similar to those presumed to be in the AU Mic disk. The morphology and *V-H* color suggested an informal designation of the HD 15115 disk as the “blue needle” by Kalas et al. (2007b).

With subsequent *HST*/NICMOS *F110W* ($1.1\ \mu\text{m}$) coronagraphic imaging, Debes et al. (2008) found, in combination with the prior *V*- and *H*-band photometry, asymmetric radial color gradients in the disk, with the redder dust particles closer to the star, and suggested a “warp” (only) on the western side of the disk. Earlier, Kalas et al. (2007b) posited that the extreme asymmetry might originate from external dynamical perturbations to the debris dust from HIP 12545, a nearby co-moving M-star member also likely a coeval member of the β Pic moving group. Debes et al. (2008) alternatively suggested that a recent collision in the disk might explain the asymmetry, color gradients, and the suggested “warp.”

More recently, Rodigas et al. (2012) obtained diffraction-limited imaging of the inner ($r < 2''.5$) portions of the HD 15115 disk with the PICES instrument on the Large Binocular Telescope (LBT) in the *Ks* band from $1''.0$ to $2''.5$, and LBTI/LMIRcam AO imaging at L' from $1''.2$ to $2''.1$. They found the disk to be more symmetrical at the longer wavelength, and a spectrally neutral *Ks* – L' color in the $\sim 2''$ region on both sides of the disk. In combination with the prior shorter-wavelength-derived colors, they suggest that the scattered light on the west side originates from a dust particle population with smaller grains, suggesting that the disk dynamics are perturbed by a local ISM wind. Rodigas et al. (2012) also suggest a “bow-like” appearance of the inner disk in both the *Ks* and L' bands offset by a few astronomical units to the north of the star.

Observations and PSF subtraction. The STIS 6R/PSFTFC observations of HD 15115 and its PSF template star executed with no anomalies or degradations of significance with reduction and template PSF subtractions. At most celestial P.A.s (inclusive of the direction along the disk major axis), an IWA of $r = 0''.4$ was achieved (green circle in Figure 11(C)). The relative orientation constraints due to *HST* guide star availabilities were a bit more restrictive than typical, so spatial coverage between $0''.4$ and $0''.8$ due to multi-roll combined Wedge-A obscuration in the sectors flanking approximately celestial P.A. 30° (and affecting the visibility of the disk minor axis) is lacking. This is the “hard black” region of “no data” in the center of the image in panels (B) and (C) in this stellocentric radial zone.

Principal results. The STIS 6R/PSFTFC AQ image of the nearly edge-on HD 15115 CS disk has (within a factor of a few) similar photon-limited sensitivity to the dust in the outer portions of the debris system as was revealed by ACS, and is seen to a similar angular extent (see Figure 11(B)). The inner disk, however, revealed at similar optical wavelengths, enabling spatially resolved image interpretation and photometry between $0''.4$ and $2''.5$ (red circle in panel (B)), is not accessible in the ACS image (panel (A)). The primary geometrical and photometric results derived from the AQ image are as follows. The total flux density from the disk at $r \geq 0''.4$ (≥ 18 AU projected on the sky) is ≈ 2.04 mJy. The $0.58\ \mu\text{m}$ disk scattering fraction ($F_{\text{disk}}/F_{\text{star}}$) at $r \geq 0''.4$ is 0.030% (approximately an order of magnitude less than HD 32297; see Section A.3) and HD 61005 (discussed in

context in Section A.5). Light-scattering material in the plane of the disk is seen to at least 320 AU to the east and 570 AU to the west of the star (panel (C)). The celestial orientation of the disk major axis appears to be $\sim 99.1^\circ$ (E of N) as determined from photometric isophote fitting along the apparent mid-plane of the disk at $r > 1''.2$ on both sides of the star.

The asymmetric “needle”-like morphology of the HD 15115 disk, as originally suggested by Kalas et al. (2007b), is also confirmed in the inner regions previously unseen in the optical. However, the suggestion of a more azimuthally symmetric disk, as inferred from the *H*-band NIRC2 image, is not confirmed and instead suggests that the symmetrical structure appearing in the ground-based Keck AO image may be an artifact from imperfect wavefront error control and starlight rejection. In addition to the E/W asymmetry in the disk extent along the major axis as previously noted, three other very significant disk features are seen in the STIS AQ image.

1. The “bowing,” or concavity, of the disk along the minor axis to the north of the star suggested by the Rodigas et al. (2012) *Ks*-band image (reproduced and astrometrically overlaid on the optical image in Figure 11(C)), is seen to continue inward to the STIS optical IWA_{effective} $r = 0''.4$ (18 AU) along the disk major axis. This suggests that we are seeing the brighter (front?) scattering surface of an inner debris ring viewed in nearly edge-on projection that may be at least partially cleared; see the schematic suggestion of a posited inner ring component in Figure 12 (inset). The “invisibility” of the opposing (southern) side suggests a very high Henyey–Greenstein (H-G) scattering asymmetry parameter *or* dust obscuration through an incompletely cleared inner disk component. Detailed modeling will be needed to arbitrate the possibilities. A similar structure, with a more well-defined central clearing and lesser “front/back” scattering asymmetry, was revealed for the nearly edge on debris system of HD 61005 (Section A.5).
2. On the east side of the debris structure, the light scattered from the dust appears to be bifurcated away from the inferred disk major axis direction. This is indicated in Figure 11(C) by the two dashed lines annotating these disk features. The origin of this feature is uncertain, but could be due to a line-of-site opacity enhancement of the dust in that direction. We note more well-defined bifurcated structures associated with the outer disk of HD 61005 as shown in Figure 23(C).
3. There is a significant out-of-plane (vertical) extension in the scattered light distribution at $r \sim 2''.5$ above (to the north of) the disk plane *only* on the west side of the star, i.e., the region in green as stretched in Figure 11(B) to the NW of the star.

All three of these asymmetries must be modeled (and reproduced) prior to ascertaining which of the prior posited asymmetry mechanisms (stellar perturbation, recent episodic collisions, or ISM interaction) may be dominantly causal, though this is beyond the scope this paper.

A.2. HD 15754

Introductory notes. HD 15745 is an ~ 100 Myr old F2V star for which an extended scattered light disk with a “fan”-shaped morphology was reported by Kalas et al. (2007a) based on ACS coronagraphic imaging (reproduced in Figure 13(A)).

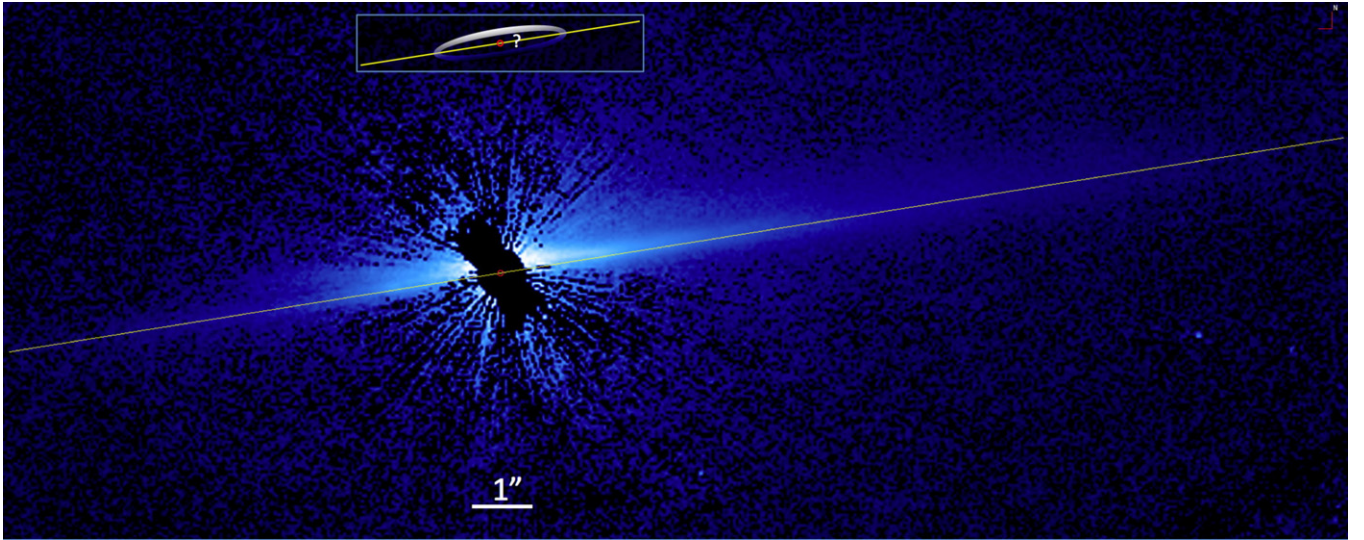


Figure 12. Mid-plane of the asymmetrically extended HD 15115 CS debris disk, found from fitting the locus of points defining the peak SB across the apparent disk major axis at $r > 1''.5$ on both sides of the star, is overplotted on the disk image. The morphology and deviation of the flux density excess at smaller stellocentric distances is to the north of the star only. This qualitatively suggests a nearly edge-on inner debris ring (schematically represented in the image inset), reminiscent of HD 61005 (see Section A.6) that may be only partially cleared in the interior. The western side of the posited ring in this image, “above” the plane of the disk, is more readily seen (compare to the illustration inset). The eastern side of the debris ring may be distorted (or partially obscured) by more complex disk features, e.g., as evidenced by the bifurcated structure highlighted in Figure 11(C), arising from extrinsic forces noted in the main text. This image north up, east left with a \log_{10} display stretch from $[-3.5]$ to $[+1.0]$ counts s^{-1} pixel $^{-1}$ {dex} (0.056 to $3532 \mu\text{Jy arcsec}^{-2}$) and $25''.4 \times 10''.4$ FOV.

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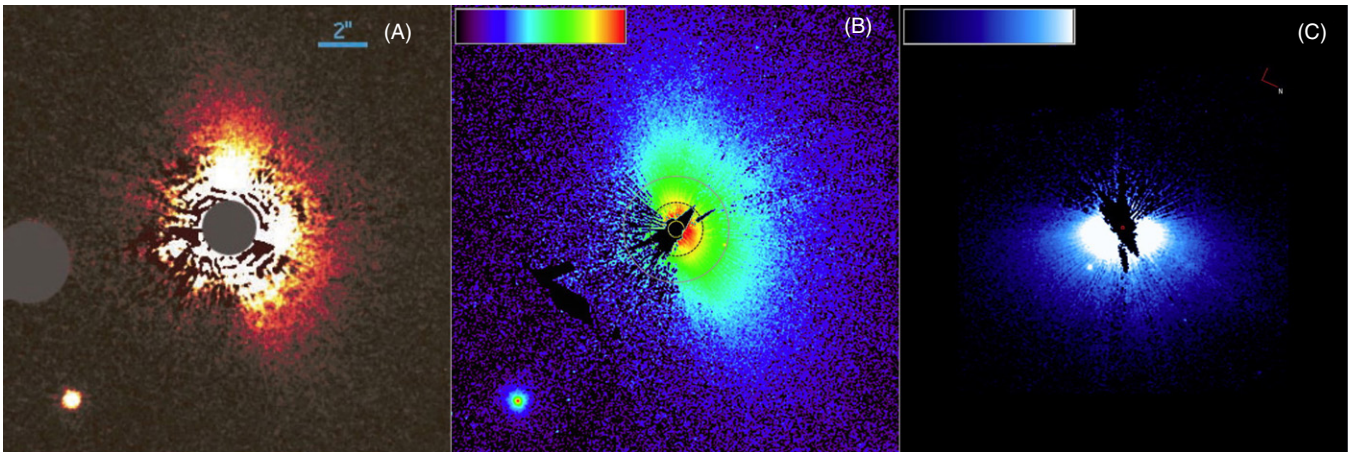


Figure 13. (A) HD 15175 ACS discovery image (from Kalas et al. 2007a; see Figure 1). The central masked region $r = 1''$ but “the presence of quasi-static speckles produces significant PSF-subtraction residuals, limiting the detection radius to ~ 2 arcseconds,” (p. 161). (B) STIS three-roll combined image (visits B5–B8 only) improves over ACS IWA $_{\text{effective}} = 2''$ undegraded from residuals by $\sim 7 \times$ to IWA $_{\text{effective}} = 0''.3$ (19 AU projected; central yellow circle). The dotted black circle corresponds to the $r = 1''$ digital mask imposed on ACS image in panel (A). The solid gray circle corresponds to the larger $r = 2''$ effective IWA in the ACS. \log_{10} display $[-2.2]$ to $[+2.5]$ {dex} cps pixel $^{-1}$. (C) Same STIS data as panel (B), but with the disk major axis rotated to image horizontal and \log_{10} monochrome display $[-1.5]$ to $[+1.5]$ {dex} cps pixel $^{-1}$. Field shown: 350×350 STIS pixels ($17''.75 \times 17''.75$). N.B.: point object (revealed as binary with STIS) is non-common proper motion background star superimposed on the disk, as determined with NICMOS (2004)/STIS (2012) two-epoch data.

(A color version of this figure is available in the online journal.)

Observations and PSF subtraction—tortured history. The first of the two initially planned sets of observations of HD 15745 and its PSF star was executed on 2011 November 5. The three disk-target orbits (V75, V76, and V78) were executed nominally. However, the guide star acquisition for V77—the contemporaneous interleaved PSF template observation—failed and no PSF data were obtained. This full set of observations was then rescheduled as visits B5–B8 with a different set of guide stars on 2012 November 3. The second set of observations (visits 71–74) was executed nominally on 2012 January 2. However, the PSF star was then found to be an a priori unknown close binary with a sufficiently bright companion to render its

unsuitable as a PSF template. Thus, prior to the repetition of the set of observations that originally failed, six orientations of target imaging data had been obtained, but without a suitable PSF template star. The 2012 November 3 observations that previously failed were re-executed nominally, with a PSF template (VB7) that proved to be an excellent match for its flanking contemporaneous visits. An attempt was made to use the VB7 PSF template for both the non-contemporaneous disk imaging data obtained in visits 71–74 and 75–78. In combination with visits B5, B6, and B8, visits 71, 72, and 74 would provide complete roll coverage. Visits 75, 76, and 78 were redundant with the repeated visits B5, B6, and B8, respectively, but could,

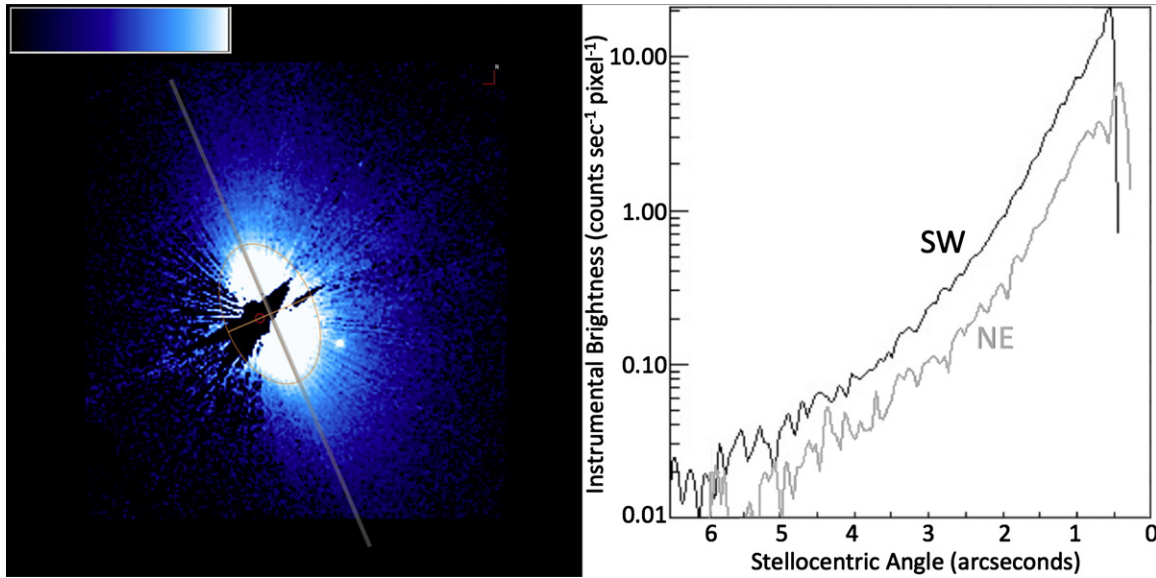


Figure 14. Left: geometrical/isophote elliptical fit to the HD 15745 CS disk, in comparison to the location of the star (red circle) offset along the disk minor axis is indicative of very strong directionally preferential scattering by the disk material. The ellipse fit is to the $0.8 \text{ counts s}^{-1} \text{ pixel}^{-1}$ SB isophote ($1 \text{ counts s}^{-1} \text{ pixel}^{-1} = 0.177 \text{ mJy arcsec}^{-2}$), which is at $r = 34.0$ pixels ($1''.73$) from the disk center along the major axis (with a minor-to-major axial ratio of 0.615). Right: radial SB profiles along the disk major axis are of common slope, but differ throughout in brightness, suggesting a higher dust density distribution on the SW side of the star.

(A color version of this figure is available in the online journal.)

in principle, improve the S/N with replicated observations. Unfortunately (though reductions were attempted), the VB7 PSF was badly mis-matched in the breathing phase relative to the visit 71, 72, and 74 data. Although a scattered-light disk could be seen in these data after its application, the PSF-subtracted images were riddled with artifacts not suitable for combination with the much more successful visit (B5, B6, and B8) PSF-subtracted images. PSF subtractions from the visit 75, 76, and 78 data were somewhat better qualitatively, but still significantly inferior to those obtained at exactly the same roll angles in visits B5, B6, and B8, and could not be used to further improve the data in combination. Other PSF templates obtained in the GO 12228 program were tested as well, but none produced visit 7* PSF-subtracted images of sufficient quality to combine for additional roll coverage with the successful visit B* images. Hence, the fully reduced AQ data images of the HD 15745 disk—unlike the other AQ images in this program—are combined from only *three* differently oriented visits (see Table 3 for details). Nevertheless, the data quality and *spatial* coverage, though not fully complete closest to the star, are both very good and the resulting 3R/PSFTSC combined image shown in Figures 13(A) and (B) was used to ascertain the principal metrical results discussed below.

Principal observational results. The STIS 3R/PSFTSC AQ image of the HD 15745 CS debris disk, to first order, appears have a “featureless,” bi-laterally symmetric morphology about the disk axes (see Figure 14, left). There are no visually apparent substructures evident in the disk, which appears to be smoothly contiguous, radially and azimuthally, with no turnover or “central clearing” of the disk radial SB to a smallest IWA_{effective} of approximately $r = 0''.3$, but with a posited break in the NE side radial SB profile at $r \sim 1''.2$ (see Figure 14) under investigation (J. H. Debes et al. 2014, in preparation). The celestial orientation of the major axis of the disk appears $\sim 22.5^\circ$ (E of N), with a bright-minor to semi-major axial ratio of the disk of ≈ 0.615 as determined from fitting photometric isophotes (e.g., see the illustration of the best-fit $1 \text{ count s}^{-1} \text{ pixel}^{-1}$ SB

contour in Figure 14, left panel). For an intrinsically circularly symmetric disk, this axial ratio implies a disk inclination of $\approx 52^\circ$ from face-on. The estimated center of the disk, also determined from isophotal ellipse fitting, appears to be offset from the position of the star by $0''.258$ (16.4 AU projected) along the disk minor axis toward the SW—indicative of strong directionally preferential scattering by the disk material. This significant offset and front/back scattering phase asymmetry in SB, in combination with longer wavelength data, can be modeled to better constrain the physical properties of the light-scattering debris.

We measured the total flux density from the HD 15745 CS disk to a stellocentric angular distance of $r \leq 10''$, excluding the small unsampled area close to the star (shown in black in Figure 14(A)) with $\text{IWD} \leq 19 \text{ AU}$ as projected on the sky, as $F_{\text{disk}} = 3.26 \text{ mJy}$ with an uncertainty of a few percent. The $0.57 \mu\text{m}$ disk-scattering fraction ($F_{\text{disk}}/F_{\text{star}}$) at $r \geq 0''.3$ is $\approx 0.092\%$. Ninety percent of the dust-scattered starlight is contained within $r < 3''.55$ (255 AU projected) of the star. Diffuse material is seen with a SB $3\sigma \text{ pixel}^{-2}$ brighter than the background noise far from the disk to a distance of $6''.8$ along the disk major axis.

With the geometry of the light-scattering disk established from the AQ image, we measured the disk major axis SB profile on both sides of the star (Figure 14, right panel). Measurements were made in single-pixel increments, in a $0''.1$ (2 pixels $\approx 1 \text{ resel}$) wide photometric aperture, $\pm \approx 6''$ in length with respect to the disk minor axis, as illustrated in Figure 14 (left panel). The SW and NE side SB profiles have identical slopes. At all stellocentric angular distances, however, the SW side of the disk along the minor axis is brighter than the NE side by factors $\approx 3\times$ at $r = 0''.6$ to $\sim 2\times$ at $r = 6''.0$. This global (large-scale) “scalable” brightness asymmetry possibly suggests a higher surface density of scattering grains on the SW side of the disk—perhaps from a recent collision or stirred by planet(s), though other morphological signatures of potentially co-orbiting planets are not apparent.

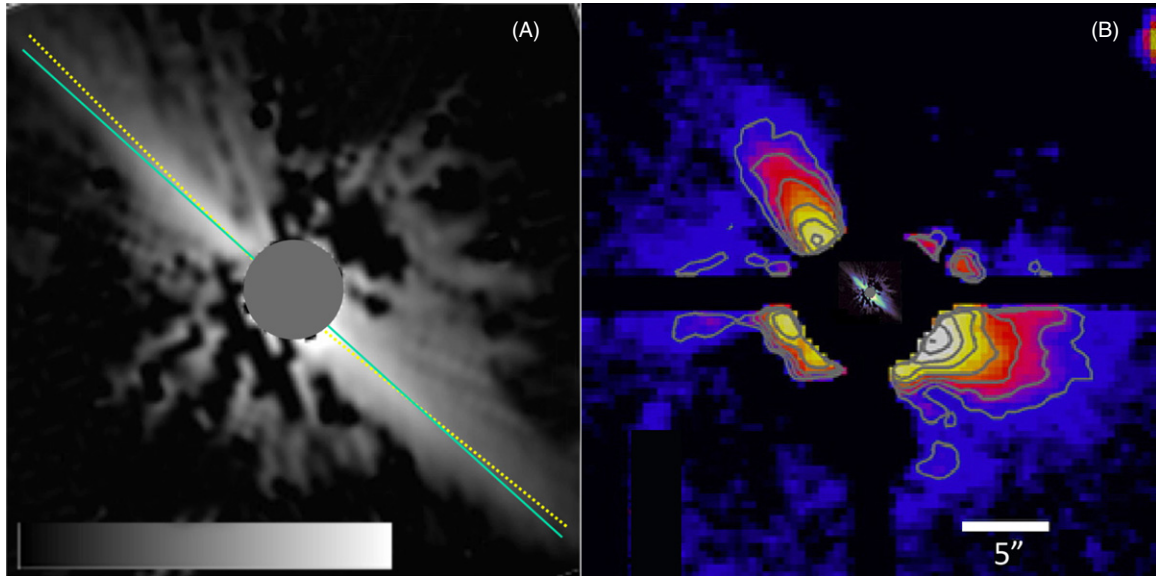


Figure 15. (A) NICMOS image of the $r \leq 3''.3$ region of the HD 32297 disk (replicated from Schneider et al. 2005; the gray circle overlays the $r = 0''.3$ coronagraphic obscuration) separately fitting image isophotes on opposite sides of the star (yellow dotted lines) indicates a departure from a linear fit to both sides of the disk (green line). (B) NICMOS image astrometrically inset with the surrounding large-scale asymmetric nebulosity seen with wider-field, ground-based imaging ($r \geq 5''$, reproduced from Kalas et al. 2005) suggested a then unseen contiguous debris structure undergoing possibly extrinsic perturbations. In both panels, residual diffraction spikes in the NICMOS image, not co-incident with the edge-on disk, have been digitally masked.

(A color version of this figure is available in the online journal.)

The spatially unresolved point-object that appears superimposed upon the disk was undetected in the ACS discovery image due to the presence of strong PSF-subtraction residuals in that region. The object was found to be a non-comoving background star based on differential proper motion measures using epoch 2004 NICMOS imaging (GO 10177) in which the inner most portion of the disk (only) was only marginally detected.

A.3. HD 32297

Introductory notes. The light-scattering debris disk about the 112 pc distant, ~ 10 Myr old A0 main sequence star HD 32297 was discovered from relatively shallow depth, two-roll PSF-subtracted $1.1 \mu\text{m}$ coronagraphy with *HST*/NICMOS (Schneider et al. 2005). The NICMOS image revealed a nearly edge-on disk with a linear morphology extending to a sensitivity-limited stellocentric distance of $r = 3''.3$, and was traced inward to the NICMOS $r = 0''.3$ IWA. The disk was noted as being asymmetrically brighter to the SE along the disk major axis at $r \sim 0''.6$ (50 AU), closely coincident with a peak in the millimeter brightness later reported by Maness et al. (2008). Schneider et al. (2005) noted a minor-axis mirrored asymmetric break in the radial $1.1 \mu\text{m}$ SB profile along the plane of the disk ~ 180 AU from the star, seen only on the NE side of the disk. A subsequent re-examination of the NICMOS image revealed a slight concavity of the disk with the ansae of the major axis slightly deflected to the NW as shown in Figure 15(A). Kalas (2005) obtained a ground-based *R*-band image of a much larger nebular structure enveloping the *HST*-resolved debris disk with the U. Hawaii 2.2 m telescope and $r = 3''.25$ coronagraphic mask, but with a larger IWA of $r \sim 5''$; i.e., the ground and *HST* images explored mutually exclusive domains in stellocentric distance. The asymmetric morphology and SB maps of the wide-field, ground-based, and narrow-angle *HST* images, however, by extrapolation were strongly spatially correlated (Figure 15(B)), and suggested an unseen extended debris structure linking the inner disk seen by NICMOS with the enveloping nebulosity seen from the ground. Kalas (2005) suggested that the swept-back

morphology of the high SB regions in the nebulosity, deviating in celestial P.A. from the linear extension of the *HST*-imaged inner disk, was due to a “sandblasting” interaction of the disk dust particles by the ISM or from perturbations from two nearby stars. The millimeter appearance of this system is discussed in Maness et al. (2008).

Follow-up *HST*/NICMOS imaging by Debes et al. (2009) at $1.6 \mu\text{m}$ (*H* band) and $2.05 \mu\text{m}$, providing color constraints on the disk dust particles, found both color and morphological asymmetries (including warps) that supported the conjecture by Kalas et al. (2005) of the HD 32297 system encountering a dense ISM gas cloud. Subsequent ground-based AO imaging confirmed several aspects of the *HST* observations. K_s -band imaging with the PALAO system and a phase mask coronagraph detected the diametrically opposed “lobes” of the disk, affirmed the $r \sim 50$ AU E/W brightness asymmetry, and provided additional color constraints on the CS grains (Mawet et al. 2009). Later AO imaging with the Keck 2 Telescope using NIRC2 with LOCI processing techniques better informed on the inner disk at the K_s band in the $0''.3 \leq r \leq 2''.5$ explored at shorter wavelengths with *HST*/NICMOS (Currie et al. 2012). Boccaletti et al. (2012) used VLT/NaCo and constrained the disk inclination to 88° , and detected an inner cleared cavity at *H* and K_s . Donaldson et al. (2013) modeled the SED, including new *Herschel Space Observatory* data. They found that the best fit to the stellar properties, including *HST*/STIS UV spectra, is for an A7V star ($T_{\text{eff}} = 7750$ K, $A_V = 0$ with $L = 5.6 L_\odot$). They fit the SED with two dust belts, a ring centered near 110 AU and an inner dust belt at 1.1 AU with an unconstrained outer edge. Their best-fit model for the outer disk was for a minimum particle size of $2.1 \mu\text{m}$, and a 1:2:3 ratio of silicates, carbonaceous material, and water ice with 90% porosity. Rodigas et al. (2014a), however, argue that this particle mixture may not be consistent with the L' imagery of the disk.

Observations and PSF subtraction—details. Due to a proposal implementation error, the contemporaneous PSF template star scheduled with our HD 32297 orbits was $B - V$ redder,

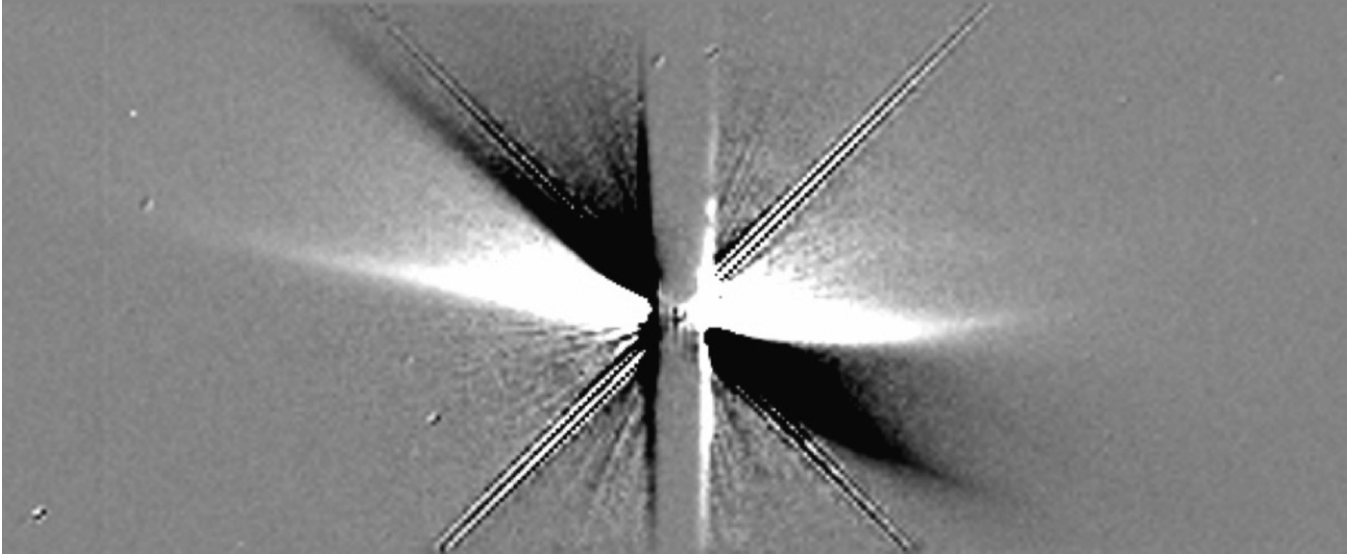


Figure 16. V05 – V06 difference image of the HD 32297 CS disk, differentially rolled by the planned, nominal, 30° with the star occulted at the $0''.6$ wide WedgeA occulter position in both orientations. In the regions where the positive and negative imprints of the disk flux do not overlap, this and similar other paired images inform us that the sensible disk flux originates only to the NW side of the star and within no more than $0''.2$ (narrower than the half-width of the WedgeA-0.6 occulter) to the SW of the plane of the edge-on disk.

$\Delta(B - V) = +0.19$, than tolerable for unbiased PSF subtraction. Fortuitously, off-rolled PSF self-subtraction, using HD 32297 observed at different celestial roll angles as its own PSF template, revealed that the out-of-plane scattered light excess from the disk itself was asymmetrical originating with significance from nearly only one “half” of the disk, to the NW of the disk major axis (see Figure 16). While the flux from the disk in such a self-subtraction may be partially entangled and not fully photometrically preserved, this suggested an alternate strategy not unlike that adopted for AU Mic, which proved successful.

As anticipated with a PSF template with $|\Delta(B - V)| > 0.1$, PSF-template subtraction resulted in a very strong chromatic artifact resulting in significant oversubtraction of the disk flux close to the star. This is illustrated in the six panels of PSF-subtracted images in Figure 17 (row (A)) of the HD 32297 disk from the six visits executed using WedgeA-0.6. (Very similar results, in terms of a repeatable chromatic residual, arises also from the WedgeA-1.0 PSF-subtracted imaging). Note that as the asymmetric disk rotates in the frame of the detector, the chromatic undersubtraction artifact (across row (A)) exposed at changing celestial P.A.’s remains invariant.

We thus created digital masks for each of the PSF-subtracted images (Figure 17 row (B)) to obscure the disk itself in each frame and enable a “build up,” in multi-image combination, of the chromatic oversubtraction signal. A median combination (with missing data exclusion) of these six images, which individually do not fully sample the space around the star, results in a nearly complete image of the oversubtraction component, as shown in panel (C1). While there is a rotationally invariant azimuthal substructure at high cyclical frequencies, at low spatial frequencies the oversubtraction pattern is globally symmetric radially about the star. Hence, we create a radial median image as shown in panel (C2), to “fill in” by approximation (without additional higher spatial frequency information) the “missing data” regions (panel (C3)). We then produce a composite C1 + C3 image as a *correction template* to apply to each of the images in row (A) to remove (most of) the chromatic oversubtractions. The PSF-subtracted disk images, corrected in this fashion, are shown in row (D). We validated the efficacy

of this approach by separately performing PSF-subtractions with a non-contemporaneously observed PSF calibration star, α Pictoris, whose $B - V$ color is a near (but not exactly) perfect match to HD 32297. Comparing the α Pic-subtracted images in row (E) to the chromatically corrected images using our misbegotten PSF star in row (D), and in contrast to the uncorrected images in row (A), clearly shows the efficacy of this approach. Note that we do not simply (alternatively) use α Pic as a PSF-template to build a multiple-roll AQ image, as in those subtractions (row (E)) breathing residuals dominate (i.e., the radial streaks in those images) that are very largely mitigated with the use of the (color mismatched, but corrected) PSF template star. The same is seen with all the WedgeA-1.0 imaging, and we proceed identically with both sets of images to produce the full WedgeA-0.6 and WedgeA-1.0 combined AQ data image.

Principal results. The STIS 6R/PSFTSC AQ image of the HD 32297 CS debris system, utilizing all data obtained from both wedge occulter positions, is illustrated over its full spatial extent in Figure 18. The STIS image indeed reveals a contiguous debris structure “connecting” the inner disk seen with *HST*/NICMOS and ground-based AO imaging, to the outer nebulosity observable with conventional ground-based coronagraphy, with unprecedented clarity.

The STIS image shows a large, highly asymmetric debris structure with complex substructures not previously revealed. Dust-scattered starlight is seen to asymmetrically extend to a distance of ~ 1560 AU from the star with 3σ resel^{-1} detection significance with respect to the instrumental background noise at the extrema of the “swept back” extensions of the disk major axis. This is physically the largest debris system in the GO 12228 10 debris-disk sample. HD 32297 is also the “brightest” of the GO 12228 debris disks in terms of its total starlight-scattering fraction with $F_{\text{disk}}/F_{\text{star}} = 0.3\%$, and is intrinsically bright with a total $0.57 \mu\text{m}$ flux density of 6 mJy. The concentration of most of this debris-scattered starlight into a narrowly viewed edge-on disk within a few arcseconds of the star (see Figure 18 logarithmic SB contour image) makes the inner regions of the disk along its mid-plane a comparatively “low contrast” target accessible to ground-based coronagraphy. The most intriguing

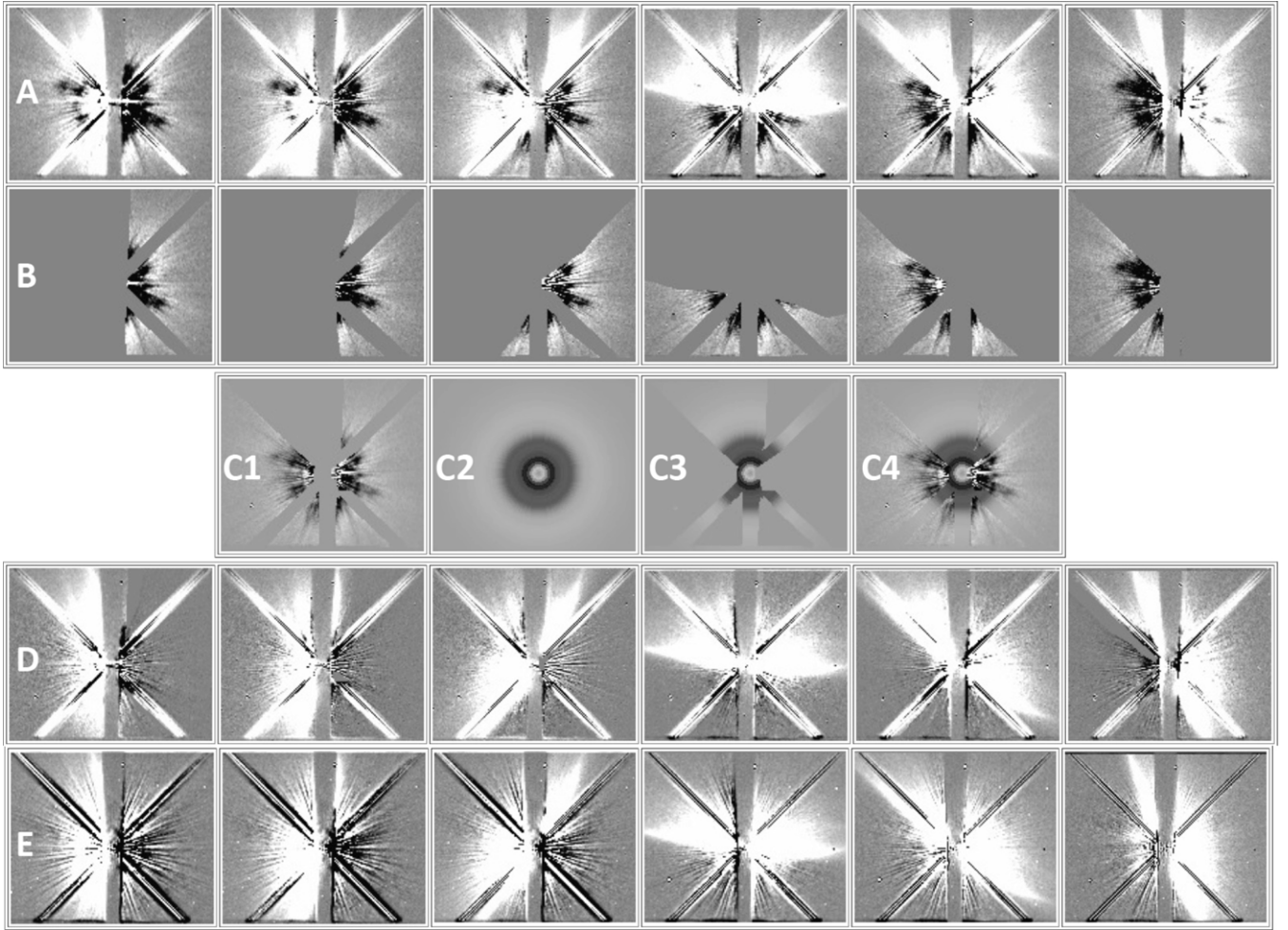


Figure 17. (A) PSF-subtracted WedgeA-0.6 images of the HD 32297 CS disk using unintentionally color-mismatched PSF star; left to right: V05, V06, V08, V01, V02, V04. (B) Masking out the disk signal from the corresponding images. (C1) Masked-median combination of all six images in row (B). (C2) Radial profile image from image C1 excluding unsampled regions. (C3) Using panel (C2) as a proxy to estimate the oversubtraction in unsampled regions. (C4) The final chromatic correction template to apply to the images in row (A). (D) PSF-subtracted images of the HD 32297 disk after chromatic correction (row A images minus C4). (E) Comparison to PSF-subtractions using a non-contemporaneous calibration PSF star, a Pic, with a closely color-matched $\Delta(B - V) = -0.02$.

and more technically challenging features of the disk over a very large imaging dynamic (contrast) range are revealed (only) in the STIS image. To facilitate discussion of the debris system features, we present the AQ image with the disk rotated to the image display horizontal in Figure 19 with three different dynamic display stretches to fully illustrate the nature of the newly revealed debris system substructures.

1. As previously noted, nearly all of the scattered-light flux from the disk originates from only one side of the star to the NW, or “above” the plane of the disk, as presented in Figure 18. The large envelope of light-scattering debris above the disk plane has a complex and asymmetrical morphology, as illustrated by the three different stretches used in panels (A)–(C), to highlight the dust at different stellocentric distances and scale heights above the disk mid-plane. As suggested for HD 61005 by Hines et al. (2007), and for HD 32297 by Kalas (2005), this, too, could possibly be due to small particles in the HD 32297 being “ejected” from the debris system as it moves through and interacts with the local ISM.
2. The most prominent (and brightest) part of the HD 32297 debris system are the two diametrically opposed, nearly linear extensions of the main disk seen in edge-on projection,

partially resolved from the ground, and whose relatively narrow width is best seen in Figure 18 (left). The radial extent of these near-linear features are highly asymmetric, reaching $\sim 6''.0$ to the SE of the star, but only $\sim 2''.8$ to the SW of the star. While Keck AO imaging with comparable spatial resolution suggested diametric co-linearity and equal radial extent of these features in SB and S/N maps obtained with LOCI processing (Currie et al. 2012; see Figure 1), this appears not to be the case as informed by the STIS images.

3. At larger stellocentric distances, these features “curve up” significantly at their extrema (best seen in panel (B)) and bifurcate into very low-contrast “V”-shaped (or “wish-bone”) features with morphological similarity to those seen more readily (with greater contrast) in HD 61005, and on just one side of HD 15115. The morphological similarity of these features modulo the perspective and line-of-site dust opacity effects perhaps suggest a common causality.
4. A long “strand” of higher SB scattering material appears “above” the plane of the disk to the NW of the star (only) roughly “parallel” to the plane of the disk, but does not originate at the location of the star. This feature itself appears to have a concavity toward the disk plane,

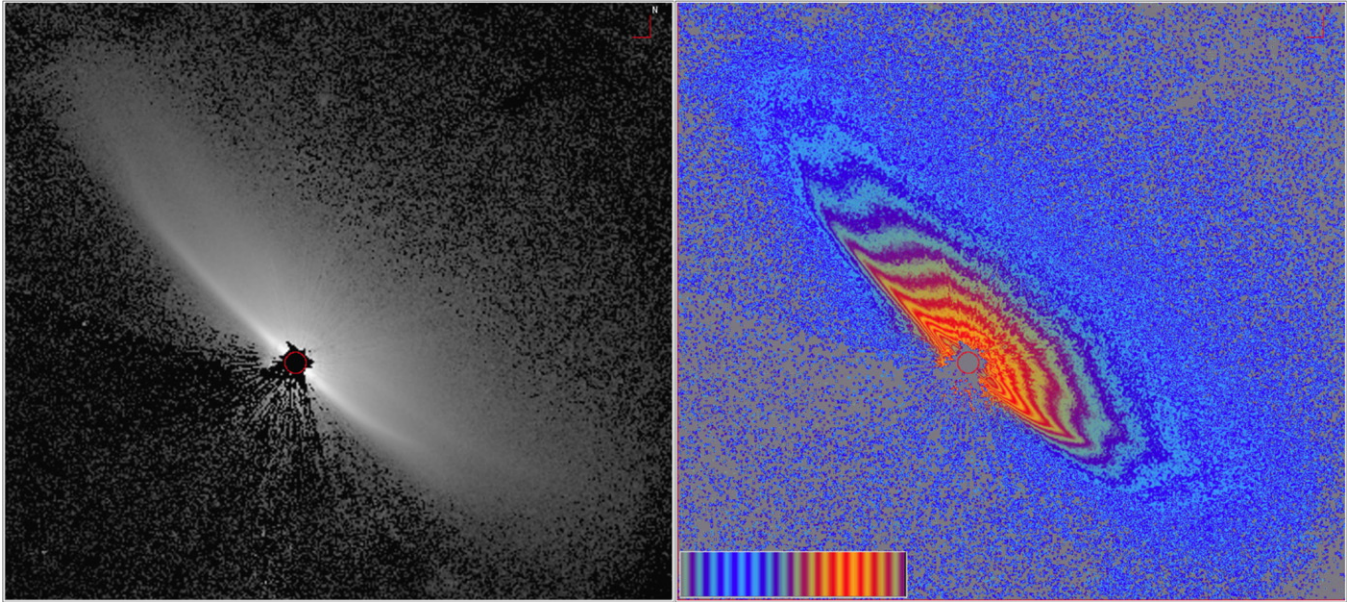


Figure 18. HD 32297 six-roll combined PSFTSC imaging (north up, east left) from all WedgeA-0.6 and WedgeA-1.0 observations. The grayscale image (left) and isophotal contours (right) both \log_{10} display from $[-3.5]$ to $[+2.15]$ dex counts $\text{s}^{-1} \text{pixel}^{-1}$ ($\approx 0\text{--}25 \text{ mJy arcsec}^{-2}$). Field: 380×340 STIS pixels ($19''.3 \times 17''.2$), north up, east left. The red circle indicates a disk major axis IWA of $r = 0''.3$ (6 STIS pixels = width of WedgeA-0.6). 1σ pixel-to-pixel noise assessed in regions far from disk is approximately $0.0032 \text{ counts s}^{-1} \text{pixel}^{-1}$ ($0.56 \mu\text{Jy arcsec}^{-2}$).

(A color version of this figure is available in the online journal.)

and extends nearly twice as far from the star than the main disk on the same disk of the star. The origin of this feature is uncertain and may represent a “Pleiades-like” foreground nebulosity rather than a dust enhancement physically associated with the HD 32297 debris disk explicitly.

The causal mechanism(s) for the extreme asymmetries and complex substructures seen in the HD 32297 disk pose a challenge to current models of CS debris disks, and may require a combination of both intrinsic and extrinsic forces to replicate and explain the observed structures.

A.4. HD 53154

Introductory notes. HD 53154, a G9V solar analog at 18.3 pc with an estimated age of ~ 1 Gyr, is the oldest star in the GO 12228 survey and possibly the oldest star (other than the Sun) for which light-scattering CS debris has been imaged. A low-SB light-scattering debris disk was discovered by Kalas et al. (2006), seen to a contrast-limited IWA of $r \geq 3''$ (55 AU projected), and to a photon-limited stellocentric angular distance of $\leq \sim 6''$ with *HST*/ASC F606W ($0.6 \mu\text{m}$) coronagraphic imaging. These authors inferred an intermediate inclination viewing geometry of the disk of $\sim 45^\circ$ and major axis P.A. = $147^\circ \pm 2^\circ$ based on the ACS discovery image reproduced herein as originally presented in Figure 20(A). With their determination of “no SB asymmetries detected between (the two) sides of the disk,” they then measured the disk’s $3'' \leq r \leq 6''$ minor-axis-mirrored SB profile along its mid-plane and found a smooth decline in SB from 22 to 24 mag arcsec^{-2} , with a shallow dip between $\sim >3''.5$ and $<5''$ (their Figure 3).

Observations and PSF subtraction. The first set of four visits (81–84) of HD 53154 and its (originally chosen) PSF star was executed on 2011 January 10, but with a target acquisition failure preceding the planned coronagraphic imaging in visit 81. The telescope was mispointed and no useful data were obtained for

that visit. A second problem plagued this first set of visits in that the wrong PSF star (too red) was associated and executed with the disk imaging visits. While the contemporaneous color-mismatched PSF template/calibration data from V83 were obtained, they were not useful due to chromatic residuals strongly polluting the resulting disk images.

The second set of visits (85–89) were executed nominally with a replacement PSF star, HD 59780, in V87 on 2011 April 11. Serendipitously, this replacement proved to be a mostly suitable calibration template also for the earlier non-contemporaneous disk-imaging visits. However, because of both “breathing” and differential target centering errors from wedge-deployment non-repeatability, the WedgeA-0.6 (only) attempted PSF subtractions from V82 and V84 (but not V81) were significantly degraded and not of any utility. The WedgeA-1.0 data from these same visits, after non-contemporaneous PSF subtraction, were only mildly degraded (with respect to the very high quality of the epoch 2 PSF-subtracted imaging using its contemporaneously observed PSF template), but are still quite useful, and in combination better tile the roll space, and overall improve the full data set.

The AQ image for HD 53143 (Figures 20(B) and (C)) thus combines WedgeA-1.0 images from all five-roll angles obtained (though mildly degraded, but fully acceptable and still very useful, incorporating visits 82 and 84), and WedgeA-0.6 PSF-subtracted images from three-roll angles (visits 85, 86, and 88) only. Because of the lack of full Wedge-0.6 roll range planned, the spatial coverage close to the star ($r \lesssim 1''.4$) is not complete—i.e., the black “butterfly” region unsampled around the IWA_{MINIMUM} $r = 0''.3$ as illustrated in Figures 20(B) and (C).

Principal results. The STIS (5/3)-roll PSFTSC imaging of the HD 53143 CS debris disk importantly improves over the ACS discovery imaging in two key aspects. First, by better unveiling an asymmetrical structure in the architecture of the outer (>55 AU) disk SE versus NW sectors (see Figure 20(B)), and second, by revealing a previously unimaged inner, azimuthally

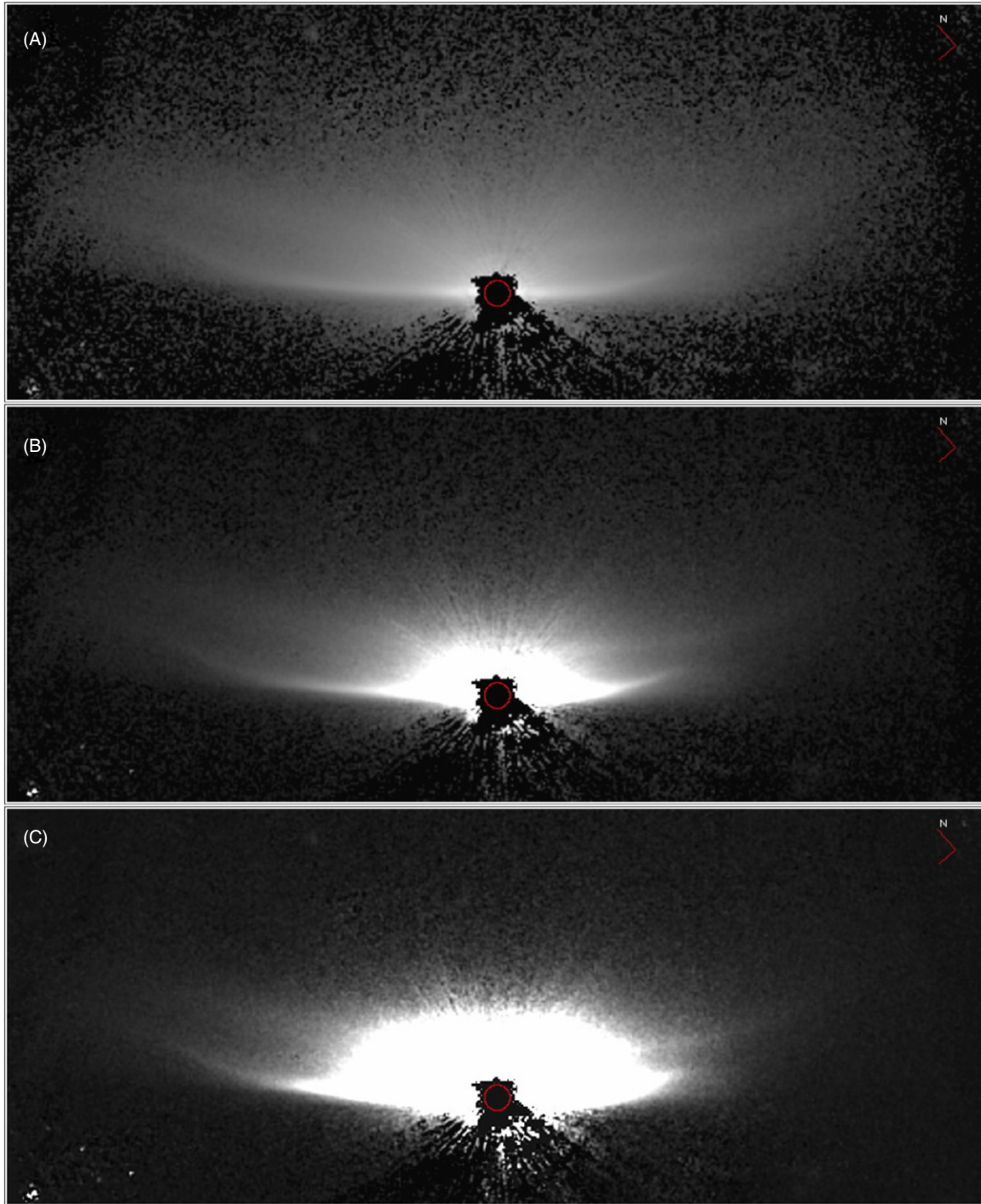


Figure 19. HD 32297 6R/PSFTSC imaging (north up, east left) from all WedgeA-0.6 and -A1.0 observations rotated 41.5° eastward from north to place the inner disk plane on the image horizontal. The three display stretches illustrate the complexity and morphology of light-scattering debris structures over a dynamic range of $13 \text{ mag arcsec}^{-2}$. Top: \log_{10} display from $[-3]$ to $[+2.13]$ dex counts $\text{s}^{-1} \text{ pixel}^{-1}$ (approximately 0.002 to $24 \text{ mJy arcsec}^{-2}$). Middle: square root display from 0 to $1 \text{ counts s}^{-1} \text{ pixel}^{-1}$ (0 – $175 \mu\text{Jy arcsec}^{-2}$). Bottom: linear from -0.01 to $+0.2 \text{ counts s}^{-1} \text{ pixel}^{-1}$ (approximately -1.8 to $35 \mu\text{Jy arcsec}^{-2}$). Field: 460×186 STIS pixels ($23''.2 \times 9''.4$). The red circle $r = 0''.3$ (6 STIS pixels = width of WedgeA-0.6).

(A color version of this figure is available in the online journal.)

symmetric, inner disk component from $r \sim 0''.3$ to $3''.0$ (~ 5.5 AU to 55 AU; see Figure 20(C)).

1. The outer disk: the ACS and STIS images of the commonly sampled (outer) regions of the HD 53143 CS disk are illustrated in Figures 20 (A) and (B) presented at the same spatial scale and celestial orientation (north up, east left). The STIS image is shown with a \log_{10} stretch to

permit the simultaneous display of both the outer and inner disk components, and to match as closely as possible the dynamic display range of the ACS discovery image as previously published by Kalas et al. (2006). In the STIS image, despite the presence of some low-SB residual light from the extended wings of the not completely suppressed stellar PSF halo, a scattered light excess in the outer disk peaking at $\sim r = 4''.7$ to the NW and SW of the star is seen,

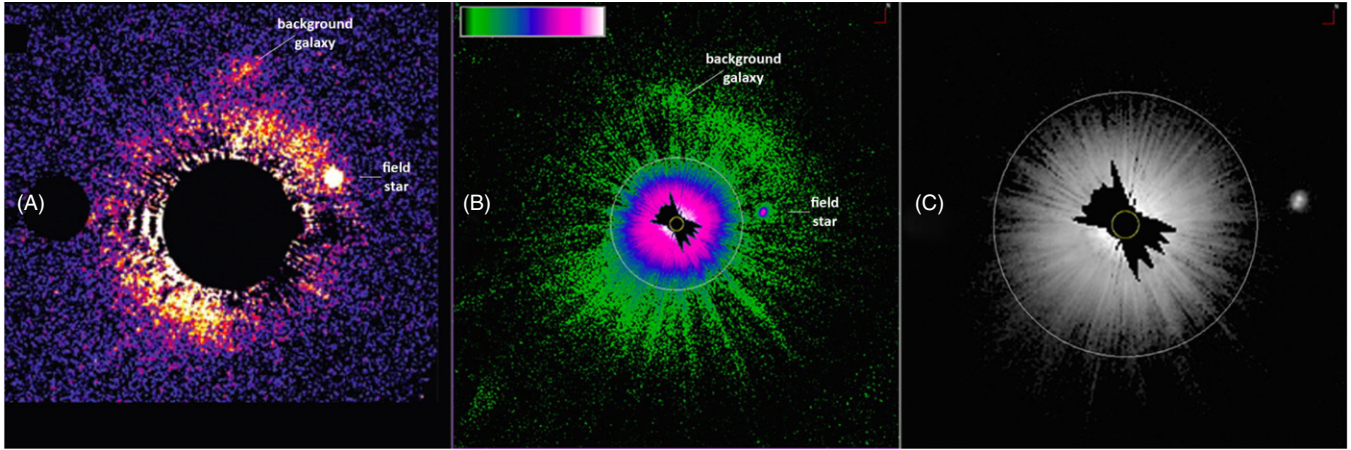


Figure 20. (A) ACS discovery image (Kalas et al. 2006, cf. Figure 1) with $IWA_{\text{effective}} = 3''.0$ digital mask. (B) STIS five-roll combined PSFTSC image. Same image scale as panel (A). Improves $IWA_{\text{effective}}$ by $\times 10$ to $IWA_{\text{minimum}} = 0''.3$ (5.4 AU projected; small yellow circle). 400×400 STIS pixels, \log_{10} display $[-2.2]$ to $[+1.5]$ {dex} counts $\text{s}^{-1} \text{pixel}^{-1}$. (C) “Inner” disk undetectable with ACS observed with STIS, at two times the spatial scale with respect to panel (B) (same data); 200×20 pixels ($11''.54 \times 11''.54$), \log_{10} display $[-1.5]$ to $[+1.5]$ {dex} counts $\text{s}^{-1} \text{pixel}^{-1}$. All images north up, east left.

(A color version of this figure is available in the online journal.)

as is also the case in the ACS image. In the STIS image, a background galaxy whose presence in a single epoch image alone would conflate the morphology of the scattered light, is unambiguously identified as a polluter (along with that of a STIS-resolved binary star superimposed within the disk) contributing spatially coincident flux that does not originate in the disk itself.

The CS scattered light in the outer disk is seen with significance in the two sectors flanking the outer disk major axis on both side of the star to a stellocentric distance $\lesssim 5''.8$ (105 AU). In the STIS image, the SB excess to the SE of the star appears to be spatially contiguous to decreasing radial distances into and interior to the ACS $r = 3''$ $IWA_{\text{effective}}$. In the diametrically opposed NW sector, however, a dearth of disk light between about $3''.5$ – $4''.4$ (63–80 AU) is apparent, which is not seen at the same distance to the SE of the star. Careful, visual examination of the Kalas et al. (2006) image may also implicate the existence of this previously unnoticed minor-axis centered asymmetry. In Kalas et al. (2006), the major-axis radial profile presented was “mirror averaged” about the disk minor axis. With that, a shallow dip in the profile at the radius of the dearth in SE disk light in the STIS image may be seen (but diluted if due to the NW outer disk alone by mirror averaging). Additionally, beyond the digital mask in the ACS image, to the E/SE of the star, there is some evidence of pollution by PSF diffraction rings that may conflate such a determination, though a relative flux deficit seems present to the NW considered by itself.

The diffuse, elongated, light-scattering structure to the NW of the star in the STIS image suggests a morphologically description as segment of an arc. This arc seems to be *marginally* at a larger stellocentric distance by a few astronomical units in the 2011 epoch STIS image when compared to the 2006 epoch ACS image, though it is not well determined or established quantitatively due to its diffuseness (though visually apparent when “blinking” co-registered images). Whether this represents a spatial re-distribution (outward migration) of the scattering dust particles in this region of the disk over this time interval, or a difference due to color-dependent scattering efficiency given the much

wider STIS spectral bandpass, or simply from measurement uncertainty, is highly speculative—but noted here in expectation of future (better informing) observations.

2. The inner disk: the STIS AQ image reveals a significantly brighter inner disk component at $r \lesssim 2''.8$ (blue/pink encoded in Figure 20(B)) that (unlike the outer disk component) appears to be circularly symmetric or nearly so (it is *marginally* brighter at equal stellocentric radii on the SE side of the star than on the NW side of the star). Ninety percent of the dust-scattered starlight is contained within $2''.03$ of the central star. An azimuthally medianed radial surface profile of the inner and immediately flanking regions of the disk from $0''.4 \leq r \leq 3''.6$ is very well fit by a single power law of index -3.7 with no significant deviations (goodness of fit $R = 0.991$).
3. The full disk: the total flux density from the disk, measured to a radius of $5''.8$ excluding the small region unsampled around the $IWA_{\text{effective}} = 0''.3$ is $F_{\text{disk}} = 6.92 \text{ mJy} \pm$ a few percent with then a $0.6 \mu\text{m}$ disk-scattering fraction, $F_{\text{disk}}/F_{\text{star}} = 0.104\%$.

Contrary to earlier speculation based on the ACS discovery image alone, these new STIS observations may implicate a nearly face-on inner disk, with an unbroken $r^{-3.7}$ power law in its SB to $\sim 3''.6$, with an anisotropic (asymmetric) dust density distribution in the outer disk, giving rise to the localized “arc”-like scattering structure seen to the NW of the star at $r \sim 5''$.

A.5. HD 61005 (“The Moth”)

Introductory notes. HD 61005 is a nearby ($d = 27.5 \text{ pc}$), 40 Myr old (Desidera et al. 2011) solar analog (spectral class G2V) with a strong thermal IR excess ($L_{\text{IR}}/L_{\text{star}} = 0.27\%$) identified as a high-priority target from the *Spitzer*/FEPS²⁰ program for *HST* follow-up coronagraphic imaging in *HST* GO 10527. From those NICMOS $1.1 \mu\text{m}$ data, Hines et al. (2007) discovered a light-scattering debris system associated with HD 61005 of unusual character. The morphology of the

²⁰ FEPS: “Formation and Evolution of Planetary Systems”—M. Meyer, PI. This was the first of two light-scattering debris disks discovered with *HST* follow-up from the *Spitzer*/FEPS program. The other was MP Mus; see Section A.11.

debris system was *highly* asymmetric (more so than any other previously seen) with nearly all of the light-scattering material confined to only one side of the star, suggestively named in the discovery literature as “the Moth.” This large scattered-light structure had what appeared to be a leading-edge “bow shock” in the direction of the stellar proper motion. The general paradigm suggested for its (then) unique morphology was causal not as a system in isolation, but posited as one interacting with an ISM wind—as we have subsequently seen with some similarities for HD 32297 (Section A.3), HD 15115 (Section A.1), and perhaps HD 15745 (Section A.2).

The morphology of the large debris structure, seen in the wake of HD 61005’s proper motion, was confirmed with higher spatial resolution ACS coronagraphic imaging by Maness et al. (2009), but left the inner region of the debris itself at $r \lesssim 1''$ unimaged. Those observations, including optical polarimetry at $r > 1''.4$, suggested from the disk colors and high degree of polarization in the disk plane, large (micron or larger size) particles in the disk plane that may have been differentiated from the larger diffuse debris structure, with the disk undergoing erosion by the ISM. Ground-based AO imaging using aggressive PSF-subtraction methods (Buenzli et al. 2010) revealed the low(er) contrast features associated with the disk itself close to the star, and suggesting a near-edge on ring-like disk (with only one side seen) and diffuse material streaming from the posited disk ansae.

Observations and PSF subtraction. *HST*/STIS 6R/PSFTSC observations of HD 61005 and its PSF star HD 56161 executed per the observing plan at two epochs: 2011 February 22 (V25–V28) and 2011 May 3 (V21–V24). We followed exactly the procedure described for reduction, calibration, and the creation of a six-roll combined image using all data from the two occulting wedge positions as described in Section 5.7. The six-roll combined image of the HD 61005 debris system (exquisitely) reveals the edge-on disk, its centrally cleared region interior to a highly inclined bright ring, as well as the conjectured “blowback” material from its interaction with the ISM from the space motion of the system. However, a low-level stellocentric circularly symmetric residual halo of light (in this case bright, rather than dark, as noted for HD 32297) was seen in the north-northwest centered sector of the image “above” the star (where the asymmetric debris structure does not appear), closely reminiscent (but to a lesser degree) of the chromatic PSF-subtraction residuals in HD 32297 prior to mitigation. This superimposed halo is an image artifact due to (small) chromatic differences in the SED of HD 61005 (which itself may be slightly dust reddened) and the PSF template star HD 56161. The halo is not reproducible through intentional PSF template mis-scaling in intensity; attempts to “null it out” by adjusting the scaling produces VERY deep (non-physical) negative residuals close to the star and along the diffraction spikes. We therefore applied the same method of chromatic correction to mitigate these residuals as we adopted for HD 32297 to produce the AQ image discussed below.

Principal results. The AQ image of the HD 61005 debris system reveals with clarity a highly inclined debris ring and its (dark) central clearing to a stellocentric IWD approaching ≈ 10 AU; see Figure 21 presented in two panels to better illustrate (A) the debris ring itself, and (B) the inner part of the skirt of material “blowing” off the ring and of its front-edge bow enhanced in visibility at the extrema from the dust opacity along the line of sight. This “skirt” of (likely small) particles emanating from the previously unresolved inner part of this complex debris

structure is likely pushed back from the disk plane as the system plows through the ISM.

As shown in Figures 22(A) and (B) (reproduced from Buenzli et al 2010), the brighter material in the likely preferentially forward-scattering half of the HD 61005 debris ring only, and ansal “streamers” to a distance of $\sim 1''.5$ beyond the ring ansae, had been detected in ground-based imaging with VLT/NACO using LOCI and ADI techniques. The more diffuse, lower spatial frequency material in the “blow-out skirt” uniquely revealed with STIS 6R/PSFTSC (panel (C) at the same spatial scale and orientation) was seen prior with lesser efficacy also with ACS and NICMOS imaging, but was undetected from VLT with LOCI and ADI processing that do not conserve the flux density in such features. The structures identified in the ground-based images as processed, however, may be directly compared to the *HST*/STIS image. To do so on a closer footing, the STIS image was high-pass filtered with a < 5 pixels ($0''.25$) spatial frequency pass-through that rejected low-spatial-frequency information in a manner that is similar (by design) to the consequences of LOCI and ADI processing, as shown in panel (D). The ability/utility of ground-based ADI/LOCI techniques to identify such (comparatively) lower-contrast features in CS disks (with the *HST* imaging as a “truth test”) is demonstrated with these observations, even though low-spatial-frequency diffuse structures (as in panel (C)) are undetectable in the ground-based images. The complexity of the starlight-scattering debris structures in the HD 61005 CS disk is manifested over a very large range of SBs. To better illustrate, in Figure 23 we present images of the HD 61005 debris system in incrementally increasing 10 times per step shown logarithmically, dynamic display ranges.

In Figure 23(A), we show the debris ring and the “wings of the moth” to roughly the same detection sensitivity as provided by the NICMOS discovery and ACS polarimetry images. Overlaid on this figure, we present a (non-physical) cartoon “model” of the debris ring to scale suggestive of what a complete model (beyond the scope of this paper) might produce, as follows. Using the debris ring ansa to define the disk major axis, we found the celestial orientation of the major axis P.A. as $\sim 70^\circ.3$. Assuming the ring ellipticity results simply from its projection onto the sky (i.e., is intrinsically circularly symmetric) by ellipse fitting we estimate the minor:major ellipse axial ratio as 0.0882:1, from which an inclination of $\sim 4^\circ.9$ from edge-on is derived. We measured the stellocentric angular distance to the center of the debris ring along the disk major axis (i.e., to the ansae, the radius of the debris ring) on average in both directions from the star as $r = 1''.57 \pm 0''.07$. We then estimate the radius of the inner edge of the debris ring (in convolution with the STIS PSF) as $\sim 1''.30$ along the disk major axis, and assume radial symmetry outward. Finally, we estimate a “front-to-back” scattering efficiency ratio for the dust along the disk minor axis of $\sim 3:1$. Physical modeling must be done to confirm/optimize these geometrically and photometrically derived parameters, but as is provides observational constraints for image interpretation. Separately, in panel (C), we call to attention a (first order) axially symmetric scattering feature originating at, and extending beyond, the ring ansae with a “Y”-shaped morphological bifurcation of the scattered light flux on the leading edge of the posited bow shock and (more diffusely) along the disk major axis. Morphologically similar features are seen in the (likely) ISM interacting disk of HD 32297 (Figure 19), and may be related to the bifurcated structure seen only on the east side of the HD 15115 disk (Figure 11(C)).

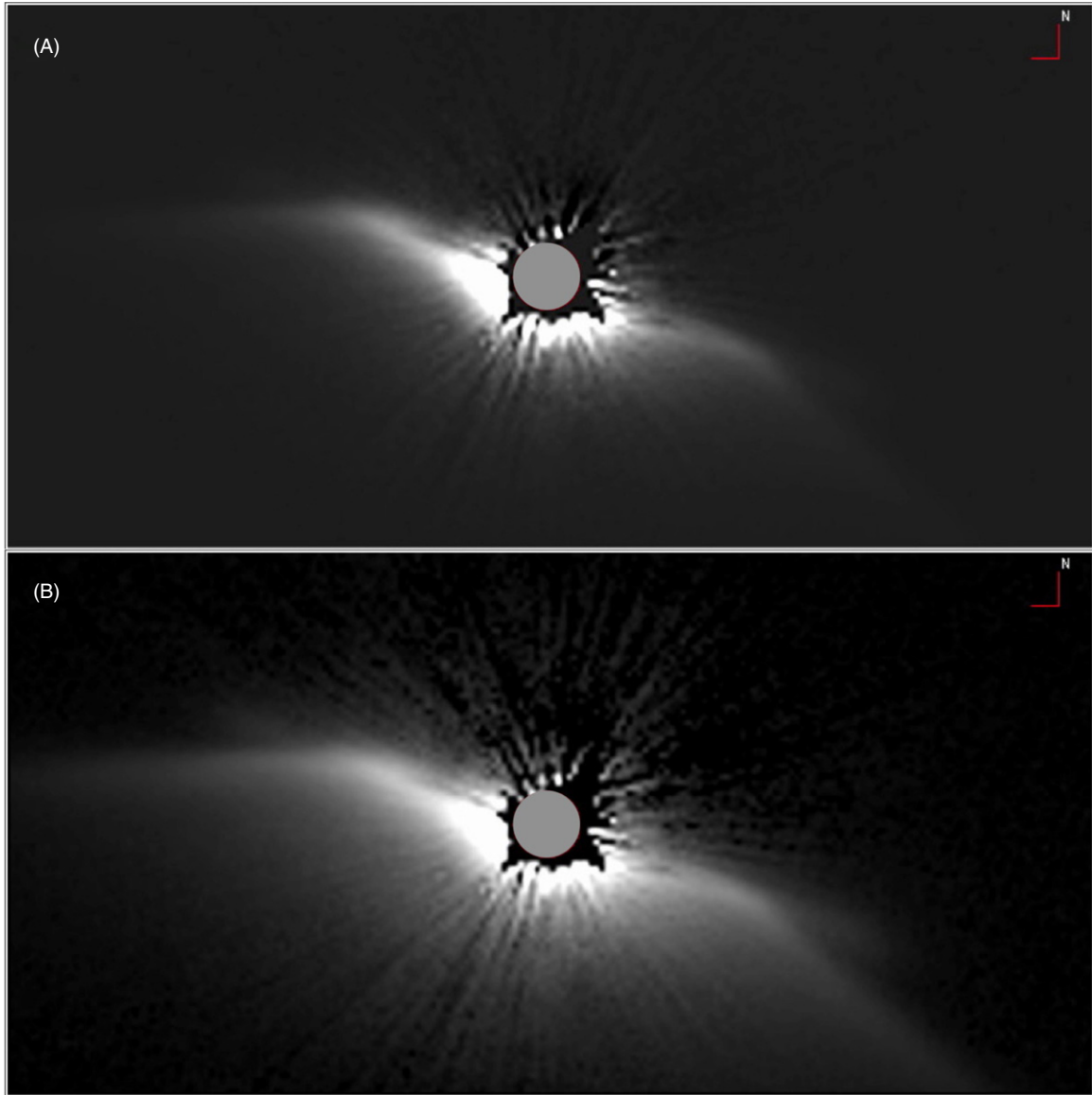


Figure 21. (A) HD 61005 debris ring, postulated as interacting with an ISM wind. The highly inclined and centrally cleared debris ring is clearly imaged to a discernable stellocentric IWD of $\sim 0''.4$ with STIS 6R/PSFTSC. Linear display in SB from -1 to $+8$ counts s^{-1} pixel^{-1} (approximately -0.177 to $1.462 \mu\text{Jy arcsec}^{-2}$). (B) Lower SB material (extending to much further stellocentric distances as shown in Figure 23(D)) is seen “streaming” off of the ring in a morphologically similar manner to the “streamers” reported previously for the HD 4976A high-inclination debris ring (Thalmann et al. 2011). Same display range as panel (A), but with a square-root stretch. FOV: $5''.07 \times 10''.14$, north up, east left. Gray circle radius, $r = 0''.3$. (A color version of this figure is available in the online journal.)

Panel (D) shows the full spatial extent of the “skirt” of (likely small) dust particles emanating from the previously unseen inner part of this complex debris structure that may be escaping the system, pushed-back from the disk-plane as the system plows through the ISM, to the full sensitivity limits imposed by the instrumental and sky background noise.

A.6. HD 92945

Introductory notes. The HD 92945 light-scattering debris disk, circumscribing its 80–300 Myr old, 21.4 pc distant, K1V central star, was discovered by Golimowski et al. (2011) with ACS coronagraphy. Their two-band optical imaging informed on the existence of a two-component axisymmetric disk inclined $\sim 28^\circ$ from edge-on seen beyond the $r > 2''$ IWA_{effective} of the ACS observations. The outer disk component was detected to a stellocentric distance of $\sim 5''.1$ (110 AU) along the disk

major axis. Golimowski et al. (2011) describe and model a diffuse, featureless outer disk component declining in SB with stellocentric distance, “truncated” at approximately $r = 5''.1$. Also suggested is a co-inclined higher SB ring extending from inner ring from $r = 2''$ to $3''$. In their F606W image (exploring the disk to smaller stellocentric angles than with their F814W image), they found the radial SB profile of the disk along its major axis is local peaked (and globally maximum) at the $r = 56$ AU in the inner ring and also at $r = 100$ – 110 AU, just interior to the outer disk truncation radius (see their Figure 4). These peaks were more pronounced on the west side of the disk than on the east side.

Observations and PSF subtraction. STIS six-roll PSFTSC observations of HD 92945 and its PSF template star HD89585 were obtained on 2011 May 6 (visits 65–68) and on 2012 February 29 (visits 61–64). All observations executed nominally, and all data were obtained according to the phase

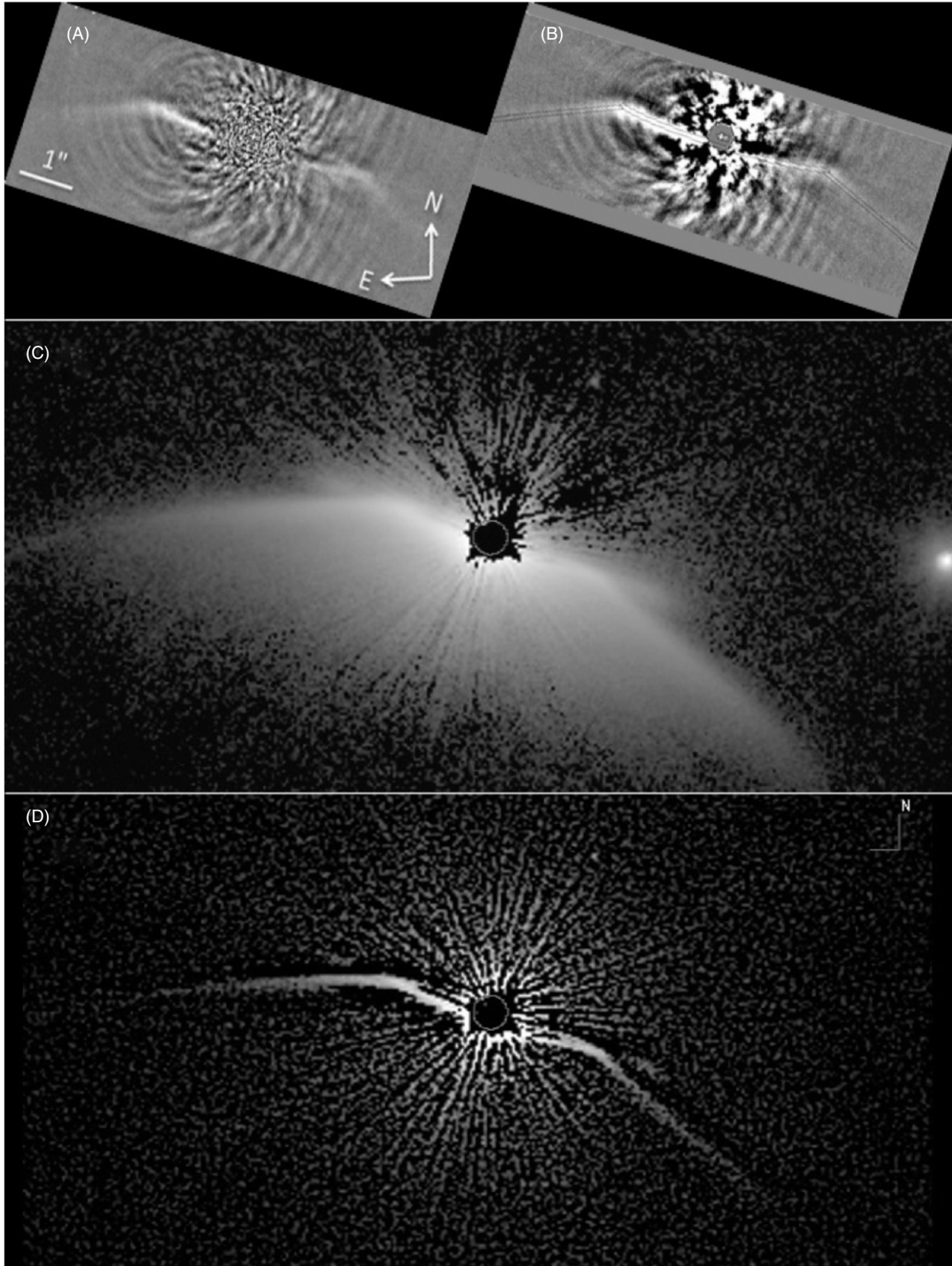


Figure 22. Panels (A) and (B) from Buenzli et al. (2010). Panel (C): \log_{10} display from $[-3]$ to $[+2.13]$ dex counts $\text{s}^{-1} \text{pixel}^{-1}$, FOV 400×250 STIS pixels ($20''.28 \times 10''.14$). Panel (D): gradient-edge enhancement by high-pass spatial filtering (at the expense of flux conservation) emphasizes the inflection of the high(er) SB ridge where the edge of the skirt joins the ring ansae. \log_{10} display from $[-3.5]$ to $[0]$ dex counts $\text{s}^{-1} \text{pixel}^{-1}$. Panels (C) and (D): red circle $r = 0''.3$ (6 STIS pixels = physical IWA angle limit of occulting WedgeA-0.6).

2 observation plan. However, the visit 63 WedgeA-0.6 and -1.0 PSF template images were locally corrupted by the presence of a close angular-proximity field star whose existence was previously unknown. That same field star is serendipitously occulted by Wedge A in the visit 67 template images, and with that had no

impact. Because of this, the visit 67 WedgeA-0.6 and -1.0 template images (with the offending field star occulted) were tested against the visit 61, 62, and 64 disk images and (fortuitously) were found to be acceptable matches to produce PSF-subtracted images for those visits. For this reason, in *addition* to the visit

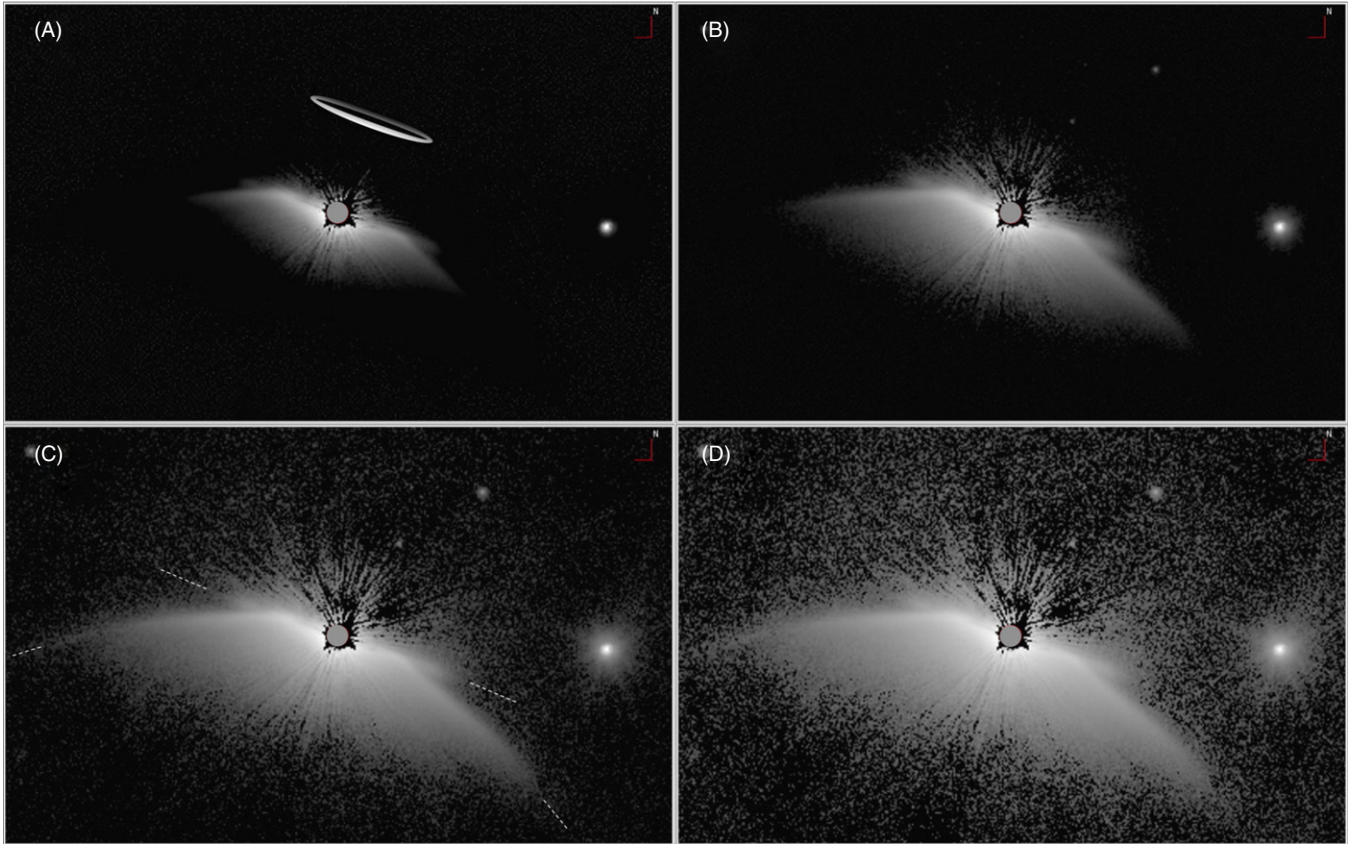


Figure 23. HD 61005 six-roll combined PSFTSC imaging (north up, east left) from all WedgeA-0.6 and WedgeA-1.0 observations. Log₁₀ display stretches A–D from [−1], [−2], [−3], [−4] to [+1.5] dex (all panels) counts s^{−1} pixel^{−1} (∼18, 1.8, 0.18, 0.018 to 5600 mJy arcsec^{−2}). Field: 400 × 250 STIS pixels (20′.28 × 10′.14)—north up, east left. Red circle $r = 0''.3$ (six STIS pixels = width of WedgeA-0.6). The morphology, geometry, and “front-to-back” SB asymmetry of a now well-seen highly inclined debris ring (previously identified in ground-based LOCI and ADI imaging by REFS) is suggested by the illustration to scale in panel (A). The nearly symmetric observed “ridges” of diffuse scattering grains, morphologically bifurcated in the skirt of lower-SB material to the SE of the star are indicated (compare also to HD 11515, panel (E)).

(A color version of this figure is available in the online journal.)

(61, 62, 64)-minus-visit 63, and visit (65, 66, 68)-minus-visit 67 PSF-subtractions, a third set of PSF-subtracted images, visit (61, 62, 64)-minus-visit 67, was added to the image reduction to provide information on the region otherwise corrupted by the background star appearing in the visit 63 template. In subsequent reduction and analysis of this data set, and the creation of an AQ image after PSF subtractions, a set of 18 PSF-subtracted images (9 WedgeA-1.0 and 9 WedgeA-0.6) was used, with the background-star-affected region in the visit 67 template digitally masked (and rejected). The resulting 6R/PSFTSC AQ image of the HD 92935 CS debris disk is shown in Figure 24(A). For completeness, we note that the inverse was also tested, i.e., using the visit 63 template for disk visits 65, 66, and 68, but was found inferior in all three visits to the visit 67 template due to differences both in breathing and centering, and so was not used in this manner.

The STIS 6R/PSFTSC AQ image reproduces the main features of the disk as described by Golimowski et al. (2011) and probes to smaller stellocentric distances for (potentially brighter) dust in the interior. The inner ring, inside $r \approx 3''$, is well seen on at least the west side of the star and with lesser visibility on the east side, through the presence of both some low-amplitude chromatic (despite $\Delta(B - V) = +0.01$ target-template color matching), and high spatial frequency radially directed breathing residuals remaining in this image. The chromatic residuals are manifested primarily as (positive and negative)

circularly concentric rings impressed upon the image detectable at approximately $r < 3''$. To partially mitigate their presence, to a simple first order, their azimuthally symmetric brightness was estimated and then subtracted in the following manner. First, a digital mask was created to exclude flux from the large, diffuse, outer disk regions (and two field objects); see Figure 25(A). Second, a digitally masked median radial profile derived from image (A) was transformed to a two-dimensional image; see Figure 25(B). To first order, this primarily represents the largely azimuthally symmetric chromatic residuals only, though other information was included in the initial image. The high spatial frequency radial “tendrils” (positive and negative emanating from the star due to uncompensated breathing errors) are conserved in total flux (or nearly so) circumferentially at any radial distance from the star, so on the azimuthal median are at (or very close to) zero. Flux from the inner disk ring itself is confined to a narrow ellipse, occupying only a small fraction of any annular zone in area, and so is not filtered out with azimuthal medianing. Third, the panel (C) difference image ((A)–(B)) subtracts the empirically determined estimation of the radial biases from chromatic residuals. While the high spatial frequency breathing tendrils are not eliminated, visibility of the inner debris ring with reduced chromatism is improved (compare panels (A) and (C)). In Figure 24(B) we separately additionally enhanced the visibility of the inner ring by low-pass spatial filtering the image in Figure 25(C) with a 0.4 pixel

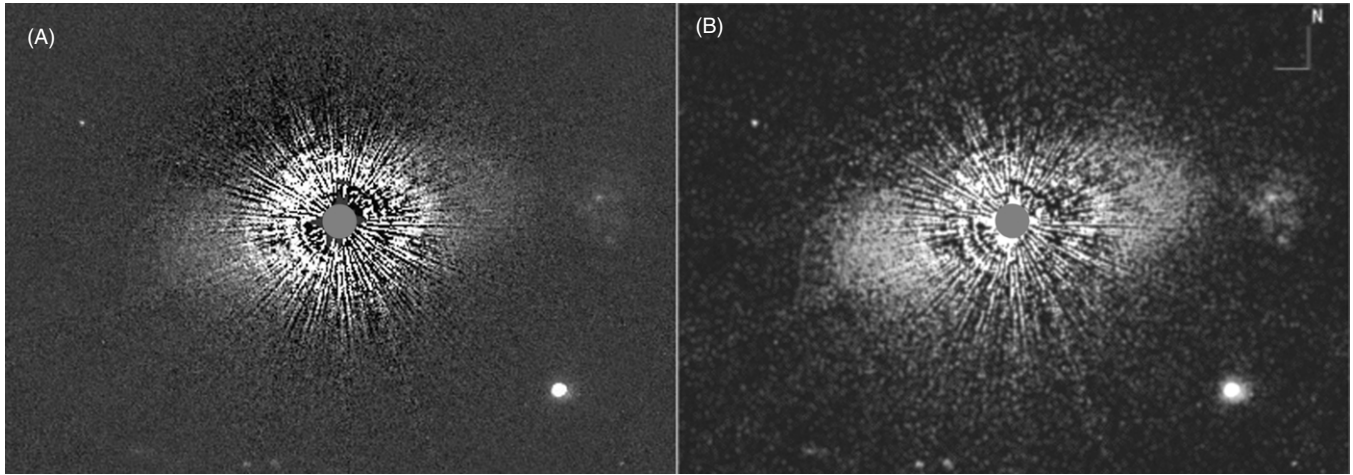


Figure 24. HD 92945 6R/PSFTSC imaging using all WedgeA-0.6 and WedgeA-1.0 observations combining 18 image-artifact masked PSF-subtracted images at six target orientations. (A) Linear grayscale display -0.02 (black) to $+0.08$ (white) counts $\text{s}^{-1} \text{pixel}^{-1}$. (B) Same data post-processed sequentially (1) to reduce chromatic residuals in the inner disk region by disk-masked radial-median image subtraction (see the main text) and (2) to additionally improve the visibility of the inner disk by low-pass (0.4 pixel Gaussian kernel) spatial filtering to reduce the amplitude of the high-frequency breathing residuals. (B) presented with a \log_{10} display from $[-2.6]$ to $[-1.1]$ dex counts $\text{s}^{-1} \text{pixel}^{-1}$. Both: $20''.3 \times 15''.0$ FOV with north up, east left. The gray circle represents $\text{IWA}_{\text{effective}} r = 0''.46$ for these observations.

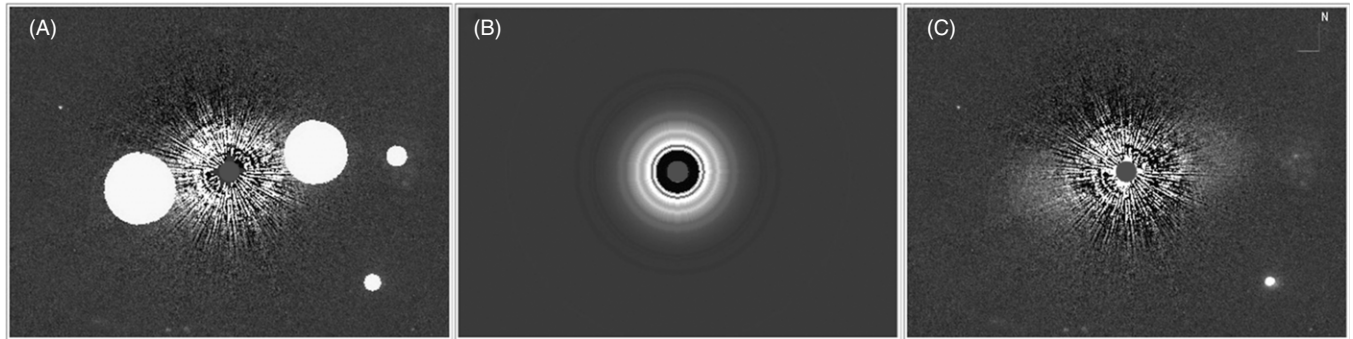


Figure 25. (Partially) Mitigating chromatic residuals. (A) Digital mask applied to the AQ image as shown in Figure 24(A). (B) Azimuthal median with digital mask applied. (C) AQ image with empirical chromatic residual mitigation.

kernel. This conservatively reduces the amplitude of the high spatial frequency tendrils without a significant effect on the lower spatial frequency ring and outer disk.

Finally, in Figure 26 (top), we more aggressively applied a 5×5 pixel ($\sim 1/4$ arcsec, similar to the 11×11 kernel used by Golimowski et al. 2011 with the ACS data given its $25 \text{ mas pixel}^{-1}$ scale) boxcar smoothing kernel to the chromatically post-processed AQ image shown in Figure 25(C). As a result, most of the high-spatial-frequency “breathing” noise that otherwise pollutes the image at approximately $1''.7 \leq r \leq 3''.5$ is mitigated.

Principal results. Our new STIS 6R/PSFTSC observations of the HD 92945, obtained at a similar central wavelength as the ACS F606W filtered observations, confirms the previously discussed observational results of Golimowski et al. (2011). The STIS AQ image reveals (at least) the eastern side of the inner ring with somewhat better clarity, and the northern extent of the ring closer to its sky-projected minor axis (e.g., compare also to Golimowski et al. 2011, see Figure 3). From the STIS data, we find the 50CCD band integrated brightness of the disk from $r = 1''.67$ (~ 36 AU on the disk major axis) outward, largely unaffected by residuals from imperfect PSF subtractions, to be 0.143 mJy , and thus a disk-scattering fraction, $F_{\text{disk}}/F_{\text{star}} \sim 0.0051\%$ at $r > 1''.67$. The uncertainty in the absolute calibration both of these measurements is estimated as $\pm 20\% 1\sigma$. This is in good agreement with what Golimowski et al. (2011) found

with $F_{\text{disk}}/F_{\text{star}} = 0.0069\%$ fitting a full disk model to their observations in the F606W band of ACS.

The STIS post-processed AQ image (Figure 26, top) clearly resolves both the HD 92945 CS debris ring at $r \sim 2''.77$ on the disk major axis, and the circumscribing “featureless” outer disk. The debris inner ring (though somewhat noisy) is well seen in 90° wide sectors flanking the disk major axis on both sides of the star. Assuming intrinsic bi-lateral symmetry, we find the celestial orientation of the disk major axis at P.A. $= 100^\circ \pm 2^\circ$ (in agreement with Golimowski et al. 2011). We then produce a radial SB profile along the disk major axis from the image (Figure 26, bottom), without quantitative error estimation) to aid in interpretation of the image.

Two significant disk asymmetries are seen. (1) The outer disk appears to be brighter on the east side than on the west side of the star beyond the inner debris ring. (2) The eastern ansa of the inner debris ring appears inwardly “distorted” in comparison to the western ansa. The latter may be due to a pericentric offset of the debris ring, but this is conflated with remaining PSF-subtraction residuals at the level of measurement uncertainties, and confirmation of this suggestion must await further improved, future, observations.

Modulo uncertainties in stellar ages (see Table 1) HD 92945 may be the oldest debris disk in our GO 12228 sample for which an inner debris ring within a larger debris system is clearly resolved. However, similar large-scale debris has been reported

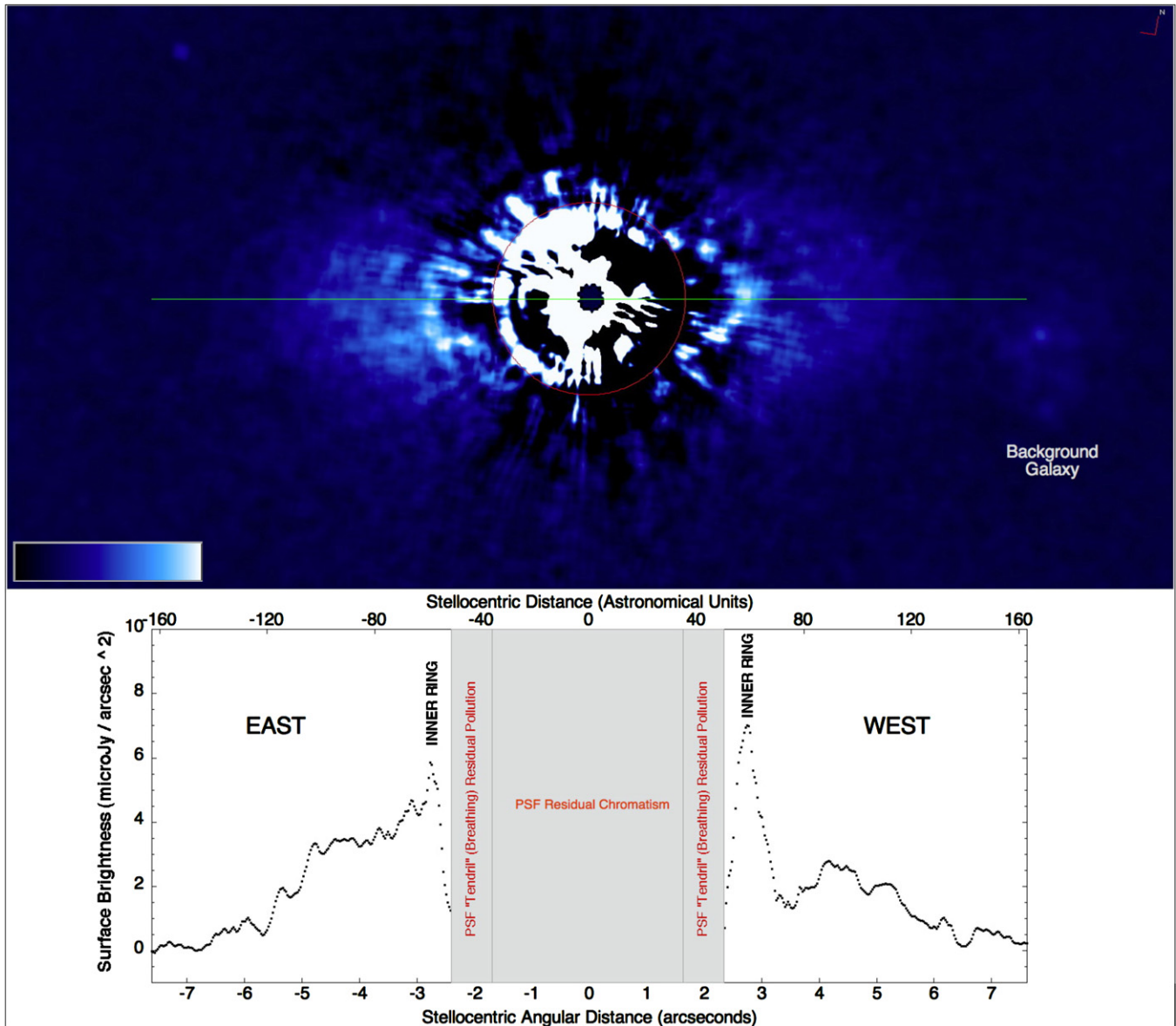


Figure 26. Top: STIS 6R/PSFTSC image of the HD 92945 CS debris disk post-processed to remove residual chromatic and high-spatial-frequency breathing residuals at $r > 1''.68$ (33 pixel radius indicated by the red circle). Linear display stretch from -0.01 to $+0.045$ counts s^{-1} pixel $^{-1}$. The green line indicates the disk major axis. Bottom: radial SB profile measured along the disk major axis with the distance scale directly corresponding to the image above.

(A color version of this figure is available in the online journal.)

for the 440 ± 40 Myr old Fomalhaut system (Kalas et al. 2008, 2013; Mamajek et al. 2013) and may be a routine feature, at low-SB levels, of many debris disks.

A.7. HD 107146

Introductory notes. HD 107146 is a 27.5 pc distant, 80–200 Myr old, close-solar analog star (spectral type G2V) with a large IR excess, $L_{\text{IR}}/L_{\text{star}} = 0.12\%$, attributed to several tenths of an Earth mass of re-radiating dust orbiting the star. The presence of a broad starlight-scattering CS debris ring about the star was reported and discussed in detail by Ardila et al. (2004, 2005). The disk, then seen in ACS/HRC PSF template-subtracted coronagraphic observations at 0.6 and 0.8 μm , was suggested to be inclined $\approx 25^\circ$ from pole on and broadly peaking in radial SB in both bands $\sim 4''.7$ (130 AU) from the central star. While the ACS observations have been successfully used to ascertain the primary characteristics of the disk, the above authors

have stated “{PSF} subtraction errors, caused by mismatches in the colors of the stars or PSF time variability, dominate the emission within $\sim 2''$ from the star and contribute light at large distances. With only one reference star, we cannot quantify very precisely the magnitudes of these two error sources”, (p. L148). The debris belt was initially detected in the millimeter (Williams et al. 2004) and has been imaged in the sub-millimeter (Corder et al. 2009) and modeled (Ertel et al. 2011). The availability of these data plus HD 107146’s status as a solar analog made HD 107146 a natural candidate for 6R/PSFTSC.

Observations and PSF subtraction. All observations of HD 107146 and its PSF reference star HD 120066 were executed nominally on 2011 February 22 and 2011 May 30. After visit-level PSF subtraction, it was obvious that the V24 PSF-subtracted images for both WedgeA-0.6 and -1.0 were clearly inferior to the other visits with “breathing”-driven PSF-subtraction residuals dominating the inner regions of the disk. The (initial) inclusion of the V24 data into a six-roll combined AQ image,

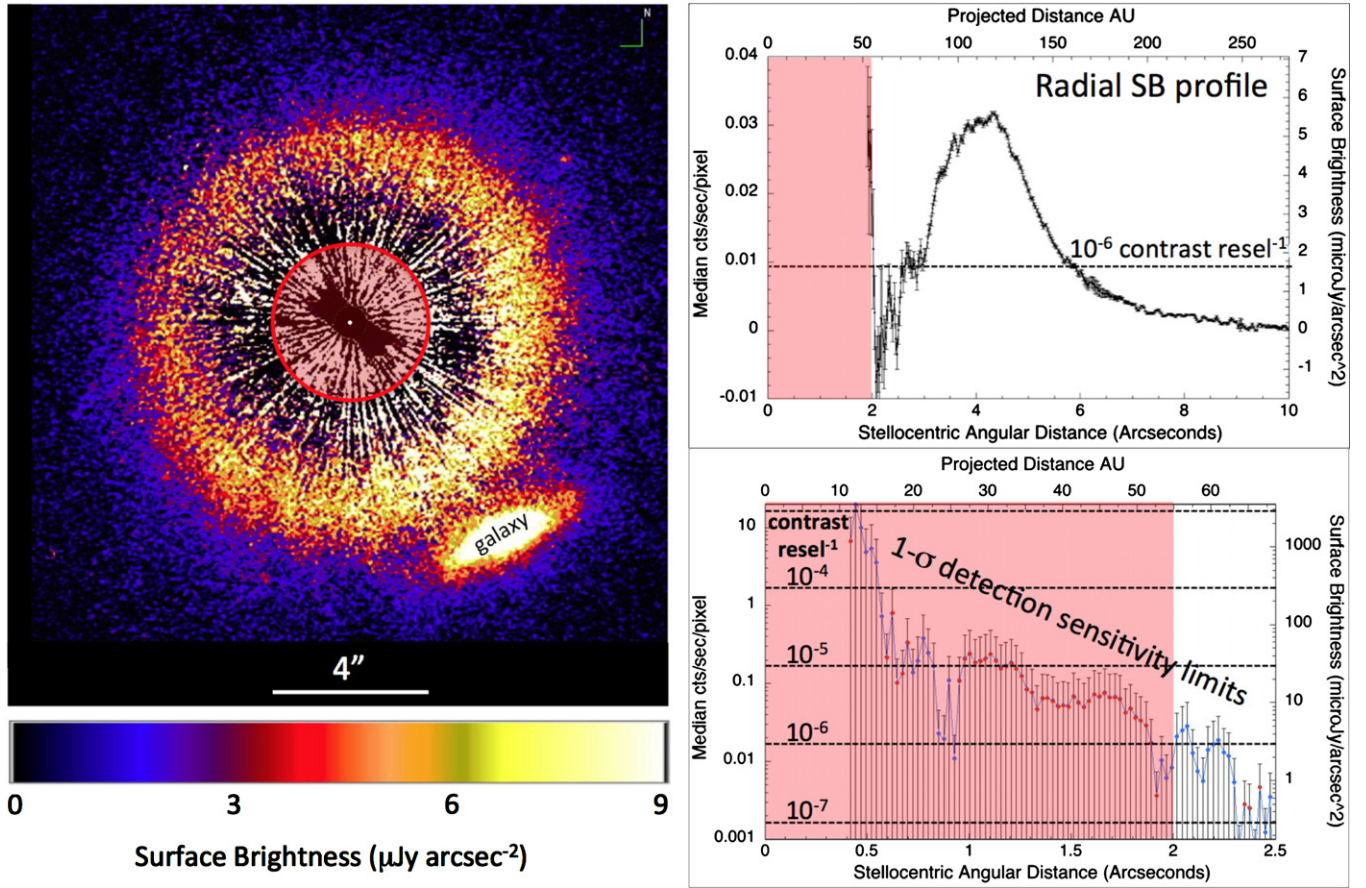


Figure 27. Left: STIS 5R/PSFTSC image of the nearly face-on HD 107146 debris disk. $15''.23 \times 15''.23''$ FOV, north up, east left. The red circle indicates the region where the brightness of PSF-subtraction residuals dominate over disk-light declining toward smaller stellocentric distances interior to radius of peak brightness at 10^{-6} resel^{-1} and lesser contrasts. Top right: azimuthal median radial SB profile in the non-contrast-limited regions of the disk. Error bars plotted on the radial SB profile represent the 1σ standard deviations about the median brightness measured in 1 pixel radial increments fully around the star. Absolute calibration uncertainties due to possible under- or oversubtractions of the template PSF of approximately a few percent at most are not included. Bottom right: 1σ contrast-limited disk detection sensitivities as a function of stellocentric distance in the presence of PSF-subtraction residuals as measured from the median absolute deviations in flux density per resel in 1 pixel wide annuli. (The red/blue points indicate positive vs. negative median residuals).

(A color version of this figure is available in the online journal.)

in a systematic inter-visit manner as described in Section 5.7, marginally degraded the final product rather than improving it compared to a five-roll combination excluding V24. Fortunately, this was not critical to the metrical analysis of this target’s angularly large, broad, and nearly face-on ring-like disk that is spatially well sampled in the five other, nearly equally superior, visit-level images included in then a five-roll combination. The five-roll combined AQ image is presented in Figure 27 (left panel). The interior region of the disk, along the direction of and flanking its major axis (SE to NW), is imaged to an $\text{IWA}_{\text{effective}} = 0''.41$, about 2 resels beyond the edge of the occulting wedge. The $r \leq 1''$ CS region in the 5R/PSFTSC combined image roughly in the direction of the disk minor axis only (depicted as the contiguously hard-black stellocentric region in Figure 27) is unsampled due to the rejection of the V24 imaging data that uniquely tiled that space. Separately, two three-roll combined images, divided by epoch inclusive of the V24 data in the second epoch, were created for the purpose of investigating potential companionship by common proper motion for point sources identified in close proximity to the disk not adversely affected by breathing degradation of the V24 disk image itself.

Principal results. The 5R/PSFTSC AQ image of the HD 107146 disk improves significantly in several regards over the discovery imaging obtained by Ardila et al. (2005, see

Figure 1). The new STIS image is presented in Figure 27 with a linear display stretch over a dynamic range and with a color table matching as closely as possible that of the ACS F606W discovery image to enable a direct visual comparison. A background galaxy, seen equally well in the Ardila et al. discovery imaging, now marginally intrudes into the outer periphery of the broad ring-like disk, as it has been “closing” on HD 107146 due to the stellar (reflex) proper motion since the 2004 epoch discovery imaging. Circa 2020, the galaxy will be superimposed upon the intrinsically brightest part of the disk, radically altering its apparent SB distribution then and for a decade to follow (a warning to future observers).

Improving on earlier ACS and NICMOS data, the STIS AQ image sensitively traces dust depletion in the inner regions of the broad HD 107146 debris ring to within ≈ 60 AU of the star, half the stellocentric distance from where the debris ring is brightest. The HD 107146 CS debris disk is intrinsically quite faint both in terms of its STIS 50CCD band integrated flux density, $F_{\text{disk}} = 0.40$ mJy at $r \geq 2''.0$ (~ 55 AU; radius of the red “here be dragons” circle in Figure 27), and its visible-light-scattering fraction $F_{\text{disk}}/F_{\text{disk}} = 7.7 \times 10^{-5}$ ($\pm 4.0 \times 10^{-6}$), and in comparison to its 16 times larger IR excess. An azimuthal median radial SB profile (measured with the background galaxy masked) is shown along side of the disk image in Figure 27 with

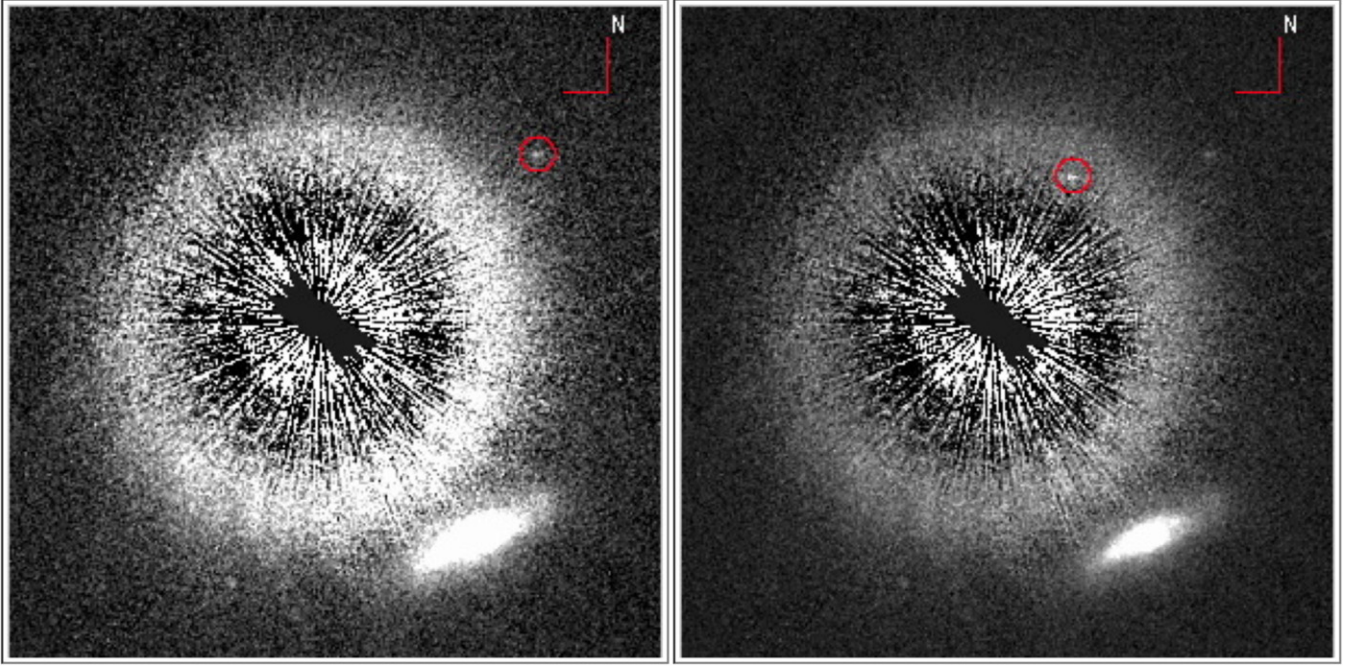


Figure 28. Identified, but rejected, companion candidates found to be background stars from their reflex motions relative to the location of HD 107146 ascertained with “X marks the spot” differential astrometry (see Section 6.4).

(A color version of this figure is available in the online journal.)

$SB_{\text{peak}} = 5.6 \mu\text{Jy arcsec}^{-2}$ at a stellocentric median distance $r = 4''.3$.

While dust depletion interior to the radius of peak SB was inferred from the prior ACS data, the extent and depth of the clearing was uncertain (e.g., see Figure 2 of Ardila et al. 2004). The STIS SB image and its azimuthal median radial profile (Figure 27) clearly show a steep decrement in SB interior to the broad ring with the region cleared of light-scattering debris to a 1σ per resel sensitivity limit of $\sim 0.2 \mu\text{Jy arcsec}^{-2}$ (image contrast $\sim 2 \times 10^{-7}$) at $2''.3$.

The ring SB declines more rapidly interior of r_{peak} toward the star and less steeply radially outward from r_{peak} . The median photometric FWHM of the ring, as one characterizing metric of the ring width, $W_{\text{FWHM}} = 2''.2$ (≈ 60 AU), with $W_{\text{FWHM}}/D_{\text{peak}} = 0.34$ (significantly broader than the HR 4796A (Schneider et al. 2009) and HD 181327 ring-like debris belts).

We note here that the debris belt and its inner clearing revealed with STIS 5R/PSFTSC exhibit a high level of bi-axial symmetry, with no obvious substructures or distortions of detectable significance around the ring, such as have been seen in sub-millimeter interferometry (Corder et al. 2009). Conversely, though not explicitly noted by Ardila et al. (2004), the ACS F606W discovery image, see Figure 1 (at similar wavelength as the STIS image), on visual examination appears to be asymmetrically extended to a greater peak radial distance along the SW disk major axis. This is also reflected (when examined in detail) in their F606W SB profile along the disk major axis, (see Figure 2) and corresponding Tau \times Omega plot (see Figure 3) derived from that image, but is not (obviously) seen in their F817W discovery image. With the higher-fidelity STIS image and lack of correlation in the asymmetry implied only in the ACS F606W image, we suggest that this was an optical artifact from pollution by incomplete PSF subtraction beyond $r > 2''$ and/or imperfect geometrical distortion correction in the ACS F606W image.

From the STIS data we can also quantitatively assess SB limits on any dust undetected between $0''.4$ – $2''.3$ (Figure 27, bottom right) where dust detection by scattered light becomes contrast-limited. Therein, we also cast this in terms of image contrast. With such dimensionless quantification, we suggest that HD 107146 would make an ideal “test case” for more aggressive state-of-the-art ground- and space-based coronagraphs now under design and development.

Two point-like objects (putative low-mass companion candidates) were detected in close proximity exterior to, and superimposed upon, the broad HD 107146 debris ring at $r = 6''.40$, P.A. = $300^\circ.3$ (Figure 28 left panel), and $r = 3''.81$, P.A. = $307^\circ.05$ (Figure 28 right panel); two stretches of the same 5R/PSFTSC image. Due to the relatively large proper motion of HD 107146 ($[-174.16, -148.90]$ mas yr^{-1}) and its annual parallax, in concert with an uncertainty in the location of the occulted star of only ~ 5 mas in the SIAF (see Section 5.4), both objects were found to be non-commoving background objects over the three-month interval between the two visit sets.

A.8. HD 139664

Introductory notes. HD 139664 is a 17.4 pc distant F5 main sequence star with a starlight-scattering CS disk that was revealed at stellocentric distances $> 3''$ (52 AU) with *HST*/ACS coronagraphic imaging (Kalas et al. 2006). Its estimated age of 300 (–200, +700) Myr, comparable to Fomalhaut and possibly as old as HD 51354 (1 Gyr, see Section A.4), makes it one of the oldest stars for which a CS debris disk has been imaged in scattered light. HD 139664 also has the smallest IR excess, $L_{\text{IR}}/L_{\text{star}} = 0.009\%$, of any of the stars on the GO 12228 sample and the general population of about two dozen CS debris disks that have been imaged in scattered light to date, and follows the general trend of IR excess declining with stellar age (Figure 9). The ACS scattered-light discovery image revealed

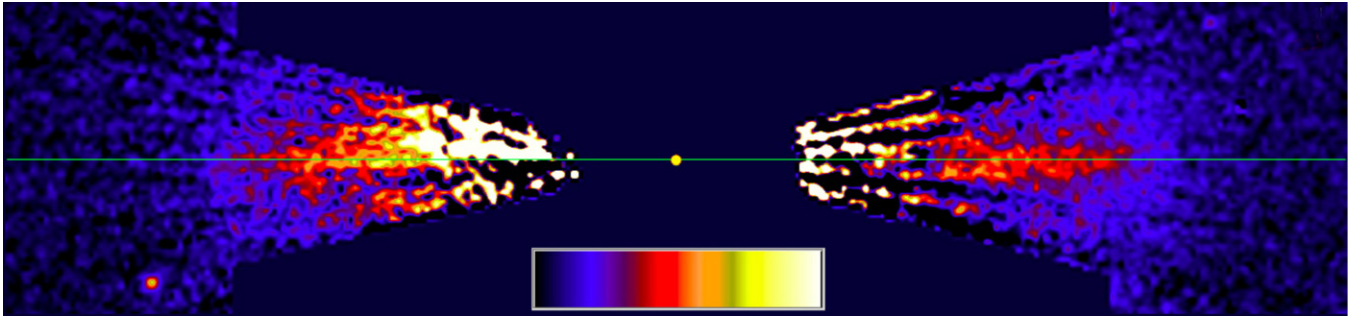


Figure 29. STIS AQ image of the HD 139664 CS edge-on debris disk oriented with the morphological major axis (green line) on the image horizontal (celestial P.A. of the disk major axis estimated from this image as $75^{\circ}5 \pm 0^{\circ}5$). Linear display stretch from -0.02 to $+0.44$ counts s^{-1} pixel $^{-1}$. FOV $15''.23 \times 3''.55$. IWA = $1''.2$. (A color version of this figure is available in the online journal.)

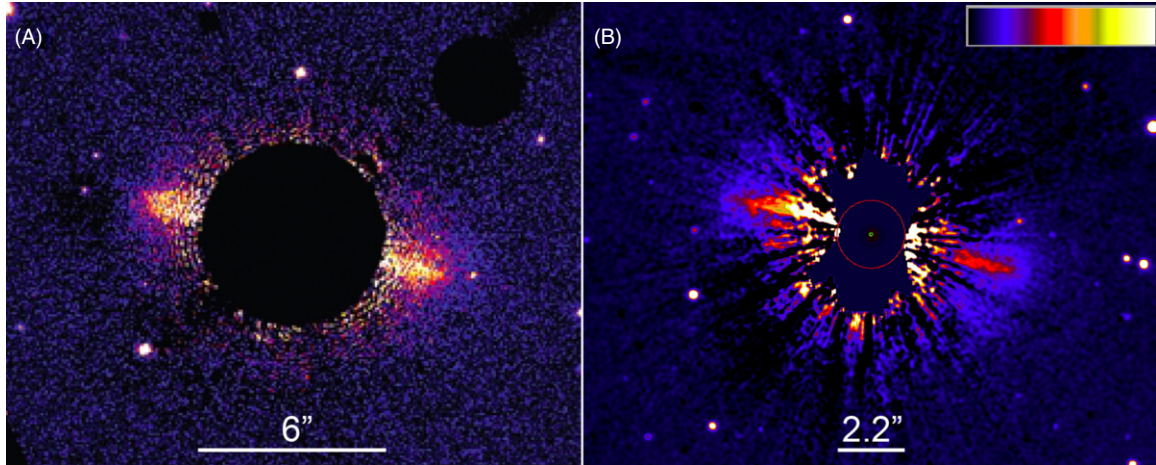


Figure 30. (A) ACS discovery image of the HD 139664 CS debris disk, $r = 3''$ IWA indicated by the black digital mask (reproduced from Kalas et al. 2006). (B) STIS 3×3 boxcar smoothed AQ image with -0.02 to $+0.44$ counts s^{-1} pixel $^{-1}$ display stretch similar to the ACS image for ease of comparison. The STIS image probes closer to the star (red circle, $r = 1''.2$), though in the presence of remaining PSF-subtraction residuals in region masked in the ACS image.

(A color version of this figure is available in the online journal.)

the ansal regions of a faint, apparently edge-on, debris disk from a contrast-limited IWA $\approx 3''$ (54 AU) to a stellocentric distance of $\approx 6''$ (104 AU) suggested by Kalas et al. (2006) to be the outer boundary of the disk (see Figure 30(A)). No SB asymmetries were detected between the two diametrically opposed sides of the disk. A minor-axis-mirrored (to improve S/N) radial SB profile suggested a peak in the dust density at $r \approx 83$ AU, the radius of a postulated inclination-obscured planetesimal belt.

Observations and PSF subtraction. The STIS observations of HD 139664 and its PSF star HD 99353 executed nominally at two epochs, 2011 May 23 (V95–V98) and 2011 July 31 (V91–V94), and would prove to be the most challenging of all the GO 12228 targets to improve over previous observations. HD 139664 itself at $V = 4.64$ is significantly brighter than our other disk target stars (see Table 1). Thus, to permit a sufficient imaging dynamic range across the inner to outer disk regions in the face of image saturation in the region adjacent to the STIS occulting WedgeA, for HD 139664 and its commensurately bright PSF star, we imaged using WedgeA-0.6 with both very short and intermediate duration exposure times, in addition to the deeper (centrally saturated as expected) WedgeA-1.0 imaging; see Table 3 for summary details.

Following our homogeneous multiple-orientation observing paradigm across the target set, even with foreknowledge of the edge-on disk major axis P.A. from the ACS observations, we did not constrain the absolute (celestial) orientation of the disk in any of the STIS visits to enable us to probe for and study any

out-of-plane material that, a priori unknowingly, might have existed (e.g., as seen in the cases of the HD 61005 (Section A.5) and HD 32297 (Section A.3) nearly edge-on disks).

All of the PSF-subtracted WedgeA-1.0 images, as well as those of intermediate exposure depth at the WedgeA-0.6 occulter position, suffered from rotationally invariant chromatic residuals from a small color mismatch ($\Delta(B - V) = -0.10$) in the PSF star (shown similarly, for example, in the case of HD 32297 in Figure 17) causing a zonal undersubtraction of the disk closer to the star. These were (separately) corrected for the data acquired at the two different wedge positions by same method described in Section A.3.

With the inter-visit relative orientations unconstrained against absolute celestial P.A., the disk major axis itself was unobscured by the STIS wedge, OTA diffraction spikes, and IWA saturation region beyond the mask edge, in only two of the six STIS visits (V96 and V98) constrained by a combination of guide star and target visibilities. This limited, in aggregate, the sensitivity depth of detection in the outer regions of the disk along the disk major axis explored with the (deeper) WedgeA-1.0 exposures. V95 was significantly degraded by breathing residuals at both wedge occulter positions after PSF subtraction, and hence was rejected in the creation of the final multi-roll (chromatism corrected) AQ image shown in Figure 29.

Principal results. The Figure 29 AQ image reproduces and confirms the main features of the debris disk observable with ACS reported by Kalas et al. (2006), with the

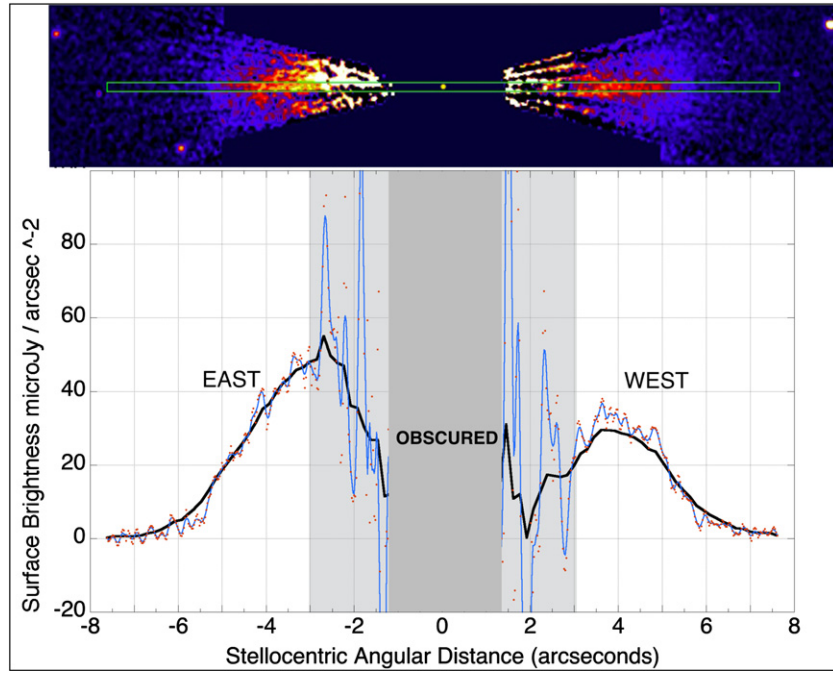


Figure 31. Radial SB profile along the disk major axis in 1 pixel radial increments (red points) in a 2 pixel (~ 1 resel) wide strip $\pm 8''$ long centered on the star (green strip superimposed on disk image at the same horizontal angular scale). The blue line indicates 3 pixel wide boxcar smoothing (corresponding to spatial filtering in Figure 30(B)). The black line further adaptively filters out the high-spatial-frequency noise. The region in light gray was inaccessible to ACS. The dispersion in the measures over small spatial scales is indicative of the relative measurement uncertainties that become breathing-dominated, when unfiltered, close to the star.

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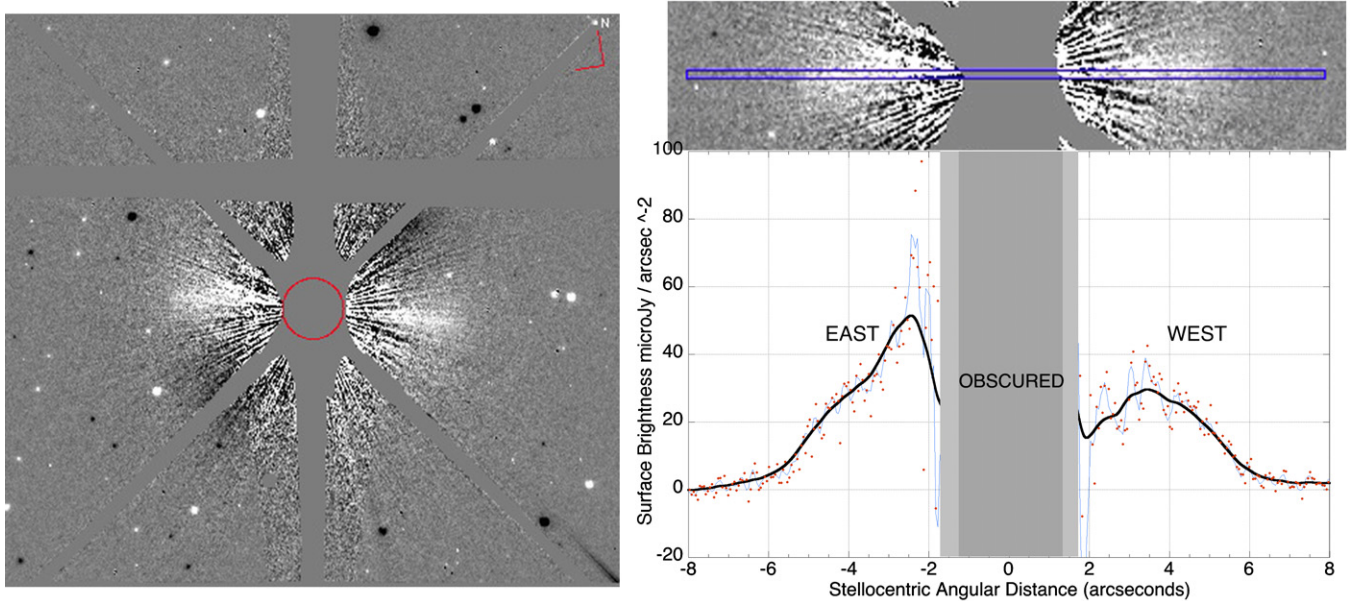


Figure 32. Left: using the off-rolled V95 images as a PSF template for the visit 98 images of the disk to intermediate exposure depth probes the CS region at a non-saturating IWA of $1''.2$ (radius of red circle) for the presence (or absence) of bright dust. Linear display stretch: -0.2 to $+0.2$ counts s^{-1} pixel $^{-1}$. Right: radial SB profile along the plane of the disk using the achromatic V95 “off-rolled” target PSF. The $1''.2 \leq r \leq 1''.7$ region masked in light gray is photometrically unreliable due to background biases from the off-rolled imprint of the disk at these stellocentric distances.

(A color version of this figure is available in the online journal.)

following differences deviating from their determination that “no surface brightness asymmetries were detected between each side of {the} disk”, (p. L57). To facilitate visual comparison, we additionally suppress still remaining high-spatial frequency breathing residuals in the fully circum-azimuthal FOV by applying a 3×3 pixel boxcar kernel as a low-pass spatial filter. The result is shown in Figure 30(B). An examination also of a radial

SB profile along the disk major axis (prior to smoothing, i.e., derived from the image in Figure 29) in Figure 31 confirms an “east/west” SB asymmetry.

Though somewhat polluted with PSF-subtraction artifacts, starlight-scattering dust on the east side of the edge-on disk is traceable in the STIS AQ image to a (masked, but saturation limit obscured) inner working distance of $r \sim 1''.2$ (21 AU),

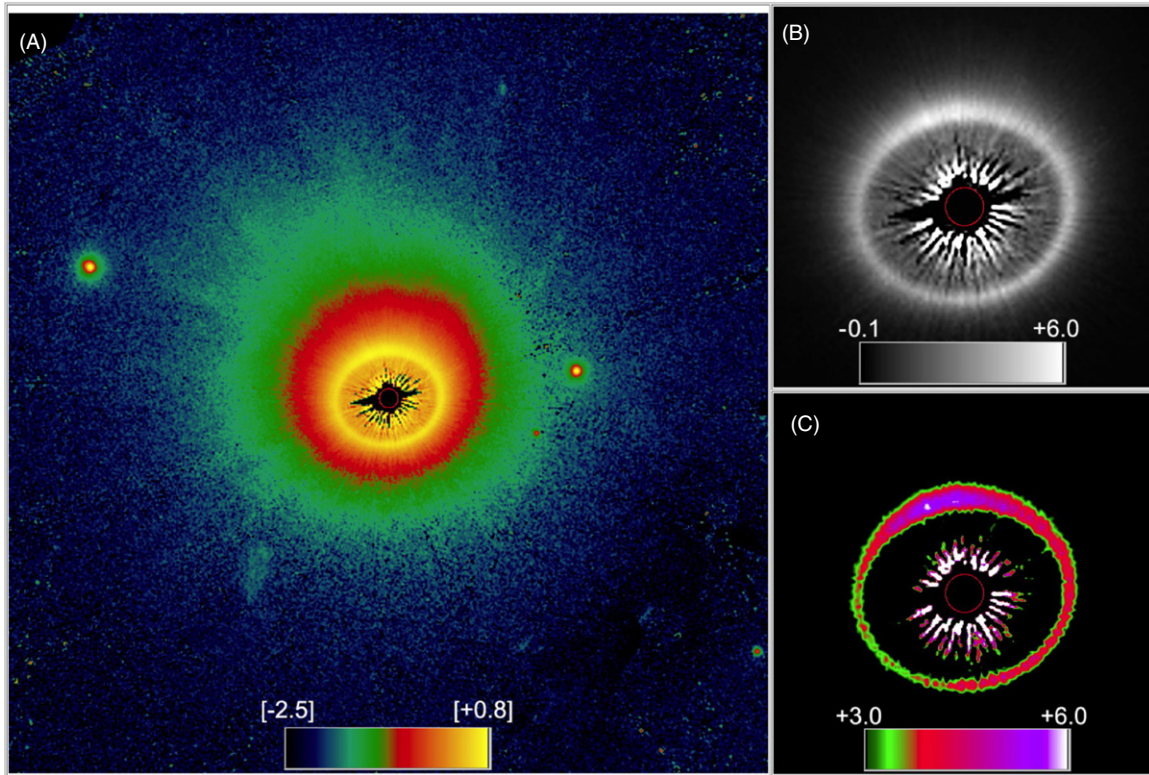


Figure 33. STIS six-roll AQ image of the HD 181327 CS debris disk (celestial north “up”) utilizing all 12 WedgeA-0.6 and WedgeA-1.0 images visit-level PSF-subtracted images in artifact-masked median combination. (A) $24''.36 \times 24''.36$ FOV centered on the occulted star. Log_{10} display stretch from $[-2.5]$ to $[+0.8]$ {dex} $\text{counts s}^{-1} \text{pixel}^{-1}$ (approximately $0.6\text{--}1100 \mu\text{Jy arcsec}^{-2}$) simultaneously revealing the debris ring and its large, asymmetric, diffuse halo. ((B) and (C)) $6''.09 \times 6''.09$ FOV centered on the occulted star in two linear display stretches (range in $\text{counts s}^{-1} \text{pixel}^{-1}$ ranges indicated) show the narrow width and “sharpness” of the bright debris ring (panel (B)), and its asymmetric SB distribution (panel (C)). Central red circle (all images): $r = 0''.3$ (15 AU).

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with a SB $\sim 50 \mu\text{Jy arcsec}^{-2}$ at $r = 2''.7$ and declines inwardly by a factor of two at $r = 1''.3$ just beyond where the disk is obscured in the STIS image. On the opposite (west) side of the star, the disk SB peaks further from the star at $r = 3''.7$, reaching a maximum SB of $\sim 30 \mu\text{Jy arcsec}^{-2}$ and declines linearly. Directly opposite to the east side peak at $r = 2''.7$, the west side has only $\sim 17 \mu\text{Jy arcsec}^{-2}$; see Figure 31.

Additionally, the vertical (out of plane) morphology of the disk to the north, “above” the disk plane as shown in Figure 29, appears different on the east and west sides of the disk and, on close examination, correlates with the ACS image. On the west side, the bright part of the edge-on disk (red in the STIS image) is a narrow linear feature defining the disk plane surrounded above, below, and exterior by dust scattered starlight at lower SB (light blue). On the east side only, dust-scattered starlight extends at higher SB to a greater distance “above” the disk plane, with greatest height in the STIS image at $r \sim 2''.8\text{--}3''.7$ where the SB peaks. A similar morphology is seen in the $r > 3''$ region in the ACS image (giving a “wedge” rather than linear shape to the outer part of the eastern side of the disk).

Together, these differences implicate a non-axisymmetric disk SB, unknown to Kalas et al. (2006) in their minor-axis mirrored reflected treatment of the posited dust density distribution. Whether this results from a stellocentric offset of a geometry-obscured debris ring or asymmetric differences in the dust density distribution is an area for future investigation.

We measure the total flux density of the disk, via background (far from the star) subtracted aperture photometry conforming to

the FOV shown in Figure 30 that captures $>98\%$ of the scattered from the disk, as $260 \mu\text{Jy}$. Given the stellar V magnitude of HD 139664, we find the $0.57 \mu\text{m}$ disk-scattering fraction, $F_{\text{disk}}/F_{\text{star}} = 0.0005\%$, to be more than an order of magnitude smaller than any of the other disks in the GO 12228 sample. This excludes the region of the disk at $r < 1''.2$ inaccessible in the STIS AQ image.

In V95, the edge-on disk (mostly obscured by one of the *HST* diffraction spikes) was sufficiently “off-rolled” (by 48°) to use separately as an independent PSF-subtraction template for the V98 disk image, mitigating residual chromatism near the star as previously discussed. This was possible to a limiting non-saturating distance of $r = 1''.2$ (similar to the two-roll combined PSF-template subtracted AQ image) in the intermediate depth exposures. This worked well, but only for the wedge-unobscured V98 image of the disk due both the large roll angle differential and serendipitously very low amplitude breathing residuals in comparison to other visits. The resulting image is shown in Figure 32 (left, wherein a negative imprint of the disk in the V95 template is seen flanking one of the diffraction spike masks and partially biases the disk flux from $1''.2$ to $\sim 1''.7$). This image (which is also combined with the WedgeA-1.0 deeper imaging) does not improve over the AQ image at $r > 1''.2$, but reveals the relative absence of bright dust close to the star between $1''.2$ and $1''.7$ (locally against the biased background) providing an additional constraint for later modeling of this older and very faint CS debris system. A radial SB profile along the disk plane from this “off-rolled” subtraction (Figure 32, right) is highly

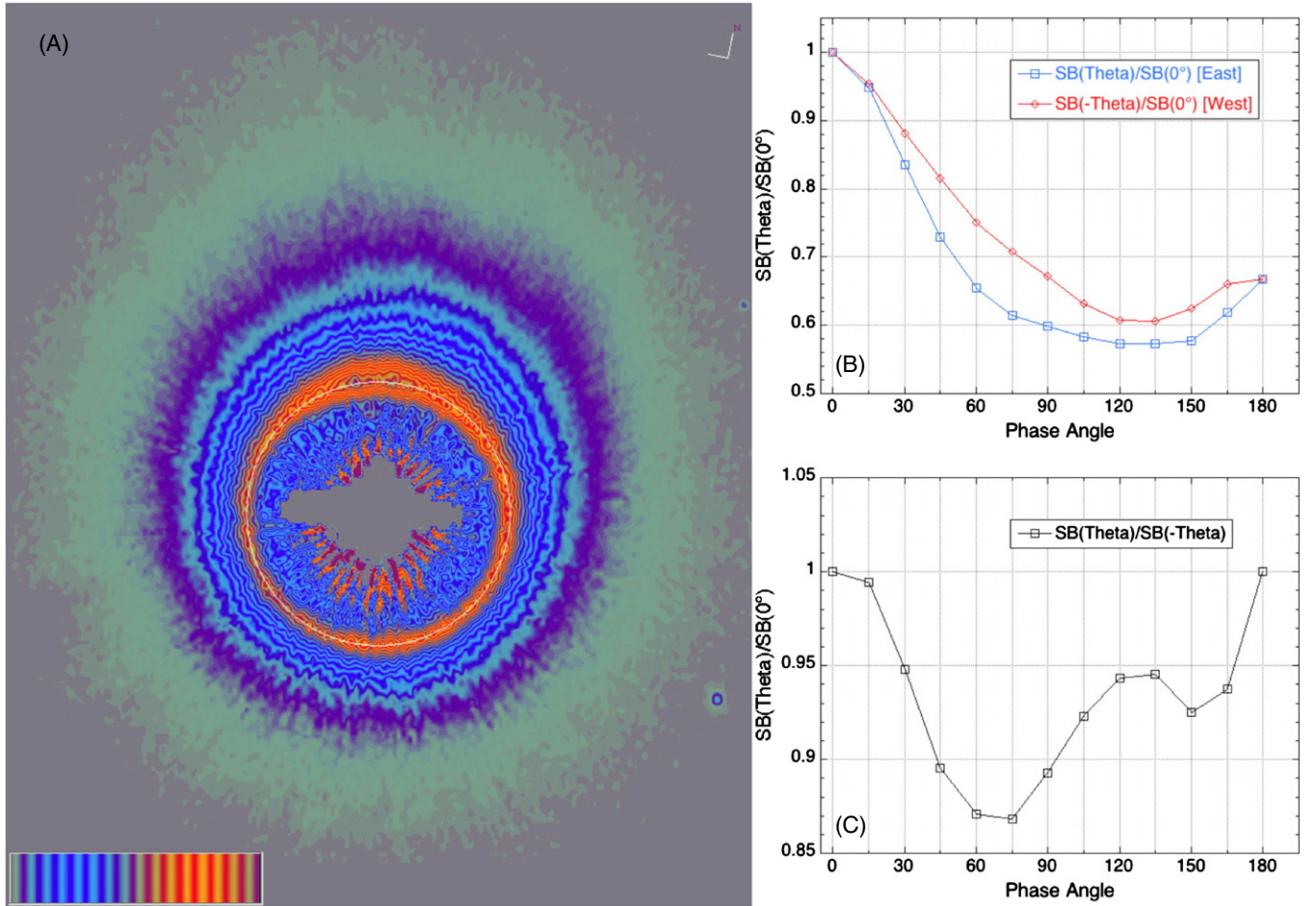


Figure 34. (A) Face-on deprojected isophotal view of the HR 181327 CS disk. 0–6 counts s⁻¹ pixel⁻¹ linear display stretch with 3.125% intensity contour intervals. Image orientation of the peak SB along ring with stellocentric deprojected azimuth angle (θ) = 0° at a celestial P.A. = 12° (“up”) in this representation. The white circle is the best circular fit to circumferential radial SB peaks. (B) Azimuthal SB profile (normalized to the θ = 0° ring-peak SB) asymmetry about the disk minor axis. (C) “Mirror symmetric” SB ratios as a function of azimuthal (θ) angle about deprojected minor axis.

(A color version of this figure is available in the online journal.)

correlated and in good agreement with the identically measured two-roll subtraction (Figure 31) where both are photometrically reliable at $r > 1''.7$.

A.9. HD 181327

Introductory notes. HD 181327 is a 51.8 pc distant F5/6 main sequence star with a strong thermal IR excess ($L_{\text{IR}}/L_{\text{star}} = 0.25\%$). HD 181327 is a member of the β Pic moving group (canonical age ~ 12 Myr), and a likely common proper motion companion to the 0.1 pc distant HR 7329 (a likely coeval star with close brown dwarf companion; see Lowrance et al. 2005). These characteristics made it a high-priority target for prior *HST* high-contrast (coronagraphic) imaging investigations at both near-IR (*HST*/GO 10177, 10847) and visible-light (*HST*/GTO 9987) wavelengths. An optically bright light-scattering debris disk circumscribing HD 181327 ($F_{\text{disk}}/F_{\text{star}} = 0.17\%$ at $1.1 \mu\text{m}$) was first reported from *HST*/NICMOS and ACS observations by Schneider et al. (2006), while mid-IR imagery along with IR spectroscopy was discussed by Chen et al. (2008). Together, those observations revealed the existence of a relatively narrow ring-like debris belt responsible for $\sim 70\%$ of the total disk flux, peaking in SB at a stellocentric distance of ≈ 86 AU. The $\sim 31^\circ 7'$ from face-on inclined debris system appeared to be largely cleared of light-scattering material closely interior to the bright ring as determined from the NICMOS images

and exterior, surrounded by a much larger, asymmetric halo of diffusely scattering low-SB particles further from the star as detected in deeper ACS optical imaging.

Observations and PSF subtraction. Our STIS GO 12228 6R/PSFTSC observations of HD 181327 are detailed in Section 5.5 as archetypical of the observation and data reduction/processing methodologies as applied to all GO 12228 targets. In particular, for HD 181327, an ensemble of 168 images targeting the inner regions of the debris ring to a WedgeA-0.6 limiting stellocentric angle of $r = 0''.3$, and 64 deeper WedgeA-1.0 images to probe the outer regions of the debris system, were obtained at six different spacecraft roll (field orientation) angles with up to ~ 12.8 ks of total integration time in most pixels (see Table 3 for details). An intermediate set of 12 “visit level” combined images (derived from multiple images, separately at each wedge position in each orbit) were created; see Figures 3 and 4. Though there is some (anticipated) small dispersion in the quality (occurrence and amplitude of PSF-subtraction residuals) between these images in the regions unaffected by the *HST*/OTA diffraction spikes and the STIS occulting wedges (e.g., qualitatively compare the “worst case” scenario to the “best case” scenario of Figures 3(A) and (E), respectively) none were excluded as outliers in subsequent multi-roll image combination. An AQ data image (with error estimate propagation), combining all 12 astrometrically registered, artifact-masked, PSF-subtracted visit-level images,

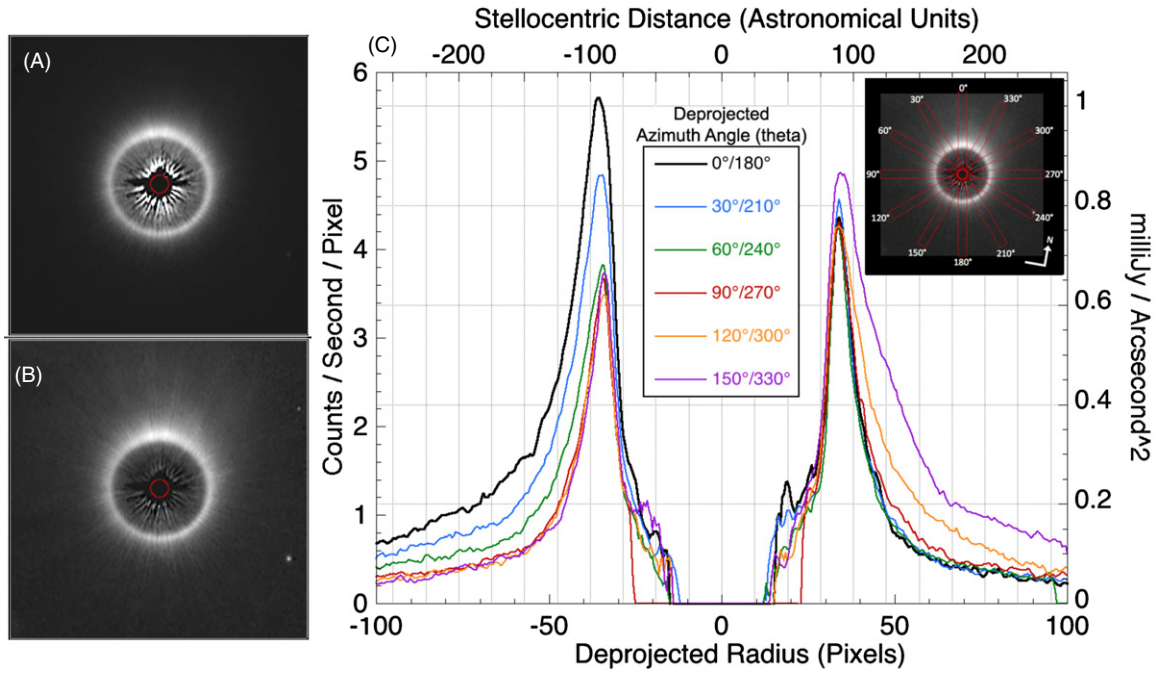


Figure 35. Face-on deprojected images of the HR 181327 debris ring (A) before and (B) after compensating for the r^{-2} dilution of the stellar radiation field (linear display stretches 0–6 counts s^{-1} pixel $^{-1}$; red circle $r = 0.3$). (C) 0.5 wide radial r^{-2} compensated brightness profile “cuts” in 30° deprojected azimuthal increments from $\theta = 0^\circ$ (direction of peak brightness along the ring). The physical units are based on *HST*/STIS photometric calibration per STScI MAST/OPUS pipeline, SIMBAD parallax, and STIS scale per instrument handbook; see Table 4.

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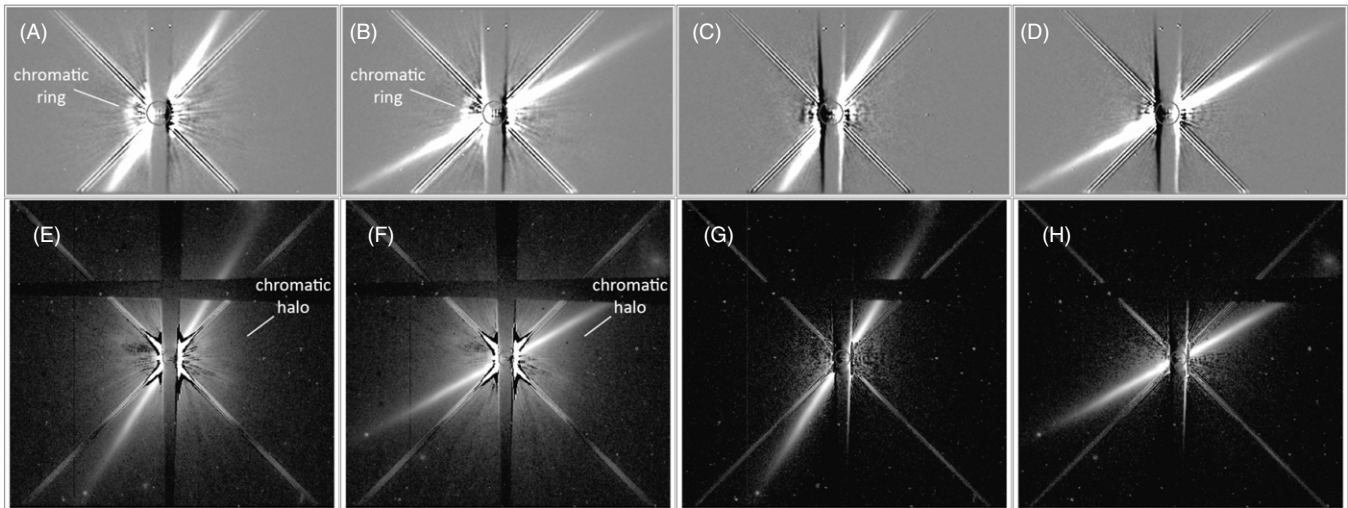


Figure 36. (Top) WedgeA-0.6 and (bottom) WedgeA-1.0 PSF-subtracted images of the HD Au Mic CD disk (V36 and V38 progressing left to right) with the HD 191847 PSF template star (left four panels) and mitigating residual chromatic effects using the disk-obscured V35 image of AU Mic itself as its own PSF template.

each oriented in the celestial (“north up”) frame, was then created with up to ~ 12.8 ks of total integration time in most pixels. The HD 181327 AQ image discussed here is presented in Figure 33.

Principal results. The new *HST*/STIS 6R-PSFTSC imaging of the HD 181327 debris system improves dramatically over the previous NICMOS (Figure 1) and ACS (Figure 6) co-discovery imaging in the range of stellocentric regions of the disk previously imaged, enabling novel science. A newly discovered radial variation in the apparent scattering phase function may imply the presence of a yet unimaged planetary-mass perturber (Stark et al. 2014). The new 6R/PSFTSC imaging of the HD 181327 debris disk (Figures 33(A) and (B)) also reveals previously unseen substructures and brightness

asymmetries along the ring (Figure 33(C)). The bright narrow ring of dust at $r = 1''.86$ (88.5 AU projected distance) exhibits non-axisymmetric SB asymmetries (Figure 33(C)) that cannot readily be explained by simple directionally preferential (e.g., Henyey & Greenstein 1941 (H-G)) scattering by the disk grains alone (Stark et al. *ibid*).

Morphologically, the HD 181327 debris disk appears as an asymmetrically bright narrow elliptical ring cleared of scattering material in the interior and exterior surrounded by a larger, fainter diffuse region. Assuming intrinsic circular symmetry, after elliptical isophote fitting of the flux density in the ring, and finding a celestial P.A. of $102^\circ \pm 4^\circ$ for the major axis of the bright ring inclined $30.1^\circ \pm 1.2^\circ$ from face-on, we deproject the HD 181327 disk to a face-on viewing geometry

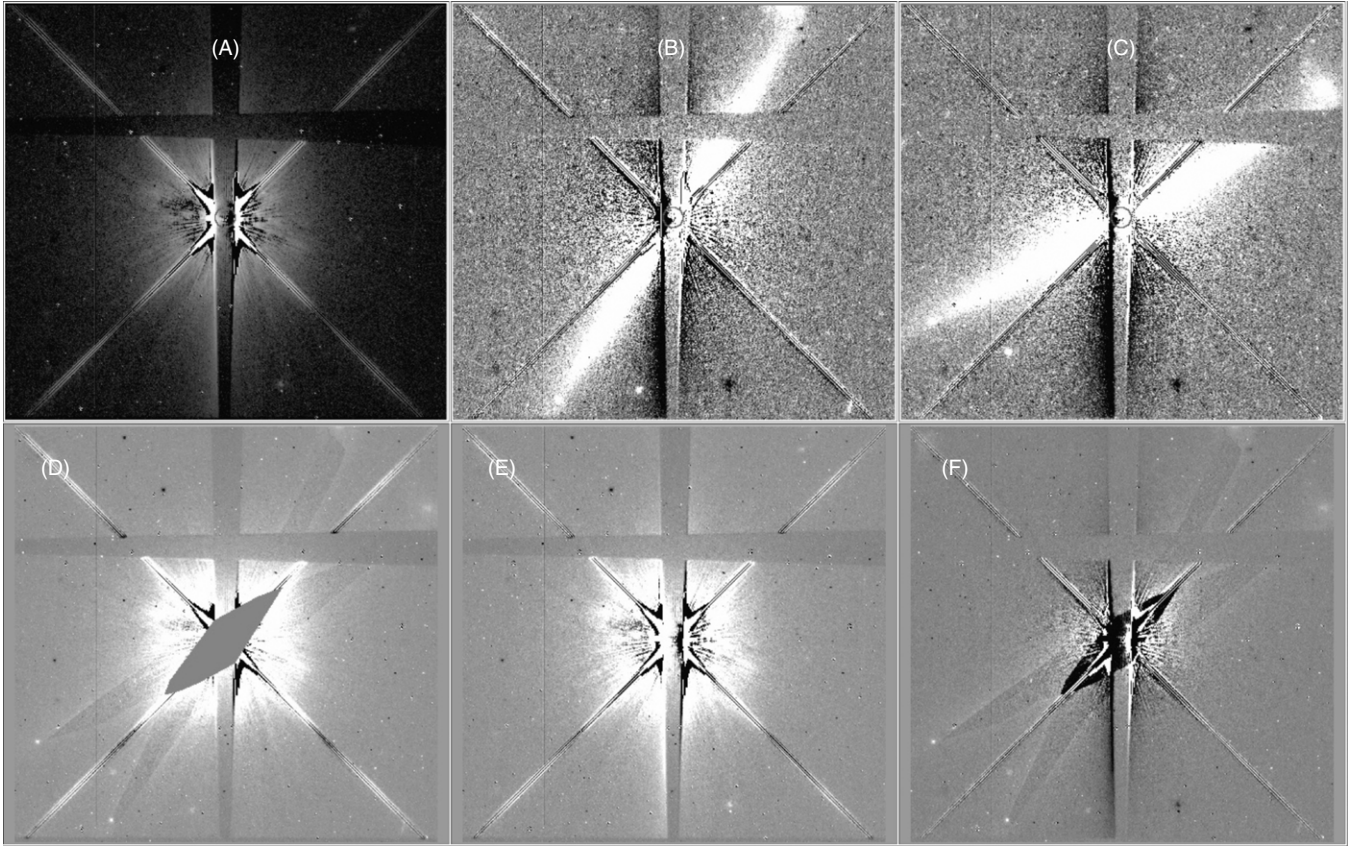


Figure 37. (A) Chromatic halo from disk-obscured V35 image minus V37 PSF template HD 19847; same stretch as Figures 36(E)–(H). ((B) and (C)) The V36 and V38 images, respectively, minus the disk-obscured V35 image stretched ± 0.02 counts $\text{s}^{-1} \text{pixel}^{-1}$. A small amount of lux from the disk in the V35 template oversubtracts the sky background close to the wedge edges (especially where the wedge taper is narrowest) but has little effect in the off-rolled plane of the disk itself. (D) Disk-masked image combination by averaging of the V36 and V38 images in the SIAF frame reproduces the halo as a rotationally invariant artifact. (E) Same as panel (A), but stretched as panel (D). (F) Difference demonstration. Panels (D) and (E) eliminates nearly all of the chromatic halo light from the disk-obscured image.

as shown in Figure 34 with a horizontal debris ring major axis. Photometric SB contours external to the bright ring deviate from perfect ellipses (circles in de-projection) and exhibit some azimuthal “skewing” with stellocentric distance about the presumed axially concentric bright debris ring.

The most immediately noticeable (and striking) features about the bright ring itself are (1) the narrowness of the ring; (2) the steeply declining SB gradients exterior, and particularly interior to the radius of (deprojected) peak SB; with (3) a dearth of dust-scattered starlight in the interior region. Additionally, (4) the diametrically opposed sides of the ring about the minor axis (vertical in Figure 34) at mirror-symmetric deprojected azimuthal phases angles (theta in Figures 34(B) and (C); theta = 0° coincident with the azimuth angle of the peak SB along the ring), are of significantly different brightness; (5) with a minimum in the SB around the ring not coincident with deprojected theta = 180° (i.e., not H-G scattering); (6) is also not diametrically opposed to the direction of the “extension” of the outer disk halo to the north.

Compensating for the r^{-2} decline of the stellar radiation field with increasing stellocentric distance, we transform the deprojected SB image of the debris ring (Figure 35(A)) into a proxy for the product of the optical depth and the scattering phase function of light-scattering particles (Figure 35(B)). The radial sharpness of this approximation for the radial dust density distribution, its abrupt clearing interior to the ring peak radius, and axial asymmetry with mirror-symmetric theta angles are illustrated (and readily apparent) in Figure 35(C).

The combination of previously unresolved substructures in the bi-axially asymmetric HD 181327 $r = 88$ AU debris ring, with a very narrow width, sharp inner edge, little or no light-scattering material interior, and large, skewed, diffuse outer halo all implicate the possible presence of yet unimagined planetary-mass perturbers.

A.10. AU MIC

Introductory notes. AU Mic is a ~ 12 Myr old member of the β Pic moving group and the spectrally latest (M0V) and lowest-mass star for which a CS debris disk has been directly imaged. Due to its only 9.9 pc fortuitous proximity to Earth, large angular extent, and $F_{\text{disk}}/F_{\text{star}} = 0.2\%$ flux ratio with natural contrast enhancement owed to the edge-on viewing geometry of its debris system, the AU Mic disk was discovered with ground-based optical coronagraphy by Kalas et al. (2004). Follow-up Keck II NIRC2/H-band images (Liu 2004) revealed not only the disk itself, but suggested some substructures along the disk major axis at stellocentric distances beyond about 25 AU, suggestive (perhaps) of being influenced by co-orbiting, unseen planets. Subsequent *HST*/ACS imaging by Krist et al. (2005), in three broad optical bands, revealed (diffraction polluted) light from the disk as close as $0''.75$ (interior to the $r = 0''.9$ focal plane mask at the *HST* spherically aberrated focal plane) from the central star (e.g., see Figure 1, top left panel, reproduced from Krist et al. 2005) and cleanly traced the disk along its edge-on major axis from approximately $r = 2''$ – $15''$. The ACS F606W ($0.61 \mu\text{m}$) image reproduces some of the substructures suggested in the

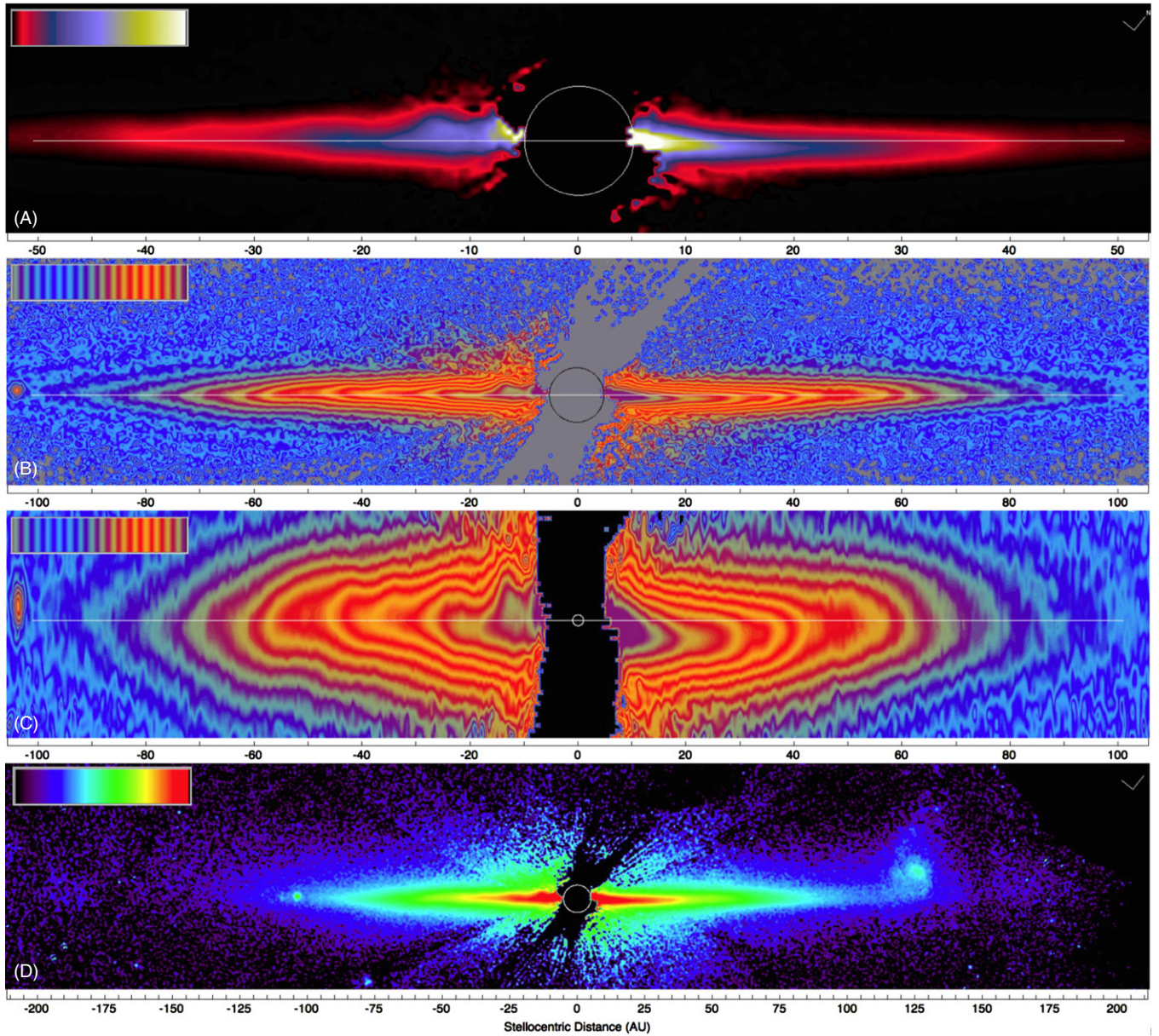


Figure 38. STIS five-roll (10 image) combined PSF-subtracted images of the AU Mic CS disk (celestial north $38^{\circ}7'$ CW of image vertical). (A) “Inner” region $5 \leq d \leq 50$ AU in $10''.62 \times 2''.67$ with a linear display stretch $0\text{--}15 \text{ counts s}^{-1} \text{ pixel}^{-1}$ (0 to approximately $2650 \mu\text{Jy arcsec}^{-2} \text{ pixel}^{-1}$). ((B) and (C)): isophotal contour images $5 \leq d \leq 100$ AU in $21''.24 \times 5''.34$ FOV. (C) With four times vertical scale expansion. (D) Full spatial extent in $42''.64 \times 16''.66$ FOV. ((B)–(D)) \log_{10} display stretch $[-3.0]$ to $[+1.0] \text{ counts s}^{-1} \text{ pixel}^{-1} \{\text{dex}\}$ (approximately 0.177 to $1770 \mu\text{Jy arcsec}^{-2} \text{ pixel}^{-1}$). The white line (panels (A)–(C)) is disk mid-plane-defined by panel (B) isophotes from $50 \leq d \leq 100$ AU. The central circle $r = 0''.5$, except panel (B), marking the star position. A background (non-commoving) star and galaxy, at this epoch, flank the opposite side of the disk along its major axis.

(A color version of this figure is available in the online journal.)

Keck II image in this radial range, though with a “relatively large discrepancy in position of feature C {that} . . . may be due to subtraction errors in either or both data sets”, (p. 1012). Krist et al. (2005) state that a three-component fit to the disk radial SB profile along the disk major axis decreases in steepness (in two breaks) with stellocentric distance. Subsequently, the disk has been modeled using multi-wavelength and polarized light data (Graham et al. 2007; Fitzgerald et al. 2007) and the signatures of grain growth probed.

Observations and PSF subtraction—details. STIS 6R/PSFTSC observations of AU Mic and its contemporaneously observed PSF star, HD 191847, were conducted at two epochs nearly a year apart: 2010 August 09 (V35–V38) and 2011 July

16 (V31–V34). In V35, the edge-on debris disk was aligned with and obscured by STIS occulting WedgeA. Thus, only five rolls in the two epochs combined contribute to the imaging of the disk along its mid-plane. Despite the excellent $\Delta(B - V) = -0.02$ color match of its PSF template star, PSF-subtractions from all six visits revealed bright residuals in the form of a chromatic ring at $r = 0''.8$ and a large diffuse chromatic halo beyond. These residuals are illustrated in Figure 36 with two representative visits each from WedgeA-0.6 (panels (A) and (B)) and WedgeA-1.0 (panels (E) and (F)). These residuals mostly arise from a mis-match in target:template photospheric SED on the red end of the very broad, unfiltered, STIS bandpass of greater impact for this reddest ($B - V = +1.46$) of all our disk hosting

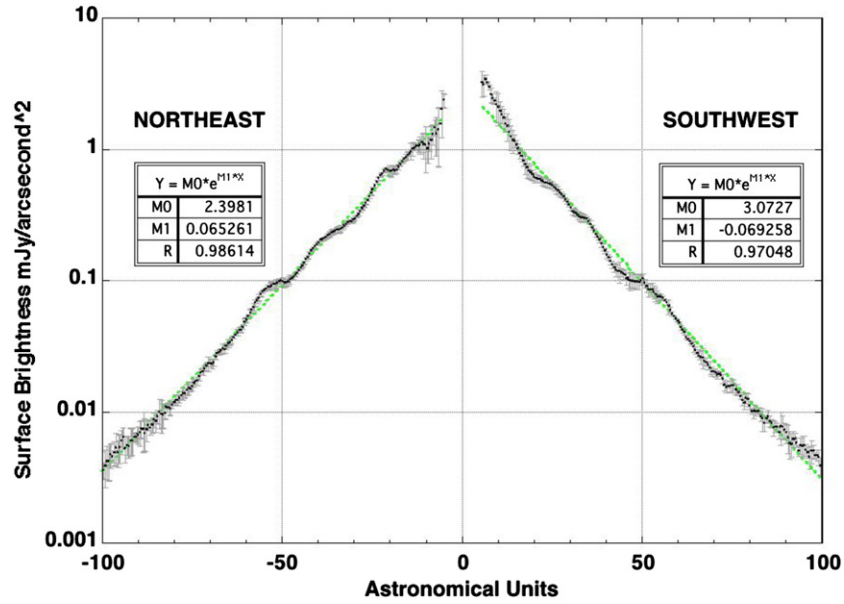


Figure 39. Radial SB profile of the AU Mic CS debris disk, corresponding to Figure 38(B), measured in a 50 mas wide central strip along the disk mid-plane ± 100 AU from the central star, separately fit on both sides of the star to single exponentials. The 1σ error bars on the observed data are computed from the dispersion in measures from the individual PSF-subtracted images about the roll-combined median image (up to 10 input images for any single pixel) and do not include likely less than a few percent uncertainty in the absolute flux density calibration. Statistically significant deviations on both sides of the star, both asymmetric (e.g., rise at ~ 20 AU to the NE and depression same distance to the SW) and symmetric (e.g., the significant “break” on both sides) about the minor axis, and other deviations elsewhere in the inner (< 50 AU) part of the disk.

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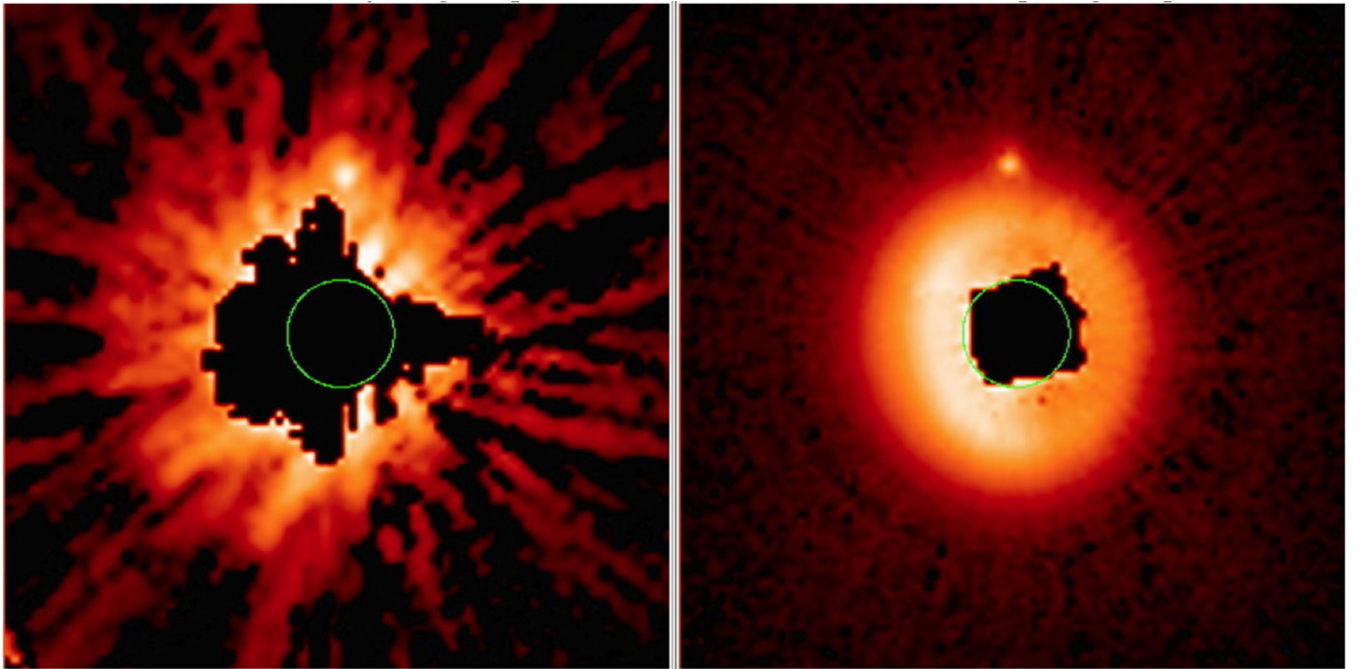


Figure 40. Left: NICMOS $1.1 \mu\text{m}$ two-roll scattered-light discovery image of the MP Mus CS disk (as discussed by Cortes et al, 2009). Right: STIS 6R/PSFTSC AQ image reproduced at the same spatial scale and orientation (north up, east left). Both \log_{10} display stretch from $[-2]$ to $[+1]$ {dex} counts $\text{s}^{-1} \text{pixel}^{-1}$. NICMOS 1 count $\text{s}^{-1} \text{pixel}^{-1} = 218 \mu\text{Jy arcsec}^{-2}$, STIS 1 count $\text{s}^{-1} \text{pixel}^{-1} = 177 \mu\text{Jy arcsec}^{-2}$. The green circle is $r = 0.41$. The prominent point-like object to the north of MP Mus, superimposed on the outer periphery of the disk, was unequivocally found to be a background star from differential proper motion measures between the NICMOS (circa 2007) and STIS (circa 2011) epochs, as previously suggested comparatively from VLT/NACO imaging (circa 2003) of the stars (but not the disk).

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stars. In panels (C)–(F) we, respectively, use the disk-obscured images of AU Mic itself obtained in V35 as its own PSF template, in which the self-mitigating effects to chromatism are apparent.

Given the above, we used the V35 disk-occulted images of AU Mic itself as its own self-referential PSF template that,

fortuitously, exhibited no significant breathing excursions with respect to the non-contemporaneous disk imaging visits 31–34. In doing so, we first tested whether or not this could oversubtract with significance any disk flux from out-of-plane light due to the disk in the template image possibly unobscured by Wedge A. In Figure 37(A), we show the large chromatic halo that

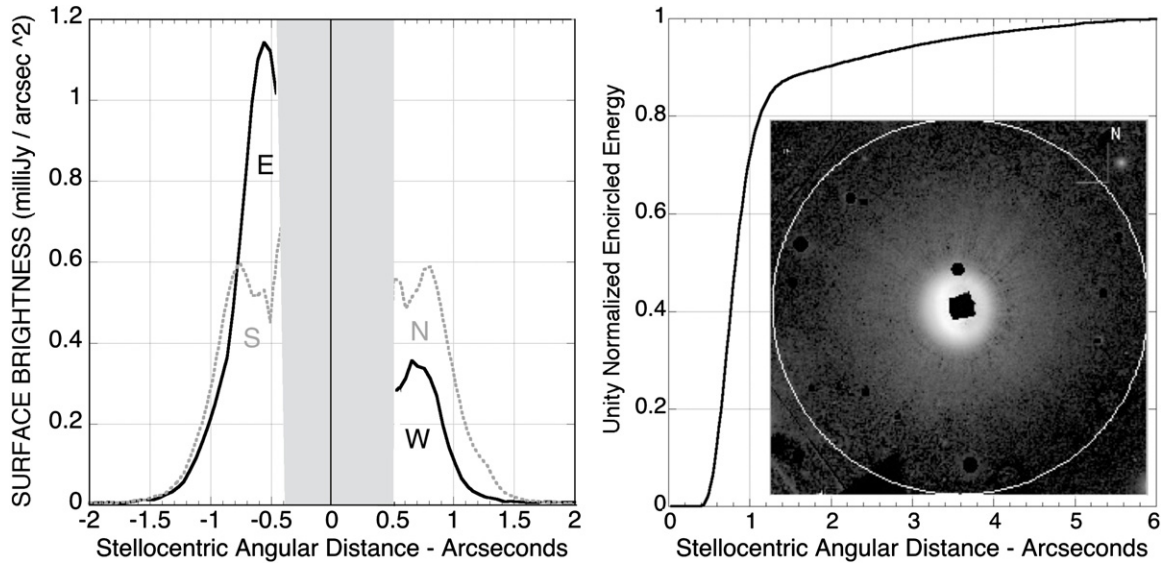


Figure 41. Left: radial SB profiles of the MP Mus CS disk, measured in two pixel wide strips along its major (dotted gray) and minor (solid black) axes, $\pm 2''$ from the star. Right: encircled energy plot to $r \leq 6''$ corresponding to the AQ image inset (white circle radius = $6''.09$) with stellar background sources masked reveals the MP Mus CS disk is surrounded by large halo of low-SB light-scattering dust. The inset image is displayed with a \log_{10} stretch from $[-3]$ to $[+1]$ {dex} counts s^{-1} pixel $^{-1}$.

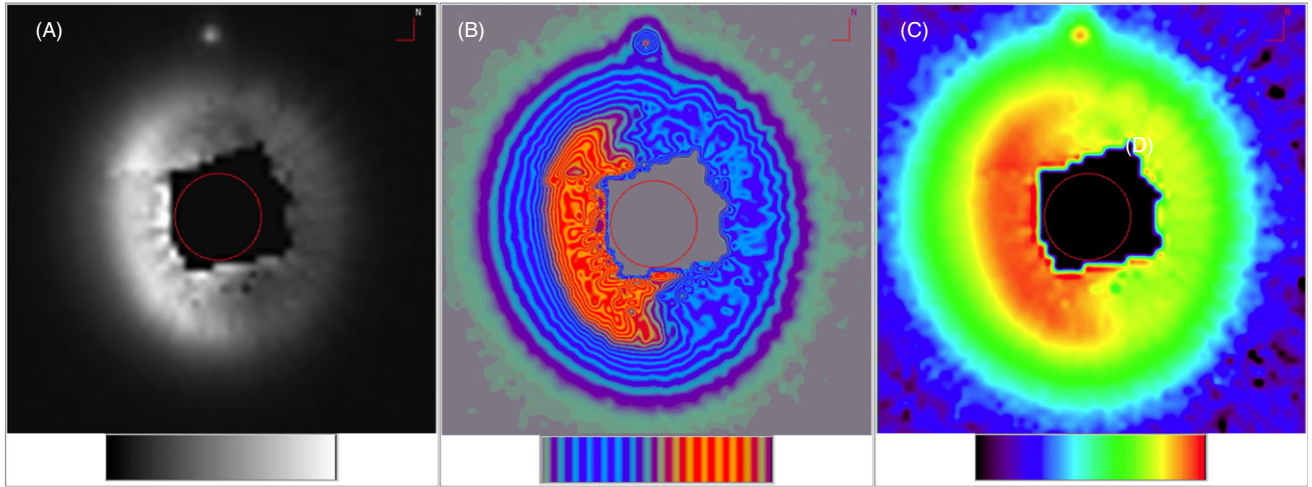


Figure 42. MP Mus disk. Field: $3''.05 \times 3''.05$ (60×60 pixels) centered on occulted star (celestial north up). Left: linear display image from -0.5 (black) to $+8.0$ (white) counts s^{-1} pixel $^{-1}$. Middle: linear display isophote map from 0 to 8 counts s^{-1} pixel $^{-1}$ with 3.125% intensity contour intervals. Right: \log_{10} image display from -2.0 to $+1.0$ {dex} counts s^{-1} pixel $^{-1}$. Red circle: $r = 6$ pixels ($0''.305$). N.B.: the point-object to the north is (unfortunately) unequivocally a background star from non-common proper motions as established with 2006 epoch NICMOS disk discovery images.

(A color version of this figure is available in the online journal.)

emerges when subtracting the HD 191847 PSF star (V37) from the disk-obscured V35 image (similar to panels (E) and (F) in Figure 36 with the disk unobscured by Wedge A). No negative oversubtraction due to light from out of the disk plane, along the edges of the A wedge, is seen at this same display stretch in which the images of the disk in Figures 36(E)–(H) are shown.

In Figures 37(B) and (C), we subtract the V35 AU Mic template from the V36 and V38 disk images as in Figures 36(G) and (H), but here shown with a very hard bi-directional linear stretch of ± 0.02 counts s^{-1} pixel $^{-1}$, from which we can see at a very low level, and quantify (nearly inconsequential) out-of-plane oversubtraction of background light far (in the vertical plane) from the disk itself. This is very significantly smaller than the level of the chromatic residuals using the HD 191847 PSF star, and thus we proceed using the AU Mic V35 image as a PSF-subtraction template.

Cautiously, as a further test illustrated in Figure 37, we confirm the chromatic halo that appears when subtracting the

PSF star is indeed an artifact (and not due to an extended disk halo) by its rotational invariance in V36-PSF and V38-PSF images (as well as in the V35-PSF image where the disk is behind the wedge). Here we combine in the detector (SIAF) frame the two PSF-subtracted images shown in panels (B) and (C) (each containing the disk visible at two different orientations, not rotated to a common celestial orientation) after separately digitally masking the visible portions of disk. This gives a slightly imperfect image of only the halo as shown in panel (D). In detail, (1) regions unmasked in both images, where the disk does not appear, were averaged; (2) regions masked in one image where the disk appears but not in the other were rejected; and (3) regions where the disk appears in both images (i.e., in uniform gray near the image center) were zeroed. The resulting combined image in panel (D) “eliminates” the disk (based upon the empirical mask of where the disk appears in each image) and shows only the posited rotationally invariant chromatic “halo.” Panel (E) then shows the

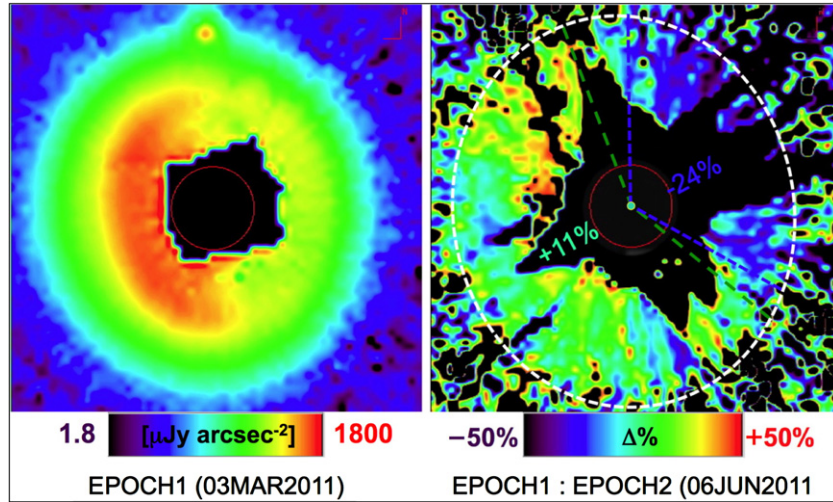


Figure 43. Left: \log_{10} -scaled SB image of the MP Mus disk using all data combined, normalized to the first epoch (visit A1) stellar brightness. Right: spatially resolved, temporal variability in the MP Mus CS disk relative to the first epoch image. The median changes in SB in the two diametrically opposed regions considered within the ellipse representing the outer boundary of the disk are noted. The central regions in black are excluded due to mutually exclusive spatial coverage between the two epochs, and in the image periphery due to insufficient S/N from rapidly declining disk flux density. Red circle radius (both panels), $r = 0''.35$.

(A color version of this figure is available in the online journal.)

chromatic halo in the V35 image of AU Mic (after subtracting the HD 191847 color mismatched template), as shown in panel (A) but here in a symmetric \pm linear display stretch. If the halo seen in the (panel (D)) two-visit combined difference image is a rotationally invariant artifact, subtracting the V35 (panel (E)) image difference should (in double differencing) null the halo. This is seen to be the case as shown in panel (F) with no significant azimuthal substructure (just noise) after differencing, indicative of a rotationally invariant image artifact in the two-roll detector-frame-combined image in panel (D); i.e., the “halo” is indeed an optical artifact, and not intrinsic to the CS environment. Therefore, the AQ data images we present and discuss here all use the V35 disk-occulted image of AU Mic itself as its own self-referential PSF template.

Principal results. The AU MIC edge-on disk is cleanly imaged with STIS 5R/PSFTSC using the disk-obscured AU Mic image as its own PSF template to a stellocentric inner working distance of 5 AU along the disk mid-plane. Here, we photometrically define the disk mid-plane (illustrated by the horizontal white line in panels (A)–(C)) by fitting the central ridge of isophotal SB maximum at stellocentric distances 50–100 AU on either side of and passing through the star (finding a celestial P.A. of $37.8 \pm 0.2^\circ$).

In the inner ($r < 50$ AU) region of the disk, the new image reveals the following.

1. A significant NW (brighter) to SE (fainter) brightness asymmetry between the two “sides” of the disk at equal stellocentric distances interior to ~ 15 –20 AU inner working distances (see Figure 38(A)) that had not been cleanly probed with earlier scattered-light imaging.
2. A prominent brightness enhancement (“bump”) on the SE side of the star “above” (to the NE) of the disk mid-plane at 13 AU (see also panel (A)). We have readily confirmed the “bump” as intrinsic to the disk (not due to background contamination) as established by its common proper motion with AU Mic itself ($+380, -361$ mas yr $^{-1}$) over the two observational epochs nearly a year apart.
3. Significant out-of-plane asymmetries on the NW side of the disk “below” (to the SW) of the mid-plane, opposite the SE-

side “bump,” with a “warp” below the disk mid-plane to a stellocentric distance of ~ 45 AU *only* on the NW side of the disk. This is shown in the isophotal contour maps in panel (B) and with a four times vertical (only) scale expansion in panel (C).

The outer regions of the disk are sensitively detected to ~ 130 AU on both sides of the star (panel (D)). In the outer regions of the disk image, in the SE a background star (confirmed from the two observational epochs as not co-moving with the disk) coincidentally falls on the disk mid-plane at a disk-projected distance of 105 AU from the star. On the NW side, at 125 AU a background galaxy just to the NE of and blending with the disk mid-plane pollutes the image.

The disk SB along the mid-plane, as measured from earlier ACS images, had been fit and discussed by Krist et al. (2005) with a three-component power law of increasing index in segregated zones with increasing distance from the star. Here, without invoking physical arguments, we separately fit the STIS-derived SB profile on both sides of the disk each to a simple exponential. We did so to better examine the apparent local deviations from a smoothly continuous radial SB profile with stellocentric distance to identify small spatial scale substructures in the disk above the local measurement uncertainties. The as-measured profiles, and the two resulting fits, are illustrated in Figure 39.

In the vertical dimension, the aforementioned “bump” at 13 AU to the north of the mid-plane on the east side of the disk is the most prominent feature. Lower SB material “above” the disk mid-plane at and beyond the radial distance of the bump on the east side of the disk is seen; e.g. in green in Figure 38(D) from -45 AU to -10 AU. This scattered light in excess of the exponential fit does not appear either south of the disk on the east side or north of the disk on the (“no bump”) west side—though visibility there is precluded at $r \lesssim 12$ AU due to the remaining obscuration by the STIS occulting wedge. The “bump” and excess to the NE of the star and “warp” in the photometric mid-plane of the disk to the diametrically opposed SW *may* have a common causality with dust in the inner part of the system orbiting non-coplanar (slightly inclined) to the outer

part of the disk, e.g., like β Pictoris, possibly (here) to a yet unseen inner planet.

A.11. MP MUS (PDS 66)

Introductory notes. Unlike the other GO 12228 targets, MP Mus (PDS 66, Hen 892) is a pre-main sequence, classical T-Tauri star (cTTS) hosting a late stage starlight-scattering protoplanetary/transition disk discovered by Cortes et al. (2009) with *HST*/NICMOS as part of a *Spitzer*/FEPS-selected coronagraph imaging survey of sunlike stars. Its age was estimated by Cortes et al. (2009) as ~ 13 Myr, as a member of the Sco Cen OB association (Lower Centaurus-Crux moving group) and by fitting stellar evolutionary tracks. Despite its advanced age for a cTTS (having been described in the literature as “mature”), CS material is still accreting onto the star, as evidenced by variability in the thermal SED and other accretion signatures typical for much younger stars. MP Mus is known to be optically variable with a 5.75 day period and peak-to-peak amplitude of ≈ 0.1 mag in *V* band (see Batalha et al. 1998); the GCVS 2007–2011 gives V_{mag} range 10.30–10.42. In combination with ACTA interferometric observations in four continuum bands from 3 mm to 6 cm, and an IR SED derived from Two-Micron All-Sky Survey, *IRAS*, and *Spitzer* IRAC and MIPS photometry and IRS spectroscopy, the grain and physical properties of the disk were partially constrained by Cortes et al. (2009), with its geometry informed by the NICMOS imaging to assess the potential for planet formation in the disk. Kastner et al. (2010) detected molecular gas (CO) orbiting MP Mus as had been found previously for the prototypical “old” cTTS TW Hya, and for V 4046 Sgr (both, like MP Mus, within 100 pc of the Earth).

The primary focus of the GO 12228 investigation and its target sample is on CS debris disks. MP Mus, however, was importantly included as a “mature” protoplanetary/transition disk to directly bridge those data spanning two decades in stellar ages to a nearby ($d < 100$ pc) protoplanetary/transition disk system, spatially resolvable with *HST*, at this earlier phase of disk (and exoplanetary system) evolution, coeval with the youngest of our debris disk targets.

Observations and PSF subtraction. STIS 6/PSFTSC observations of MP Mus and its PSF template star executed per the observation plan at two epochs three months apart with no issues on 2011 March 3 (VA1–VA4) and 2011 June 6 (VA5–VA8). Unlike our other targets, in the case of MP Mus and its PSF star, our observing strategy included, by design, only WedgeA-0.6 imaging. The combination of the relative faintness of the star ($V = 10.44$) and efficiently scattering disk dust particles, as known from the NICMOS imaging, allowed us to expose more deeply with WedgeA-0.6 (only) without risk of saturation close to the star. This, and the anticipated smaller angular size of this 86 pc distant protoplanetary disk, obviated the need for WedgeA-1.0 imaging (and the additional longer overheads incurred from its utilization).

MP Mus’s intrinsic optical variability informed of the possibility of significant/detectable stellar brightness changes at some light curve phases even over the ≈ 7 hr period of each of our two sets of visits, in addition to the three-month timescale between the two observational epochs. Hence, in multi-roll combination, the contributing single-visit PSF-subtracted images were re-normalized to the stellar brightness in visit A1 (when the star was found to be brightest); i.e., so inter-visit PSF-subtracted images of the disk are constant in $F_{\text{disk}}/F_{\text{star}}$ as F_{star} varies.²¹

²¹ MP Mus is optically variable with $\Delta V \sim 0.1$ mag and a 5.75 day period. In combining and comparing the image data obtained at the two observational

The inter-visit brightness of the occulted star was determined by ratioing the flux from the unocculted OTA diffraction spikes, before PSF subtraction, beyond the WedgeA-0.6 mask. In detail, we found the following relative brightnesses by visit: (epoch 1): $A1 = \times 1.0000$, $A2 = \times 1.0041$, $A4 = \times 1.0081$; (epoch 2): $A5 = \times 1.0784$, $A6 = \times 1.0714$, $A8 = \times 1.0554$.

In the thusly weighted six-roll combined image, the disk is very cleanly resolved to an instrumentally asymmetric IWA of as close as $0''.305$ from the disk-hosting star. Additionally, AQ images partitioned from the first and second observational epochs were separately created to permit the investigation of spatially resolved temporal changes in the disk surface illumination over time (and thus SB) in regions commonly sampled at both epochs.

Principal results. The improvement in image quality from the NICMOS discovery image reproduced in Figure 40(A) to the STIS 6R/PSFTSC AQ image in panel (B) is striking, and bespeaks the utility of this observational technique not only for studying optically thin debris disks, but for optically thick protoplanetary disks as well. From this image, we have, more accurately than from the NICMOS image, redetermined the principal characterizing attributes of the disk.

The disk geometry was explored by elliptical isophote fitting in the stellocentric region $r \geq 0''.863$ (17 STIS pixels) where the disk isophotes are elliptically concentric (see Figure 42(B)). At this inner distance the disk brightness, in instrumental units, is ~ 2.0 counts s^{-1} pixel $^{-1}$ on the eastern semi-minor axis and ~ 1.5 counts s^{-1} pixel $^{-1}$ on the western semi-minor axis. With an assumption of apparent ellipticity due to a circularly symmetric disk seen in the inclined projection, the disk isophotes were found to be offset from the location of the star (by the “X marks the spot” method, Section 5.4) by $+0.66$ (farther from the star) ± 0.2 pixels along the brighter disk semi-minor axis. This photometric (isophotal) offset does not necessarily imply a physical offset in the particle density distribution as the large scattering fraction ($F_{\text{disk}}/F_{\text{star}}$, see below) is consistent with a disk that is at least partially optically thick and non-isotropic scattering may account for this offset (as later modeling can explore).

We revisited the disk geometry, applying this observed isophotal stellocentric offset to define the disk center. We find that (1) the celestial orientation (P.A.) of the disk major axis is $10^\circ \pm 2^\circ$ east of north; (2) the minor:major axial ratio is 0.889 ± 0.026 , implying (3) an inclination of $27.3 \pm 3.3^\circ$ from a face-on viewing geometry (in statistical agreement with, but improving upon the earlier determination of $i = 32^\circ \pm 5^\circ$ by Cortes et al. (2009) from the NICMOS discovery image).

Photometrically, the disk is truncated along the major axis at $r \sim 1''.3$ (see axial radial profiles in Figure 41, left panel), which may be roughly considered the outer edge of the main disk. However, only 90% of the total starlight scattered by the MP Mus CS disk (excluding the central region obscured by the STIS wedge in the STIS six-roll combine image) is within $2''.0$ of the star (see encircled energy plot in Figure 41, right panel). A diffuse low-SB halo surrounds the main disk to $r \leq 6''$ (see Figure 41 right panel inset). The flux density of disk measured in an annulus $0''.36 \leq r \leq 6''.09$ (120 STIS pixels), normalized to the temporally variable stellar brightness in Visit A1, is 1.86 mJy. Beyond this stellocentric distance, any posited additional contribution to total disk flux density is at or below the

epochs, the disk flux density was first re-normalized on a per-visit basis to a compensate for a $\pm 3.5\%$ stellar in stellar brightness between the two observational epochs.

level of the noise from the sky-plus-instrumental background. *N.B.*: a similar measure in the F110W band found a total disk $1.1\ \mu\text{m}$ disk flux density of $2.4 \pm 0.4\ \text{mJy}$ (as reported by Cortes et al. 2009). Given the V magnitude of the star, we find the total optical disk plus diffuse halo scattering fraction, $F_{\text{disk}}/F_{\text{star}}$ ($r < 6''.09$) $\sim 0.73\%$. To a conservatively inclusive distance beyond the outer boundary of the optically thicker main disk of $r = 2''.5$, $F_{\text{disk}}/F_{\text{star}}$ ($r < 2''.5$) $\sim 0.68\%$.

The detailed spatially resolved morphology of the disk is better revealed in a series of three representations of the 6R/PSFTSC AQ image as shown in Figure 42. The disk apparently exhibits strong directionally preferential scattering centered at a celestial P.A. $\sim 92^\circ$, which is offset from the disk's brighter semi-minor axis by -8° ; e.g., at $r = 0''.6$ along the disk minor axis, the disk SB is ~ 3.5 times brighter on the eastern side of the disk compared to its diametrically opposed western side. The 6R image is composed of images from two epochs with strong evidence, discussed below, for time-variable illumination resulting in spatially localized fluctuations of disk SB at levels of tens of percent that must be considered in analytic modeling of the disk.

Our STIS six-roll PSFTSC observations were obtained at two epochs, in two sets of three-target orbits three months apart. In the spatial regions commonly sampled at both epochs, a significant change in the disk SB with stellocentric azimuth angle was seen, separate from an anticipated correlation of the global disk brightness with stellar variability. By taking the ratio of the first epoch to (globally renormalized) the second epoch AQ images, a strong signature of azimuthally differentiated SB variability unassociated with the (global) stellar variability is seen. Specifically, as illustrated in Figure 43, in an annulus from $0''.6 \leq r < 1''.4$ in regions commonly sampled at both epochs, the first epoch to second epoch globally renormalized ratio of median disk SB varied significantly on opposite sides of the disk. In particular, in sector from approximately celestial P.A. $\approx 20^\circ$ – 230° (roughly centered on the presumed forward-scattering direction to the SE of the star; green dashed lines in Figure 43) the median disk SB increased by $+11\%$ relative to the star. Contemporaneously, in near diametric opposition from celestial P.A. approximately 240° – 360° (blue dashed lines) to the NW, the median disk SB declined by -24% relative to the star.

Given the orbital timescales, these changes cannot be from a change in the surface density distribution of light-scattering grains. We interpret this as possibly being due to time-variable shadowing (or “beaming”) of starlight to the outer disk beyond, either by opaque material orbiting close to the star, or the “puffing” (change in vertical height) of a unseen inner cavity wall, interior to the inner working distance ($\sim 35\ \text{AU}$) that could probed by the STIS coronagraph. Alternatively, some variation in the illumination pattern onto the disk might arise from differentiated brightness or localized on the stellar photosphere from accretion “hotspots.” These possibilities might be arbitrated with properly sampled, time-resolved imaging on the stellar rotation versus longer timescales, which our two-epoch data inform, but are insufficient to differentiate between the possibilities.

REFERENCES

- Albertsson, T., Semenov, D., & Henning, Th. 2014, *ApJ*, **784**, 39
- Ardila, D. R., Golimowski, D. A., Krist, J. E., et al. 2004, *ApJL*, **617**, L147
- Ardila, D. R., Golimowski, D. A., Krist, J. E., et al. 2005, *ApJL*, **624**, L141
- Augereau, J. C., Nelson, R. P., Lagrange, A. M., Papaloizou, J. C. B., & Mouillet, D. 2001, *A&A*, **370**, 447
- Backman, D. E., & Paresce, F. 1993, in *Protostars and Planets III*, ed. E. H. Levy & J. I. Lunine (Tucson, AZ: Univ. of Arizona Press), 1253
- Batalha, C. C., Quast, G. R., Torres, C. A. O., et al. 1998, *A&AS*, **128**, 561
- Beust, H., Augereau, J.-C., Bonsor, A., et al. 2014, *A&A*, **561A**, 43
- Beuther, H., Klessen, R., & Henning, Th. 2014, in *Protostars and Planets VI*, ed. H. Beuther, R. Klessen, C. Dullemond, & Th. Henning (University of Arizona Space Science Series; Tucson, AZ: Univ. Arizona Press), in press
- Boccaletti, A., Augereau, J.-C., Lagrange, A.-M., et al. 2012, *A&A*, **544**, A85
- Boley, A. C., Payne, M. J., Corder, S., et al. 2012, *ApJL*, **750**, L21
- Booth, M., Kennedy, G., Sibthorpe, B., et al. 2013, *MNRAS*, **428**, 1263
- Botte, W. F., Durda, D. D., Nesvorný, D., et al. 2005, *Icar*, **179**, 63
- Buenzli, E., Thalmann, C., Vigan, A., et al. 2010, *A&A*, **524**, L1
- Carpenter, J. M., Wolf, S., Schreyer, K., Launhardt, R., & Henning, Th. 2005, *AJ*, **129**, 1049
- Chen, C. H., Sheehan, P., Watson, D. M., Manoj, P., & Najita, J. R. 2008, *ApJ*, **701**, 1367
- Chiang, E., Kite, E., Kalas, P., Graham, J. R., & Clampin, M. 2009, *ApJ*, **693**, 734
- Choquet, É., Pueyo, L., Hagan, J. B., et al. 2014, *Proc. SPIE*, **9143**, 914357
- Clampin, M., Krist, J. E., Ardila, D. R., et al. 2003, *AJ*, **126**, 385
- Corder, S., Carpenter, J. M., Sargent, A. I., Zauderer, B. A., et al. 2009, *ApJL*, **290**, L65
- Cortes, S. R., Meyer, M. R., Carpenter, J. M., et al. 2009, *ApJ*, **697**, 1305
- Currie, T., Rodigas, T. J., Debes, J., et al. 2012, *ApJ*, **757**, 28
- Debes, J. H., Perrin, M. D., & Schneider, G. 2013, *STScI Newsl.*, **30**, 24
- Debes, J. H., Weinberger, A. J., & Kuchner, M. J. 2009, *ApJ*, **702**, 318
- Debes, J. H., Weinberger, A. J., & Song, I. 2008, *ApJL*, **684**, L41
- Decin, G., Dominik, C., Malfait, K., Mayor, M., & Waelkens, C. 2000, *A&A*, **357**, 533
- Desidera, S., Covino, E., Messina, S., et al. 2011, *A&A*, **529**, A54
- Donaldson, J. K., Lebreton, J., Roberge, A., Augereau, J.-C., & Krivov, A. V. 2013, *ApJ*, **772**, 17
- Duchêne, G., Arriaga, P., Wyatt, M., et al. 2014, *ApJ*, **784**, 148
- Eiroa, C., Marshall, J. P., Mora, A., et al. 2013, *A&A*, **555**, 11
- Ertel, S., Wolf, S., Metchev, S., Schneider, G., et al. 2011, *A&A*, **533**, A132
- Ertel, S., Wolf, S., & Rodmann, J. 2012, *A&A*, **544**, A61
- Fedele, D., van den Ancker, M. E., Henning, Th., & Jayawardhana, J. M. 2010, *A&A*, **510**, 72
- Fitzgerald, M. P., Kalas, P. G., Duchene, G., Pinte, C., & Graham, J. R. 2007, *ApJ*, **670**, 536
- Golimowski, D. A., Ardila, D. R., Krist, J. E., et al. 2006, *AJ*, **131**, 3109
- Golimowski, D. A., Krist, J. E., Stapelfeldt, K. R., et al. 2011, *AJ*, **142**, 30
- Grady, C. A., Polomski, E. F., Henning, Th., et al. 2001, *AJ*, **122**, 3396
- Grady, C. A., Woodgate, B. E., Bowers, C. W., et al. 2005, *ApJ*, **630**, 958
- Grady, C. A., Woodgate, B., Bruhweiler, F. C., et al. 1999, *ApJL*, **523**, L151
- Graham, J. R., Kalas, P. G., & Matthews, B. C. 2007, *ApJ*, **654**, 595
- Greaves, J. S. 2012, in *From Atoms to Pebbles: Herschel's View of Star and Planet Formation*, ed. J.-C. Augereau (Grenoble: ESA), http://herschel.esac.esa.int/Publ/2012/Grenoble/23-0900_GreavesJ_DebrisDisks-HerschelOverview.pdf
- Greaves, J. S., Holland, W. S., Moriarty-Schieven, G., et al. 1998, *ApJ*, **506**, 133
- Greaves, J. S., Kennedy, G. M., Thureau, N., et al. 2014, *MNRAS*, **438**, L31
- Haisch, K. E., Jr., Lada, E. A., & Lada, C. J. 2001, *ApJL*, **553**, L153
- Heap, S. R., Lindler, D. J., Lanz, T. M., et al. 2000, *ApJ*, **539**, 435
- Heney, L. G., & Greenstein, L. J. 1941, *ApJ*, **93**, 70
- Hines, D. C., Schneider, G., Hollenbach, D., et al. 2007, *ApJL*, **671**, L165
- Hinkley, S., Oppenheimer, B. R., Soummer, R., et al. 2009, *ApJ*, **701**, 804
- Holland, W. S., Greaves, J. S., Zuckerman, B., et al. 1998, *Natur*, **392**, 788
- Jewitt, D. C., Weaver, H., Agarwal, J., Mutchler, M., & Drahos, M. 2010, *Natur*, **467**, 817
- Kalas, P. 2005, *ApJL*, **635**, L169
- Kalas, P., Duchene, G., Fitzgerald, M. P., & Graham, J. R. 2007a, *ApJL*, **671**, L161
- Kalas, P., Fitzgerald, M. P., & Graham, J. R. 2007b, *ApJL*, **661**, L85
- Kalas, P., Graham, J. R., Chiang, E., et al. 2008, *Sci*, **322**, 1345
- Kalas, P., Graham, J. R., & Clampin, M. 2005, *Natur*, **435**, 1067
- Kalas, P., Graham, J. R., Clampin, M. C., & Fitzgerald, M. P. 2006, *ApJL*, **637**, L57
- Kalas, P., Graham, J. R., Fitzgerald, M. P., & Clampin, M. 2013, *ApJ*, **775**, 56
- Kalas, P., Liu, M. C., & Matthews, B. C. 2004, *Sci*, **303**, 1990
- Kastner, J. H., Hily-Blant, P., Sacco, G. G., Forveille, T., & Zucherman, B. 2010, *ApJL*, **723**, L248
- Kenyon, S. J., & Bromley, B. C. 2004, *AJ*, **127**, 513
- Kriss, G. A., Wolfe, M. A., Aloisi, A., et al. 2013, in *STScI Instrument Science Report, STIS 2013–03*, (Baltimore: STScI), http://www.stsci.edu/hst/stis/documents/isrs/2013_03.pdf
- Krist, J. E., Ardila, D. R., Golimowski, D. A., et al. 2005, *AJ*, **129**, 1008

- Kuchner, M. J., & Holman, M. J. 2003, *ApJ*, **588**, 1110
- Lagrange, A.-M., Milli, J., Bocatelli, A., et al. 2012, *A&A*, **546**, A38
- Lebreton, J., Augereau, J.-C., Thi, W.-F., et al. 2012, *A&A*, **539**, A17
- Liu, M. C. 2004, *Sci*, **305**, 1442
- Löhne, T., Eiroa, C., Augereau, J.-C., et al. 2012, *AN*, **333**, 441
- López-Santiago, J., Montes, D., Crespo-Chacón, I., & Fernández-Figueroa, M. J. 2006, *ApJ*, **643**, 1160
- Lowrance, P. J., Becklin, E. E., Schneider, G., et al. 2005, *AJ*, **130**, 1845
- MacGregor, M. A., Wilner, D. J., Rosenfeld, K. A., et al. 2013, *ApJL*, **762**, L21
- Mamajek, E. E., Bartlett, J. L., Seifahrt, A., et al. 2013, *AJ*, **146**, 154
- Mamajek, E. E., & Hillenbrand, L. A. 2008, *ApJ*, **687**, 1264
- Maness, H. L., Fitzgerald, M. P., Paladini, R., et al. 2008, *ApJL*, **686**, L25
- Maness, H. L., Kalas, P., Peek, K. M. G., et al. 2009, *ApJ*, **707**, 1098
- Mawet, D., Serabyn, E., Stapelfeldt, K., & Crepp, J. 2009, *ApJL*, **702**, L47
- Moór, A., Ábrahám, P., Derekas, A., Kiss, Cs., et al. 2006, *ApJ*, **644**, 525
- Moór, A., Juhász, A., Kóspál, Á., et al. 2013, *ApJL*, **777**, L25
- Morbidelli, A., Levison, H. F., Tsiganis, K., & Gomes, R. 2005, *Natur*, **435**, 462
- Morbidelli, A., Lunine, J. I., O'Brien, D. P., Raymond, S. N., & Walsh, K. J. 2012, *AREPS*, **40**, 251
- Nesvorný, D., Vokrouhlický, D., & Morbidelli, A. 2013, *ApJ*, **768**, 45
- Nordström, B., Mayor, M., Andersen, J., et al. 2004, *A&A*, **418**, 989
- Panić, O., Holland, W. S., Wyatt, M. C., et al. 2013, *MNRAS*, **435**, 1037
- Patel, R. I., Metchev, S. A., & Heinze, A. 2014, *ApJS*, **212**, 10
- Perrin, M. D., Duchene, G., Millar-Blanchaer, M., et al. 2014a, *ApJ*, in press (arXiv:1407.2495)
- Perrin, M. D., Hines, D. C., Wisniewski, J. P., & Schneider, G. 2014b, in *Polarization of Stars and Planetary Systems*, ed. L. Kolokolova (Cambridge: Cambridge Univ. Press), in press
- Raymond, S., Quinn, T., & Lunine, J. 2004, *Icar*, **168**, 1
- Raymond, S. N., Kokubo, E., Morbidelli, A., et al. 2014, in *Protostars and Planets VI*, ed. H. Beuther, R. Klessen, C. Dullemond, & Th. Henning (Tucson, AZ: Univ. Arizona Press), in press
- Ricarte, A., Moldvai, N., Hughes, A. M., et al. 2013, *ApJ*, **774**, 80
- Roberge, A., Welsh, B., Kamp, I., Grady, C., & Weinberger, A. 2014, *AAS*, **223**, 40905
- Roccatagliata, V., Henning, Th., Wolf, S., et al. 2009, *A&A*, **497**, 409
- Rodigas, T. J., Debes, J. H., Hinz, P. M., et al. 2014a, *ApJ*, **783**, 21
- Rodigas, T. J., Hinz, P. M., Leisenring, J., et al. 2012, *ApJ*, **752**, 57
- Rodigas, T. J., Malhorta, R., & Hinz, P. 2014b, *ApJ*, **780**, 65
- Schneider, G., Silverstone, M. D., & Hines, D. C. 2005, *ApJL*, **629**, 117
- Schneider, G., Silverstone, M. D., Hines, D. C., et al. 2006, *ApJ*, **650**, 414
- Schneider, G., Silverstone, M., Stobie, E., & Rhee, J. 2011, in *HST Calibration Workshop*, ed. S. Deustra & C. Oliveira (STScI: Baltimore), 154
- Schneider, G., Smith, B. A., Becklin, R. R., et al. 1999, *ApJL*, **513**, L172
- Schneider, G., Weinberger, A. J., Becklin, E. E., Debes, J. H., & Smith, B. A. 2009, *AJ*, **137**, 53
- Song, I., Caillault, J.-P., Barrado y Navascues, D., Stauffer, J. R., & Randich, S. 2000, *ApJL*, **532**, L41
- Sommer, R., Pueyo, L., & Larkin, J. 2012, *ApJL*, **755**, L28
- Stark, C. C., Schneider, G., Weinberger, A. J., et al. 2014, *ApJ*, **789**, 58
- Strom, R. G., Malhotra, R., Ito, T., Yoshida, F., & Kring, D. A. 2005, *Sci*, **309**, 1847
- Su, K. Y., Bryden, G., Rieke, G., Stapelfeldt, K. R., & Balog, Z. 2013, *AAS*, **221**, 144.19
- Telesco, C. M., Fisher, R. S., Piña, R. K., et al. 2000, *ApJ*, **530**, 329
- Thalmann, C., Janson, M., Buenzli, E., et al. 2011, *ApJ*, **743**, 6
- Tsiganis, K., Gomes, R., Morbidelli, A., & Levison, H. F. 2005, *Natur*, **435**, 459
- Walsh, J., & Testi, L. 2012, *Msngr*, **148**, 32
- Wahhaj, A., Liu, M. C., Biller, et al. 2104, *A&A*, accepted (arXiv:1404.6525)
- Williams, J. P., Najita, J., Liu, M. C., et al. 2004, *ApJ*, **604**, 414
- Wolf, S., Moro-Martín, A., & D'Angelo, G. 2007, *P&SS*, **55**, 569
- Wyatt, M. C. 2003, *ApJ*, **598**, 1321
- Wyatt, M. C. 2005, *A&A*, **440**, 937
- Wyatt, M. C. 2008, *ARA&A*, **46**, 339
- Wyatt, M. C., Dermott, S. F., Telesco, C. M., et al. 1999, *ApJ*, **527**, 918
- Wyatt, M. C., Smith, R., Greaves, J. S., et al. 2007a, *ApJ*, **658**, 569
- Wyatt, M. C., Smith, R., Su, K. Y. L., et al. 2007b, *ApJ*, **663**, 365
- Zuckerman, B., Song, I., Bessell, M., & Webb, R. 2001, *ApJL*, **562**, L87